

Observation of periodic variable stars towards the Galactic spiral arms by EROS II^{*,**}

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Abstract. We present the results of a massive variability search based on a photometric survey of a six square degree region along the Galactic plane at ($l = 305^\circ$, $b = -0.8^\circ$) and ($l = 330^\circ$, $b = -2.5^\circ$). This survey was performed in the framework of the EROS II (Expérience de Recherche d'Objets Sombres) microlensing program. The variable stars were found among 1 913 576 stars that were monitored between April and June 1998 in two passbands, with an average of 60 measurements. A new period-search technique is proposed which makes use of a statistical variable that characterizes the overall regularity of the flux versus phase diagram. This method is well suited when the photometric data are unevenly distributed in time, as is our case. 1362 objects whose luminosity varies were selected. Among them we identified 9 Cepheids, 19 RR Lyræ, 34 Miras, 176 eclipsing binaries and 266 Semi-Regular stars. Most of them are newly identified objects. The cross-identification with known catalogues has been performed. The mean distance of the RR Lyræ is estimated to be $\sim 4.9 \pm 0.3$ kpc undergoing an average absorption of $\sim 3.4 \pm 0.2$ mag. This distance is in good agreement with that of disc stars that contribute to the microlensing source star population. Our catalogue and light curves are available electronically from the CDS, Strasbourg and from our Web site^{**}.

Key words. techniques: photometry – stars: variables: general – galaxy: stellar content

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* Full Tables 4 and 5 are only available in electronic form at the CDS via anonymous ftp to

cdsarc.u-strasbg.fr (130.79.128.5) or via

<http://cdsweb.u-strasbg.fr/cgi-bin/qcat?J/A+A/389/149>

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1. Introduction

In 1996, the EROS II collaboration started an observation program towards the Galactic Spiral Arms (GSA) dedicated to microlensing events. Since then, four regions of the Galactic plane located at large angles with respect to the Galactic Centre are being monitored to disentangle the disc, bar and halo contributions to the microlensing optical depth. Seven microlensing event candidates have already been published, based on three years (1996–98) of observations (Derue et al. 1999, 2001), (hereafter Papers I and II). The distance of the source stars used in these papers to compute the expected optical depths was deduced from a detailed study of our colour-magnitude diagrams. It was thus found that the source star population is located ~ 7 kpc away, undergoing an interstellar extinction $A(V)$ of about 3 mag (see Mansoux 1997 for more details). This distance estimate is in rough agreement with the distance to the spiral arms obtained by Georgelin et al. (1994) and Russeil et al. (1998), but its uncertainty is limiting further interpretation of our microlensing optical depth estimates. It was therefore desirable to seek more information on the distance distribution of the source stars – whether these stars belong to the disc or to the spiral arms – and on the reddening along our observation line of sights. This led us to perform a dedicated variable star search between April and June 1998, on a subset of our Galactic plane fields. The analysis was restricted to the brightest stars of this subset.

Among the wide variety of variable stars, periodic ones are of particular interest. The properties of Cepheids make them well suited to trace the Galactic spiral arms. Their reddening is measurable as well as their distance via the period-luminosity (PL) relation. RR Lyræ stars are old stars, well suited to trace the disc population. One can infer their mean dereddened magnitude and their absolute magnitude (Gould & Popowski 1998). The infrared PL(K) relation for Miras and Semi-Regular variable stars can be calibrated using a comparison of DENIS and EROS LMC giant stars (Cioni et al. 2001). Finally detached eclipsing binaries also offer the opportunity to measure their stellar parameters and their distance (Paczynski 1996).

This paper presents the results of this particular campaign that led to a catalogue containing a large number of new variable objects in the Galactic plane. Section 2 gives the basic features of the observational setup, Sect. 3 gives details on a new algorithm used to search for periodic variations of the luminosity. Section 4 describes the catalogue and the cross-identification process. In Sect. 5 we use the selected RR Lyræ to estimate the mean reddening of our fields and we give the distance distribution of these stars.

2. Experimental setup and observations

The MARLY telescope and its two cameras, the way we carry out our observations, as well as our data reduction sequence are described in Paper I and references therein.

Table 1. Characteristics of the 6 one square degree fields monitored for this study. The equatorial (J2000) and galactic coordinates of the field centres, the number of photometric measurements per light curve per colour N_m and the number of analysed light curves N_{obs} (in millions) are given for each field. The observation duration was ~ 100 days.

Field	α	δ	b°	l°	N_m	N_{obs}
Norma (γ Nor)					59	1.30
gn450	16:09:45	-53 : 07 : 03	-1.17	330.49	63	0.28
gn453	16:22:28	-52 : 06 : 20	-1.69	332.24	52	0.36
gn455	16:26:52	-52 : 21 : 02	-2.35	332.54	57	0.30
gn459	16:15:51	-54 : 48 : 45	-2.86	329.82	61	0.36
Musca (θ Mus)					60	0.61
tm550	13:27:04	-63 : 02 : 18	-0.47	306.98	61	0.27
tm551	13:31:18	-63 : 34 : 41	-1.07	307.37	60	0.34
Total					191	

The two EROS passbands are non standard. The so-called EROS-red passband R_E is centred on $\bar{\lambda} = 762$ nm, close to I_C Cousins, with a full width at half maximum $\Delta\lambda \simeq 85$ nm, and EROS-visible passband V_E is centred on $\bar{\lambda} = 600$ nm, close to V_J Johnson, with $\Delta\lambda \simeq 78$ nm. The EROS II colour magnitude system is defined as follows: a zero colour star with $V_E - R_E = 0$ (a main sequence A0 star) will have its R_E magnitude numerically equal to its Cousins R_C magnitude and its V_E magnitude numerically equal to its Johnson V_J magnitude. The colour transformation between the EROS II system (V_E, R_E) and the standard Johnson-Cousins (V_J, I_C) system is then:

$$I_C = R_E + 0.01 \times (V_E - R_E) \quad (1)$$

$$V_J = V_E + 0.39 \times (V_E - R_E).$$

The colour coefficients are obtained from the study of our passbands based on Landolt standards and on one of the EROS II LMC fields observed simultaneously in $BV_J R_C I_C$ with the Danish 1.54 m (at ESO-La Silla) and with the MARLY. The zero points are established with tertiary standards in $BVRI$ taken with the Danish 1.54 m (Regnault 2000). We have cross-checked our R_E photometry with the I_C one of DENIS (Fouqué et al. 2000). Furthermore, using Eqs. (1), the mean magnitudes of the LMC red-giant clump stars agree within 0.1^{mag} with determinations made by Harris & Zaritsky (1999) and Udalski et al. (1998). We thus estimate that the precision of the zero points of the MARLY calibration is $\sim 0.1^{\text{mag}}$.

Among the 29 fields of the EROS GSA microlensing program, the six fields considered here were monitored about once per night between April and June 1998. They represent 2 square degrees towards θ Mus and 4 towards γ Nor (named after the closest bright star). Table 1 gives their coordinates, the number of images N_m taken in each direction and the number of analysed light curves N_{obs} . To avoid CCD saturation by the brightest stars ($I_C \sim 9$), by Cepheids or Miras in particular, we have reduced the exposure time to 15 s instead of the 120 s used in the microlensing survey. As a consequence the catalogue is incomplete as far as faint stars are concerned, but it could be

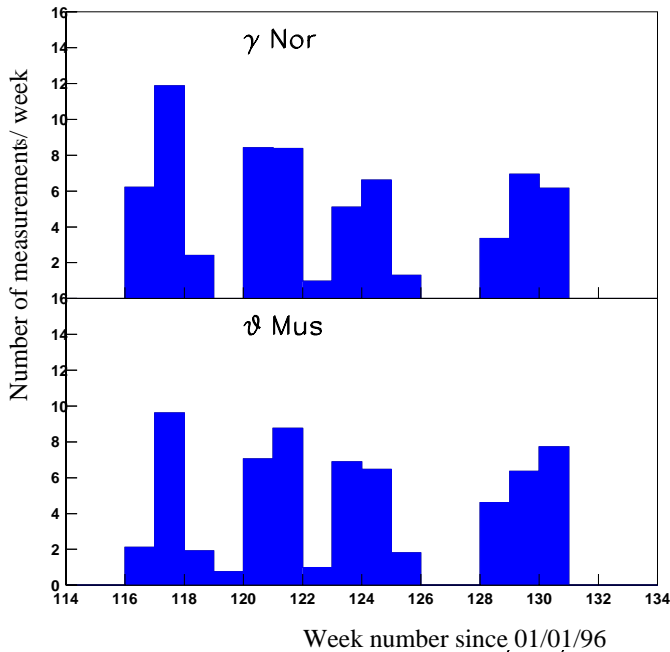


Fig. 1. Average time sampling for the 6 fields monitored (2 towards γ Nor, upper panel, and 4 towards θ Mus, lower panel), in number of measurements per week and per field.

updated later by using the total set of available GSA images. Figure 1 shows the average time sampling. Three gaps can be seen in our data: the first two (around weeks 119 and 123) were due to bad weather conditions while the third one (around week 127) corresponds to the annual maintenance of our setup.

3. Search for periodic stars

3.1. Reconstruction of the light curves

Since the EROS photometry is described in detail in Ansari (1996) only the main features of the PEIDA++ package are summarised below. For each field, a template image is first constructed using one exposure of very good quality. A reference star catalogue is set up with this template using the CORRFIND star finding algorithm (Palanque-Delabrouille et al. 1998). For each subsequent image, after geometrical alignment with the template, each identified star is fitted together with its neighbours, using a PSF determined on bright isolated stars and imposing the position from the reference catalogue. A relative photometric alignment is then performed, assuming that most stars do not vary. Photometric errors are computed for each measurement, assuming again that most stars are stable, and parameterised as a function of star brightness and image sequence number. Figure 2 shows the mean point-to-point relative dispersion of the measured fluxes along the light curves as a function of R_E and V_E . The photometric accuracy is $\sim 15\%$ at $R_E \sim 18$, and about 2% for the brightest stars. Finally, using the PEIDA++ photometric package, we reconstruct the light curves of 1913576 stars.

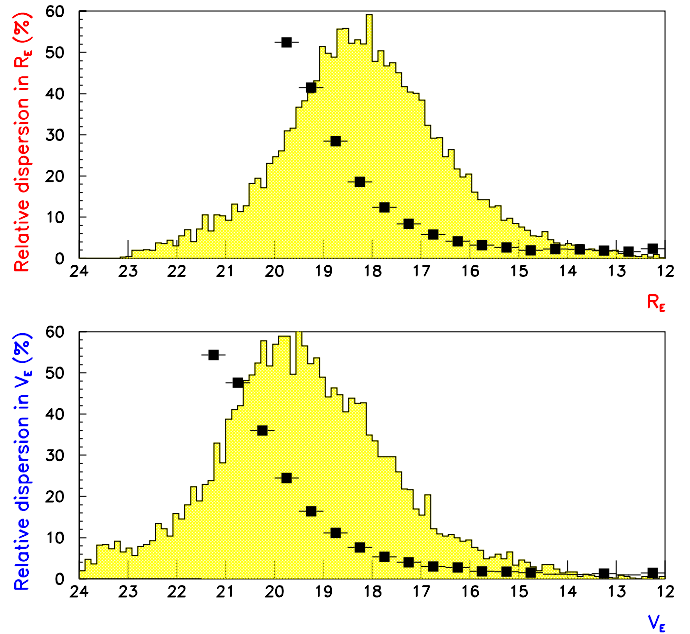


Fig. 2. Average value of the relative frame to frame dispersion of the luminosity measurements versus R_E (upper panel) and V_E (lower panel), for stars with at least 30 reliable measurements in each colour. This dispersion is taken as an estimator of the mean photometric precision. The superimposed hatched histograms show the EROS magnitude distribution of monitored stars.

Table 2. Impact of each selection criterion on the data. For each cut, the number of remaining light curves is given.

Cut	Criterion	
	Total analysed	1913576
1	$N_m > 30$	1299690
2	$R_E < 17$	330089
3	Pre-filtering	41545
4	Period search	2553
5	Aliasing	2424
6	Visual inspection	1362

3.2. Pre-selection

Each one of the light curves is subjected to a series of selection criteria in order to isolate a small sub-sample on which we will apply the time consuming period search algorithms. These analysis cuts are briefly described hereafter (see Derue 1999 for more details) and their effect on the data is summarised in Table 2:

- cut 1:* At least 30 measurements should be available in both passbands and the base flux must be positive;
- cut 2:* The search is restricted to stars whose magnitude is $R_E < 17$ which corresponds to a photometric accuracy in R_E better than $\sim 10\%$.
- cut 3:* A non specific pre-filter is applied which retains most variable stars. It selects light curves satisfying one or both of the following criteria:
 - the relative dispersion of the flux measurements is 25% larger than the average one for the set of stars having the same magnitude;

- the distribution of the deviations with respect to the base flux is incompatible with the one expected from a stable source with Gaussian errors during the observation period (Kolmogorov-Smirnov test).

These cuts are tuned to select $\sim 10\%$ of the light curves. We have checked that this procedure allows one to retrieve the previously known Cepheids observed by EROS in the Magellanic Clouds. We also keep a randomly selected set of light curves ($\sim 2\%$) to produce unbiased colour-magnitude diagrams, for comparison purposes;

At this stage a set of 41 545 light curves remains which is then subjected to a periodicity search.

3.3. Light curve selection

We use three independent methods to extract periodic light curves. The first two are classical methods already described in the literature: method 1 is based on the Lomb-Scargle periodogram (see Scargle 1982) while method 2 makes use of the One Way Analysis of Variance algorithm (see Schwarzenberg-Czerny 1996). Both provide the probability for false periodicity detection. In method 1 one computes the Fourier power over a set of frequencies. It is therefore well adapted to identify sinusoidal light curves. It can be improved by incorporating higher harmonics in order to detect any kind of variability such as eclipsing binaries (Grison 1994; Grison et al. 1995).

We developed a new method, the third one, in order to extract periodic light curves in a way which is insensitive to the particular shape of the variation. This method also provides a probability for false periodicity detection.

It consists in searching for a frequency such that the corresponding phase diagram, i.e. the series of fluxes F_i versus phases φ_i in increasing order of φ , displays a regular structure significantly less scattered than for other frequencies. Let T_{obs} be the observation duration (~ 100 days in this analysis). We span the frequency domain from T_{obs}^{-1} to $(0.2 \text{ days})^{-1}$ with a constant step of $(4 \times T_{\text{obs}})^{-1}$. The total number of test-frequencies is thus $N_{\text{test}} \sim 2000$. This sampling ensures that the total phase increment over T_{obs} is $\pi/2$ for two adjacent test-frequencies.

For each value of the test-frequency we compute the corresponding phase diagram. We calculate a χ^2 from the weighted differences of F_i and the fluxes interpolated between F_{i-1} and F_{i+1} :

$$\chi^2 = \sum_{i=1}^{N_m} \left(\frac{F_i - (1 - R_i) \times F_{i-1} - R_i \times F_{i+1}}{\sigma_i} \right)^2 \quad (2)$$

where $R_i = (\varphi_i - \varphi_{i-1}) / (\varphi_{i+1} - \varphi_{i-1})$ and N_m is the number of measurements. The uncertainty σ_i takes into account the errors ϵ_i on the flux F_i and on the interpolated flux: $\sigma_i^2 = \epsilon_i^2 + (1 - R_i)^2 \epsilon_{i-1}^2 + R_i^2 \epsilon_{i+1}^2$. Expression (2) can be interpreted as the χ^2 of the set of differences between the odd measurements with respect to the line joining even

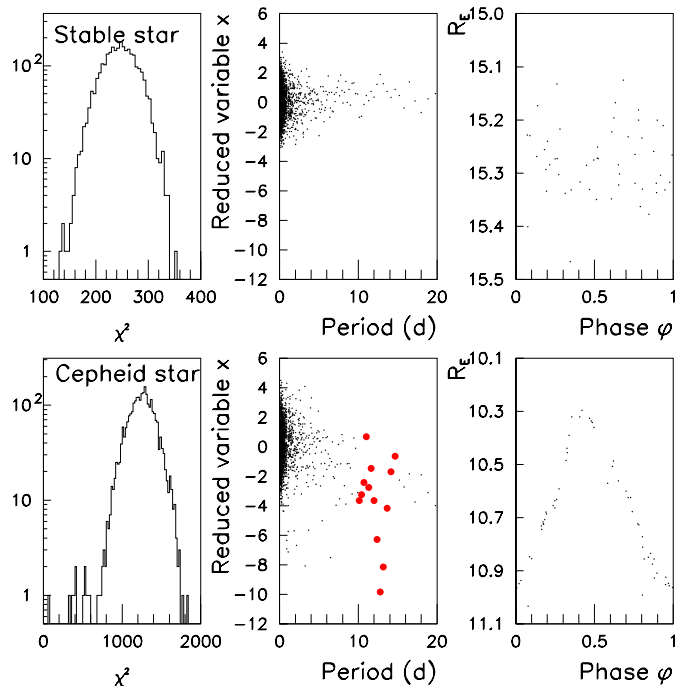


Fig. 3. Distributions of χ^2 (left panels) and variation of x with the value of the test-period in days (middle panels) obtained with the third method for a stable star and for a typical Cepheid candidate. The bold dots are pointing to the actual period of the star. The right panels show the light curve (in R_E) obtained once the period has been folded in.

ones, added to the χ^2 of the set of differences between even measurements with respect to the line joining odd ones. If the star is measured in both colours, we add the χ^2 obtained in each colour (N_m is then twice as large). For a given stable star with Gaussian errors, each phase diagram can be considered as a random realisation of the light curve. When the test-frequency spans the search domain, the distribution of the χ^2 parameter defined by Eq. (2) is the one of the standard χ^2 with N_m degrees of freedom. Since N_m is large enough, this distribution is close to a Gaussian with average N_m and variance $2N_m$ (see upper left panel of Fig. 3). For a periodic variable star, the χ^2 distribution displays a main cluster, when the test-frequency results in a phase diagram with non-correlated point to point variations, and a few lower values when the test-frequency corresponds to a phase diagram with a regular structure (see lower left and middle panels of Fig. 3). In practice, instead of using the parameter defined by Eq. (2), we use the reduced variable:

$$x = \frac{\chi^2 - \langle \chi^2 \rangle}{\langle \chi^2 \rangle / N_m} \times \frac{1}{\sqrt{2N_m}} \quad (3)$$

where $\langle \chi^2 \rangle$ is the average of the realisations of χ^2 for all test-frequencies. For a stable star, the distribution of this variable x is a Gaussian centred at zero, with unit variance. If the errors are correctly determined $\langle \chi^2 \rangle / N_m$ is close to unity; if the errors are all systematically overestimated (or underestimated), then including this term ensures a global renormalisation of the

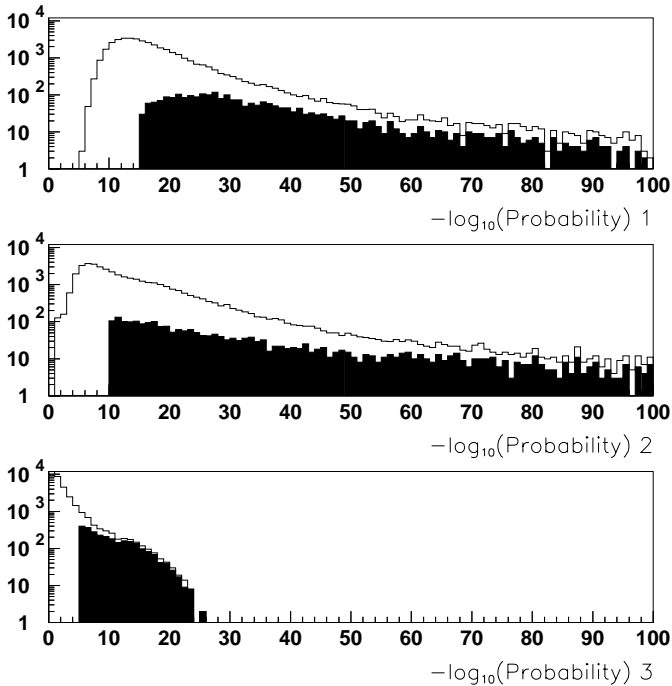


Fig. 4. Distribution of $-\log(\text{Prob})$ for method 1 (top), 2 (middle) and 3 (bottom). The histograms show the distribution for all 41 545 stars on which the periodicity search is done (in white) and for the 2553 selected light curves (in black).

errors in Eq. (3), and the distribution of our reduced variable x is also a normal distribution. Let x_{\min} be the smallest value of x calculated among all test-frequencies for a given star. Under the hypothesis that the light curve is produced by a stable star, the probability to obtain at least one value $x \leq x_{\min}$ in a series of N_{test} realisations is $\text{Prob}(x < x_{\min} | N_{\text{test}}) = 1 - \left[1 - \frac{1}{\sqrt{2\pi}} \int_{-\infty}^{x_{\min}} e^{-x^2/2} dx \right]^{N_{\text{test}}}$. If this probability is small, then:

$$\text{Prob}(x < x_{\min} | N_{\text{test}}) \simeq \frac{N_{\text{test}}}{2} \text{erfc} \left(\frac{-x_{\min}}{\sqrt{2}} \right). \quad (4)$$

If the light curve exhibits periodic variations, then there exist test-frequencies for which x is significantly smaller than typical values of this variable (see middle and right panels of Fig. 3), and the probability for false detection is then extremely small. Figure 4 displays the probability distribution obtained with the three methods for the set of filtered light curves¹.

We apply the three algorithms which all give a probability for no periodicity. A star is accepted only if selected by all three methods with the following thresholds: $P(\text{method 1}) < 10^{-15}$, $P(\text{method 2}) < 10^{-10}$, $P(\text{method 3}) < 10^{-5}$, tuned in order to allow one to retrieve the previously known Cepheids observed by EROS in the Magellanic Clouds (see Derue 1999 for more details). This procedure (*cut 4*) selects a sample of 2553 stars.

¹ The third method gives a probability which is always larger than 10^{-30} . Indeed, the smallest value of x which can be obtained takes place when $\chi^2 = 0$, then $|x| \sim \sqrt{N_m} < 10$.

Ten times more stars would have been selected if we had used method 1 or method 2 only (with the same thresholds), most of them being spurious variables. The third method would have added far less candidates if used alone, but still a factor 2 more. Combining independent methods has thus the advantage of considerably reducing the noise background (mostly due to aliases of one day or noisy measurements) while giving redundant information about the period. To obtain individual periods we perform Fourier fits with five harmonics:

$$\text{flux} = \text{flux}_0 + \sum_{l=1}^5 a_l \cos\left(\frac{2\pi}{P} l(t - t_0) + \phi_l\right), \quad (5)$$

where P is the period, ϕ_l the phase and a_l the amplitude. We define the amplitude ratios $R_{kl} = a_k/a_l$, and the phase differences (defined modulo 2π) $\Phi_{kl} = \phi_k - k\phi_l$, with $k > l$. Objects with non-significant harmonic amplitudes (i.e. with almost sinusoidal light curves) have $R_{21} \sim 0$ and their Φ_{21} is ill defined.

The selection of periodic variable stars is complicated by aliases. Some of the stars with periods equal to a simple fraction or a low multiple of one day may be badly phased because of the nightly sequence of measurements. These aliased periods are seen in Fig. 5 as vertical groups of dots at $2/3$, 2 and 3 days. To eliminate them we demand (*cut 5*) that the fitted periods are not within $\pm 1\%$ of these values. One can also notice some vertical groups of points around 25 days which correspond to data gaps in our sample (see Fig. 1). Once these objects are removed, 2424 stars remain. The flux values of the remaining stars are folded using each period obtained with the three methods. The resulting phase diagrams are visually inspected. Some of them display an obvious spurious periodic or quasi-periodic variability due to a low photometric quality. After this final visual selection (*cut 6*) the list of variable stars includes 1362 candidates which exhibit unambiguous periodic variability.

3.4. Type of variability

The classification of the selected stars among different types of variability cannot be based on the position of the objects in the colour-magnitude diagram since the spread in distance of these stars entails a spread in magnitude and colour. It is desirable however to classify the various light curves according to some physical parameters. In the following we mainly use criteria based on the period P of the luminosity variations and on the amplitude ratio $\Delta V_E/\Delta R_E$. For each selected type of variable star the phase diagram of a typical candidate is displayed in Fig. 6. Three groups are distinguished depending on their period:

- 1st: stars with a period larger than 60 days (543 objects): For these objects an entire period has not been observed. There is thus no warranty that these objects are periodic ones. 34 display a nearly linear light curve and are thus catalogued as Miras candidates.

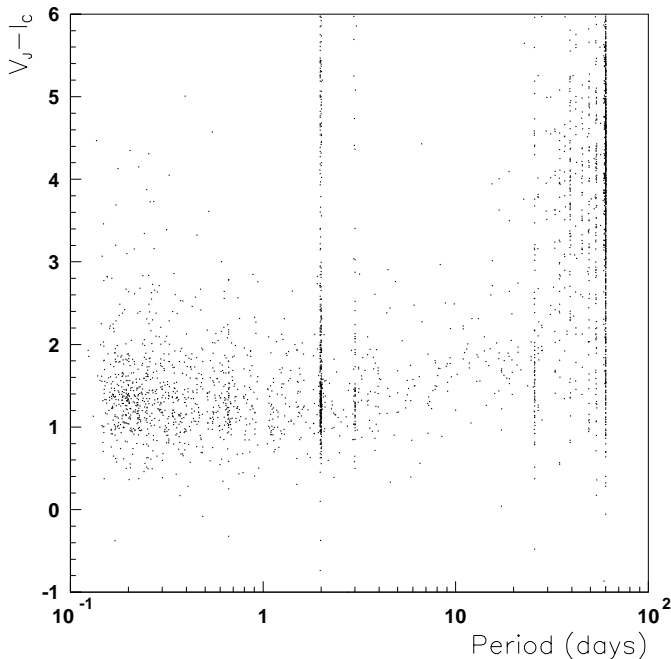


Fig. 5. Period-colour diagram (P in days vs. $V_J - I_C$) of the 2553 selected candidates before cut 5.

The other 509 objects are catalogued as Long Period Variable stars (LPVs).

2nd: stars with a period between 30 and 60 days (387 objects):

264 Semi-Regular variable stars are selected by requiring: $\Delta V_E / \Delta R_E > 1.2$ to select pulsating stars (see below), $V_J - I_C > 2.5$ to discriminate from bluer variable stars and $\Delta R_E < 1.0$ to avoid possible Miras or LPVs wrongly phased. The long term stability of these stars is not known. Some of the reported periods may change from season to season, as a result of their semi-regular behaviour. The remaining 123 objects are catalogued as miscellaneous variable stars.

3rd: stars with a period smaller than 30 days (432 objects):

The colour change for a Cepheid in standard passbands is $\Delta V_J / \Delta I_C > 1.3$ (Madore et al. 1991) which corresponds to $\Delta V_E / \Delta R_E > 1.2$ in the EROS system. Two sets are thus distinguished based on this criterion:

– The pulsating variable stars (60 objects):

For stars with period $P < 1$ day, two samples of RR Lyræ are identified: the RRc have $R_{21} < 0.4$ (14 objects) and the RRab have $R_{21} > 0.4$ (5 objects). We adopt the morphological classification proposed by Antonello et al. (1986) and classify as s -Cepheids the stars that lie in the lower part of the $R_{21} - P$ plane, and as classical Cepheids the remaining stars. s -Cepheids pulsate in the first overtone and classical Cepheids in the fundamental mode (see e.g. Beaulieu et al. 1995; Beaulieu & Sasselov 1996). We use the

Table 3. Number of selected objects for each type of variability.

Period range	Type	Number of objects
$P > 60$ d		543
	LPV	509
	Miras	34
$60 > P > 30$ d		387
	Semi-Regular	264
	miscellaneous	123
$P < 30$ d		432
	pulsating	60
	RRc	14
	RRab	5
	classical-Cepheids	6
	s -Cepheids	3
	miscellaneous	32
	non-pulsating	372
	EA	130
	EB	35
EW	11	
	miscellaneous	196

empirical function $R_{21}^{\text{cut}}(P) = 0.4 - (P/30 \text{ days})$ in order to separate these pulsation modes (see Fig. 7). Among the five objects that pass the s -Cepheid cut, only three belong to the $R_{21} - P$ and $\Phi_{21} - P$ distributions of galactic s -Cepheids and their phase parameter Φ_{21} is poorly constrained. Besides most of classical Cepheids have amplitudes larger than $\Delta R_E > 0.4^{\text{mag}}$ (see e.g. Afonso et al. (1999)). For stars with period $P > 1$ day, three samples are then identified: the classical Cepheids have $R_{21} > R_{21}^{\text{cut}}$ and $\Delta R_E > 0.4^{\text{mag}}$ (6 objects); the s -Cepheids have $R_{21} < R_{21}^{\text{cut}}$ and $3 \text{ rad} < \Phi_{21} < 6 \text{ rad}$ (3 objects); the remaining 32 objects are catalogued as miscellaneous pulsating stars.

– The non-pulsating variable stars (372 objects):
The remaining objects have similar amplitudes in both passbands. We classify them according to the following criteria. Algol systems (130 stars, type EA) display well-defined eclipses whose secondary one has a depth lower than half the primary one, and possibly flat light curve between them. EB type objects (35 in total) show a secondary eclipse equal to half the primary one. The EW type (11 objects) is characterised by similar depths of the two eclipses. The members of a residual sample of 196 objects do not look like convincing eclipsing binaries and are catalogued as miscellaneous variable stars.

As emphasised by Udalski et al. (1999a) a large number of variable objects show small amplitude sinusoidal variations, such as ellipsoidal binary variable stars. A contamination of the sample of pulsating stars by eclipsing binaries is thus possible.

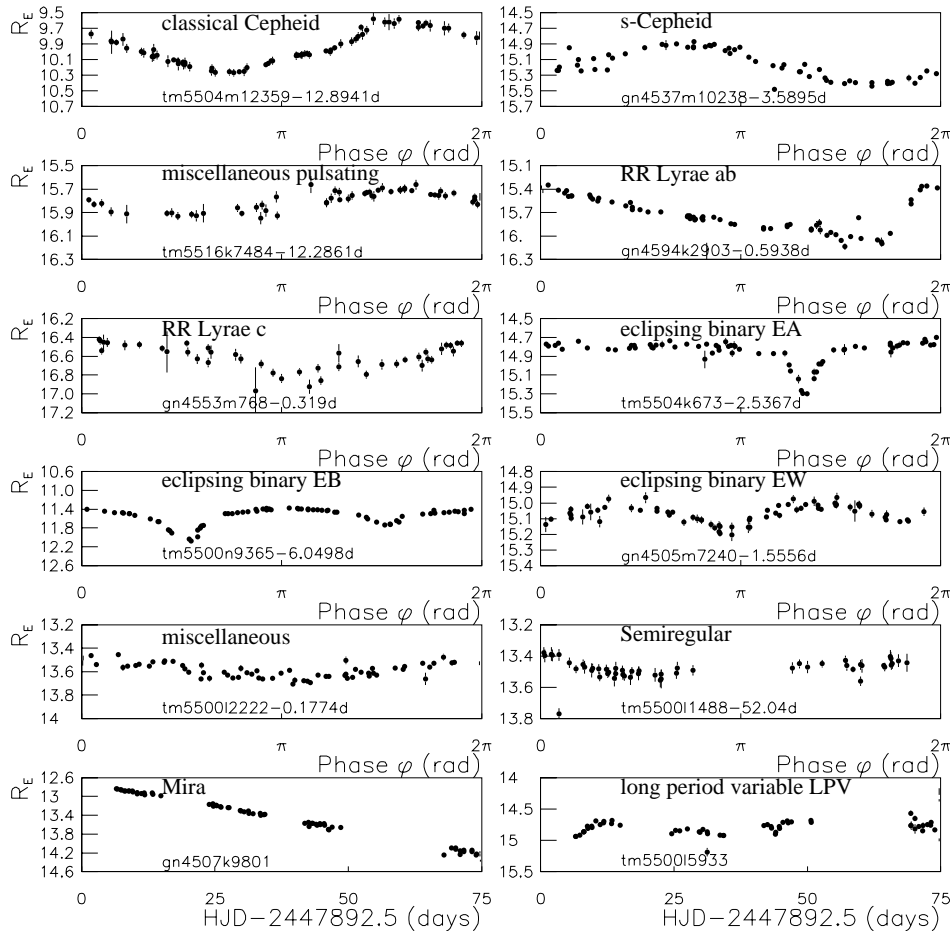


Fig. 6. Phase diagrams (or light curves for the two lowest panels) for typical variable stars of our catalogue in R_E magnitude.

Figures 8 and 9 show the location of the selected variables in the colour-magnitude diagrams. Also plotted are 10000 stars located in the central part of the two fields tm550 and gn450. Most of the Cepheids are much brighter than our magnitude threshold (cut 2); this is not so for RR Lyrae (see Fig. 9, lower panel). Our catalogue is thus not complete for this type of variable stars, as already mentioned.

4. The catalogue

The catalogue is composed of two tables containing objects with periods P smaller or larger than 30 days, respectively. The identifier of each star is given according to the recommendations of the IAU Commission 5 in *The Rules and Regulations for Nomenclature* (see the Annual Index of A&A). The general acronym used in the catalogue is EROS2 GSA followed by J2000 equatorial coordinates in the format JHHMMSS \pm DDMSS. The remainder of the identifier in parentheses gives some information relating to the internal organisation of the EROS database: gnnnn or tmnnn is the name of the field, followed by the CCD number and the location on the image following the EROS II nomenclature. The remaining number is the star identifier used in the EROS database. As an example, J132630-630945(tm5504m12359) is the name of

the 12359th star observed in quarter m of CCD 4 in the field tm550. The J2000 equatorial coordinates of this star are 13:26:30.11, -63:09:45.66.

The equatorial coordinates (J2000) of individual stars have been obtained as follows. First, we have inserted the suitable WCS keywords into the header of the EROS II reference images using the WCSTOOLS package (Mink 1999). Whenever possible, the cross-identification of each star with previously known objects within a $10''$ search radius has been done using the SIMBAD and VIZIER databases available at the CDS, Strasbourg.

The tables contain the following information:

1. Identifier;
2. Right ascension α (J2000);
3. Declination δ (J2000);
4. $\langle R_E \rangle$ mean magnitude in EROS-red;
5. ΔR_E amplitude peak to peak in R_E ;
6. $\langle V_E \rangle$ mean magnitude in EROS-visible;
7. ΔV_E amplitude peak to peak in V_E ;
8. Period in days. Note that periods longer than 30 days are given with less accuracy since the time span of the measurements does not allow a precise determination. Measured periods which are longer than 60 days (i.e. $2/3$ of the observation period) are flagged by writing “ $P > 60$ d” and have no warranty to be true periodic

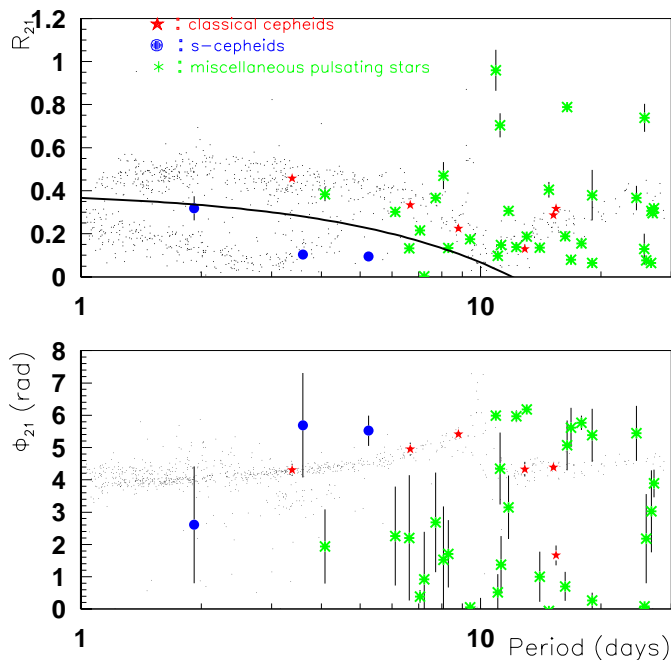


Fig. 7. Distinction between classical Cepheids, *s*-Cepheids and miscellaneous pulsating stars in the $R_{21}-P$ and $\Phi_{21}-P$ planes. The individual uncertainties are reported. The cloud of points represent the Cepheids observed in the LMC (adapted from Afonso et al. 1999). The curve corresponds to the empirical function $R_{21}^{\text{cut}}(P) = 0.4 - (P/30 \text{ days})$ used to distinguish between *s* and classical Cepheids.

variable stars. The peak to peak amplitude of these stars could be meaningless and the mean magnitude is determined with low accuracy.

For Cepheids and RR Lyræ the results of the Fourier fit are given:

9. Fourier coefficient ratio R_{21} ;
10. Fourier coefficient ratio R_{31} ;
11. Phase difference ϕ_{21} (in rad);

Also given when possible:

13. Type of variability (C=classical Cepheids, S=*s*-Cepheids, puls.=miscellaneous pulsating stars, EA, EB, EW = eclipsing binaries, misc = miscellaneous variable stars, SR = Semi-Regular variables, M = Miras, LPV = long period variables);
14. Name of cross-identified object(s) within a search radius of $10''$.

Tables 4 and 5 show a sample of the catalogue that we plan to install at the CDS (see also our Web site <http://eros.in2p3.fr/>).

A by-product of such a survey is the possibility to update the coordinates given in older catalogues. As an example, it was found that the well known classical Cepheid OO Cen with the SIMBAD identifier V*OO Cen was $20''$ away from an EROS object, namely J132630-630945(tm5504m12359), as shown in the finding chart of Fig. 10. The period and coordinates we report are compatible with the ones given by the GENERAL CATALOGUE

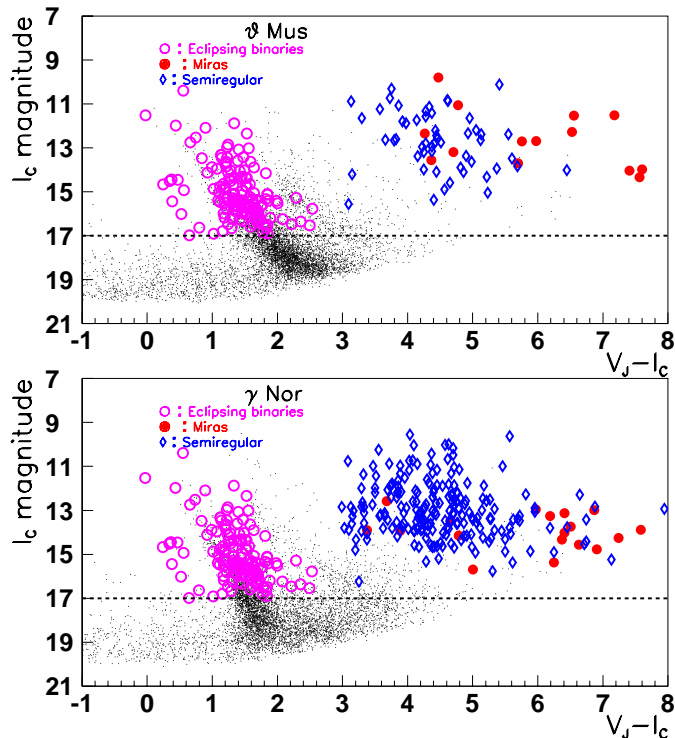


Fig. 8. Colour-magnitude diagrams (I_C vs. V_J-I_C) for the Miras (represented by dots ●), Semi-Regular variable stars (diamonds ◇) and eclipsing binaries (open circles ○). The dotted line corresponds to cut 2 on the luminosity of the stars.

OF VARIABLE STARS (GCVS)² (Kholopov 1985) (see Table 6). It seems that an error occurred when the coordinates of this particular star were filled in the SIMBAD. The light curve of this object is shown in Fig. 6.

We have performed several cross-identifications of our catalogue with those previously available, namely the IRAS POINT SOURCE CATALOG (Beichman et al. 1998), the MSX5C INFRARED ASTROMETRIC CATALOG (Egan et al. 1996), the TWO MICRON ALL SKY SURVEY (2MASS) (Skrutskie et al. 1997), the CKS91 catalogue (Caldwell et al. 1991) and the GENERAL CATALOGUE OF VARIABLE STARS (GCVS) (Kholopov 1985). A total of 38 IRAS sources and 220 MSX5C objects have been thus retrieved. The overlap with the available 2MASS catalogue exists only for the EROS field gn459, representing 255 stars. A total of 233 2MASS objects have been thus retrieved among them 37 objects classified as Semi-Regular variables.

An overlap exists with the CKS91 catalogue (Caldwell et al. 1991). These authors have searched for bright Cepheids and other variable stars with $I < 14$, towards Crux and Centaurus, during 42 days with less than 10 measurements per star. A small overlap exists between this survey and the GSA fields tm550 and tm551. Unfortunately this overlap involves our CCD #2 which was not operational at the time of the

² Queries on this catalogue are possible at <http://www.sai.msu.su/cgi-bin/wdb-p95/gcvs/stars/form>

Table 4. Catalogue of stars with $P < 30$ days (sample).

EROS2 GSA ID	RA(J2000)	DEC(J2000)	$\langle R_E \rangle$	ΔR_E	$\langle V_E \rangle$	ΔV_E	$P(\text{days})$	R_{21}	$\sigma_{R_{21}}$	R_{31}	$\sigma_{R_{31}}$	ϕ_{21}	$\sigma_{\phi_{21}}$	Type	Note
J132357-631158(tm5504k673)	13:23:57.99	-63 : 11 : 58.88	14.73	0.58	15.7	0.64	2.5367							EA	
J132402-633439(tm5506k1244)	13:24:02.90	-63 : 34 : 39.66	14.20	0.57	16.03	0.70	8.8020	0.224	0.004	0.215	0.002	5.41	0.06	C	
J132433-624202(tm5500l6773)	13:24:33.35	-62 : 42 : 02.40	12.69	0.65	13.37	0.50	1.7601							EA	V* NX Cen
J132449-631304(tm5504k9109)	13:24:49.24	-63 : 13 : 04.60	14.56	1.24	15.72	1.41	3.5099							EA	
J132457-633032(tm5506k8804)	13:24:57.51	-63 : 30 : 32.00	12.55	0.54	13.66	0.52	1.5775							EA	
J132500-631337(tm5504k10856)	13:25:00.23	-63 : 13 : 37.21	14.64	0.22	15.91	0.31	14.8285	0.405	0.036	0.151	0.008	-0.06	0.03	puls	
J132531-633526(tm5506m1888)	13:25:31.96	-63 : 35 : 26.48	14.96	1.36	16.15	1.47	1.6274							EA	
J132554-631029(tm5504m6315)	13:25:54.62	-63 : 10 : 29.34	13.71	0.25	14.71	0.20	0.4349							misc	
J132630-630945(tm5504m12359)	13:26:30.11	-63 : 09 : 45.66	10.18	0.52	11.77	0.64	12.8941	0.129	0.001	0.146	0.001	4.33	0.23	C	V* OO Cen
J132721-630110(tm5503l3554)	13:27:21.68	-63 : 01 : 10.80	10.98	0.80	12.67	0.97	15.2278	0.286	0.006	0.219	0.003	4.39	0.15	C	V* V881 Cen

Table 5. Catalogue of stars with $P > 30$ days (sample).

EROS2 GSA ID	RA(J2000)	DEC(J2000)	$\langle R_E \rangle$	ΔR_E	$\langle V_E \rangle$	ΔV_E	$P(\text{days})$	Type	Note
J132359-623229(tm5500l1488)	13:23:59.86	-62 : 32 : 29.63	10.83	0.14	14.18	0.22	52.04	SR	MSX5C G306.6758+00.0933
J132400-632057(tm5504l1003)	13:24:00.60	-63 : 20 : 57.82	12.34	0.20	16.06	0.38	45.83	SR	[CKS91] 13206-6305
J132419-632717(tm5506k3687)	13:24:19.25	-63 : 27 : 17.42	11.56	0.34	14.94	5.43	>60 d	LPV	MSX5C G306.5942-00.8179
J132427-631542(tm5504k5591)	13:24:27.10	-63 : 15 : 42.08	11.75	0.18	14.88	0.33	>60 d	LPV	[CKS91] 13211-6300
J132431-630539(tm5504k6408)	13:24:31.78	-63 : 05 : 39.51	15.34	0.27	17.19	0.50	55.70		
J132433-632201(tm5504m10330)	13:24:33.05	-63 : 22 : 01.64	11.96	0.24	9.49	0.43	58.46		[CKS91] 13229-6252
J132437-623119(tm5500k7345)	13:24:37.80	-62 : 31 : 19.09	12.4	0.21	15.94	0.47	>60 d		
J132440-633236(tm5506k6457)	13:24:40.22	-63 : 32 : 36.78	11.67	0.34	15.46	0.42	>60 d		MSX5C G306.6222-00.9099
J132444-632224(tm5504l7786)	13:24:44.82	-63 : 22 : 24.20	10.79	0.44	14.13	1.19	55.00		IRAS 13214-6306; MSX5C G306.6524-00.7426
J161332-541002(gn4590k724)	16:13:32.50	-54 : 10 : 02.22	13.14	0.84	16.49	1.79	>60 d	LPV	2MASS 1613326-541004
J161332-541453(gn4590k673)	16:13:32.45	-54 : 14 : 53.91	13.87	0.18	16.12	0.26	38.29	SR	2MASS 1613325-541454

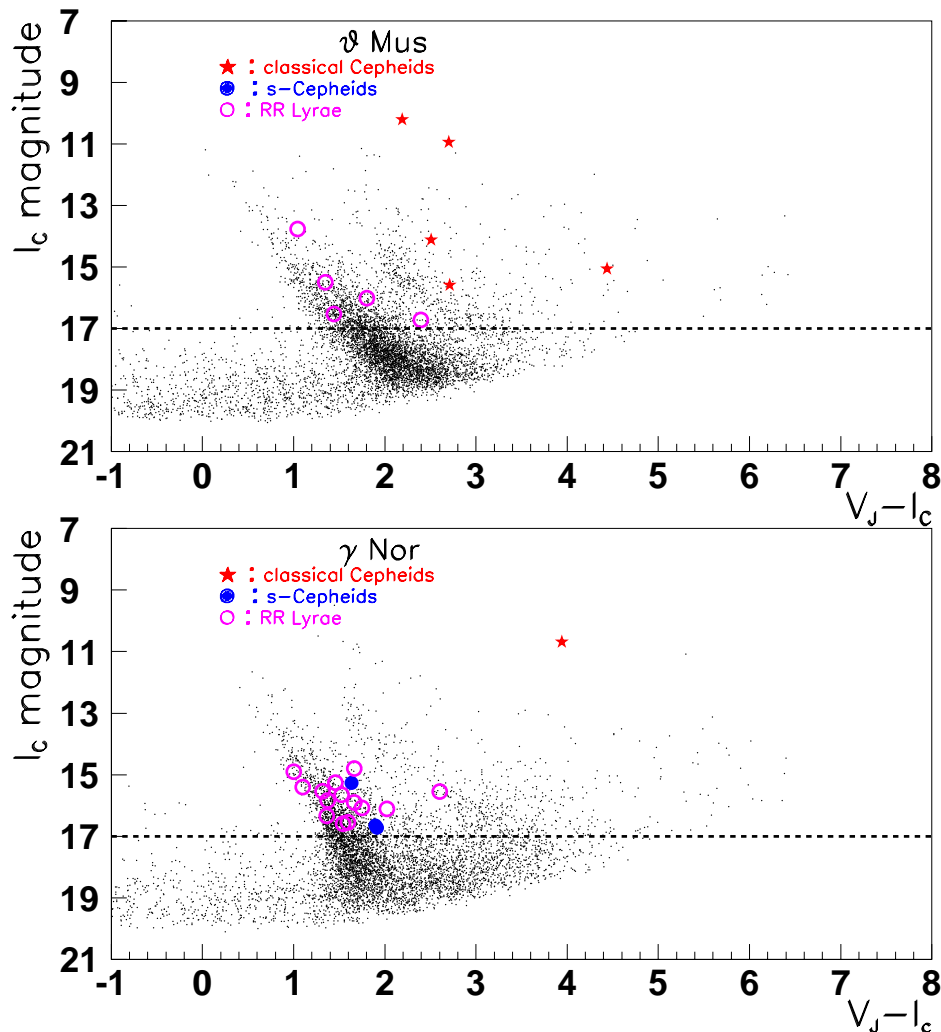


Fig. 9. Colour-magnitude diagrams (I_C vs $V_J - I_C$) for the classical Cepheids (represented by stars \star), s -Cepheids (dots \bullet), and RR Lyrae (open circles \circ).

Table 6. Period and coordinates obtained on OO Cen by SIMBAD, GCVS and this study.

Source	SIMBAD	GCVS	This study
object	V* OO Cen	OO Cen	J132630-630945(tm5504m12359)
P (days)	n/a	12.8805	12.894
α (B1950)	13:23:06.85	13:23:09	13:23:08.8
δ (B1950)	-62 : 54 : 24.9	-62 : 54.0	-62 : 54 : 11

observations. Therefore the comparison can only be carried out on 0.7 square degree. Furthermore, as pointed above, this comparison is restricted to stars with magnitude $I_C < 14$. A total of 118 EROS objects, 6 GCVS and 23 CKS91 objects lie in this region. Three objects are common to the GCVS and CKS91 catalogues one. The overlap between the three catalogues represents only 19 objects. Among them one finds the most interesting ones, such as Cepheids OO CEN and V881 CEN (see Table 7). Some objects show a large difference between the two magnitude determinations. These are long period variables for which the mean magnitude is measured on only a part of the whole period, and thus ill determined in

both surveys. Only 7 known variable stars are not recovered by our analysis (see Table 8). Conversely, 99 objects of our catalogue are not listed by CKS91. All of them are labelled Long Period Variable stars (LPV).

5. Discussion

The motivation of our search was to improve our knowledge of the distance distribution and of the extinction of the microlensing source stars used in Papers I and II. In the following we use the RR Lyrae which are well-known distance indicators and have been observed in all six directions that we investigated. The GSA fields having a high

Table 7. Previously known variable stars recovered by this survey. For each star we give its name, the original catalogue, the type of variability, the period and the I magnitude given (if any) in the catalogue and in this survey.

Object	Catalogue	Type	P (days)	I	EROS ID	P (days)	I_c
OO Cen	GCVS	C	12.8805	9.96	J132630-630945 (tm5504m12359)	12.8941	10.20
V881 Cen	GCVS	C	-	10.57	J132721-630110 (tm5503l3554)	15.2278	10.97
V608 Cen	GCVS	EB	1.6287	12.95	J132948-630634 (tm5505m13922)	1.7601	12.05
CKS91 13246-6259	CKS91	EA	-	13.29	J132758-631449 (tm5505l9598)	5.4918	13.03
CKS91 13297-6244	CKS91	E	-	13.05 1	J133311-630023 (tm5511m5801)	0.2141	13.20
CKS91 13206-6305	CKS91	SR	-	12.35	J132400-632057 (tm5504l1003)	45.83	12.34
CKS91 13211-6300	CKS91	LPV	-	11.73	J132427-631542 (tm5504k5591)	$P > 60$	11.75
CKS91 13214-6306	CKS91	SR	-	10.51	J132444-632224 (tm5504l7786)	55.00	10.79
CKS91 13214-6256	CKS91	SR	-	12.85	J132447-631201 (tm5504k8841)	52.39	12.77
CKS91 13215-625	CKS91	LPV	-	10.43	J132454-631437 (tm5504k9938)	$P > 60$	10.23
CKS91 13219-6301	CKS91	LPV	-	11.81	J132519-631720 (tm5504l13138)	$P > 60$	13.00
CKS91 13229-6252	CKS91	LPV	-	12.51	J132433-632201 (tm5504m10330)	58.46	11.96
CKS91 13230-6251	CKS91	LPV	-	12.48	J132625-630645 (tm5504m11603)	$P > 60$	12.41
CKS91 13238-6227	CKS91	LPV	-	12.63	J132714-634338 (tm5507l1058)	$P > 60$	10.78
CKS91 13238-6254	CKS91	LPV	-	12.21	J132719-633644 (tm5507l1752)	$P > 60$	11.91
CKS91 13246-6306	CKS91	M	-	12.46	J132803-632200 (tm5505l10106)	$P > 60$	13.99
CKS91 13247-6244	CKS91	SR	-	12.66	J132813-630013 (tm5510k2095)	44.80	12.72
CKS91 13264-6256	CKS91	LPV	-	12.37	J132948-631227 (tm5510n4312)	$P > 60$	12.56
CKS91 13289-6308	CKS91	M	-	12.65	J133222-632339 (tm5513k12129)	$P > 60$	12.65

non-uniform absorption, we give only an average reddening towards our fields, based on EROS data alone.

For each selected RR Lyræ we estimate the extinction $A(V) = 2.4 \times E(V - I)$ using the standard extinction coefficients (Schlegel et al. 1998; Stanek 1996). The colour excess $E(V - I)$ is derived from the colour that we measured and their intrinsic colour: $(V - I)_0^{\text{RRab}} = 0.4$ and $(V - I)_0^{\text{RRc}} = 0.2$ (Alcock et al. 1998). The uncertainty on individual extinctions is estimated to be 0.5^{mag} . This error includes the 0.2^{mag} uncertainty on

the EROS colour measurement (see Eq. (1)) and a 0.1^{mag} magnitude uncertainty on the intrinsic colour. The absolute magnitude of RR Lyræ is $M_V^{\text{RR}} \sim 0.71$ with a precision of 0.07^{mag} (Gould & Popowski 1998). Finally, we estimate the distance to each star simply by using the relation $m - M = 5 \times (\log(d) - 1) + A$. The typical uncertainty on the distance is 20%. The left panels of Fig. 11 show the obtained extinction $A(V)$ versus the calculated distance. The mean extinction of the RR Lyræ is $A(V) \sim 3.6 \pm 0.2$ mag towards θ Mus and 3.3 ± 0.1 towards γ Nor, with a

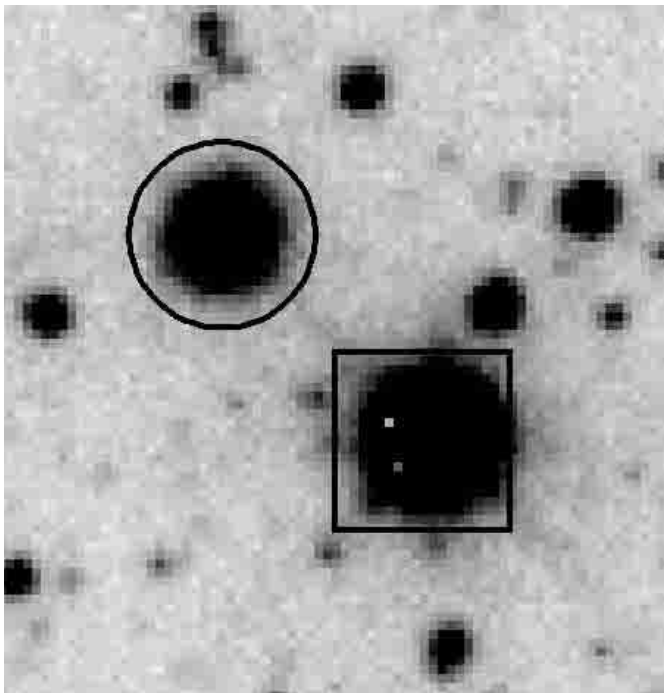


Fig. 10. Finding chart of the star OO Cen – J132630-630945(tm5504m12359). The position given by the SIMBAD database is shown by the square which is $10''$ wide, while the EROS position is shown by the circle.

Table 8. Previously known variable stars *not* recovered by this survey. For each of them we give (where possible) their type of variability, their period and the reason for their absence in the catalogue.

Object	Type	$P(\text{days})$	catalogue	comment
HQ Nor	EB	90.9	GCVS	too long period
HY Nor	Mira	236	GCVS	close to the CCDs gap
UW Nor	EA	8.4860	GCVS	fails cut 6
13214-6256	LPV		CKS91	fails cut 6
13218-6254	LPV		CKS91	fails cut 6
13248-6249	LPV		CKS91	fails cut 6
13232-6249	LPV		CKS91	fails cut 6

dispersion of 1 mag. The mean distance of the RR Lyræ is $\sim 4.7 \pm 0.3$ kpc towards θ Mus and 5.3 ± 0.2 kpc towards γ Nor. The dispersion of the values is 1.4 kpc which reflects the spread in distance of disc stars. The disc population contributes to the sources of the microlensing events that we observe in the Galactic plane. A model based on the distribution of matter in the disc and the luminosity function of neighboring stars has been used in Derue (1999) to estimate the distance of the star population from the disc. The distance of RR Lyræ is in good agreement with the one obtained with this model.

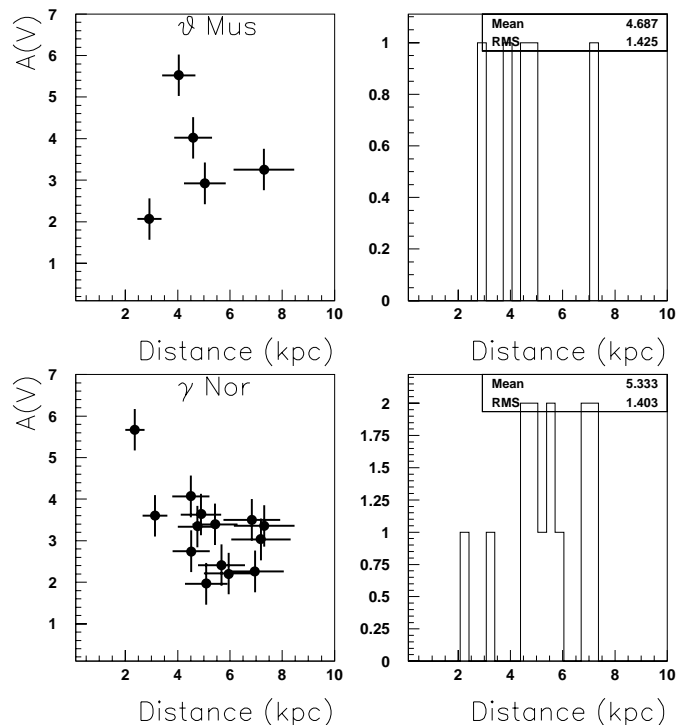


Fig. 11. Extinction $A(V)$ versus the distance of RR Lyræ (in kpc) towards θ Mus and γ Nor. The right panels show the inferred distance distribution.

6. Conclusion

In the course of our program dedicated to microlensing events, we have devoted a fraction of observing time to the search for variable stars in six directions of the Galactic plane. This exploratory campaign, that lasted three months, led to the discovery of 1362 variable stars. Among them we identified 9 Cepheids, 19 RR Lyræ, 34 Miras, 176 eclipsing binaries and 266 Semi-Regular variable stars. We have set up a catalogue of all of the 1362 stars and cross-identified it with several other catalogues. In particular a comparison with the GCVS and the CKS91 catalogues shows that only a small fraction ($\sim 15\%$) of the objects that we have identified appear in those two. Among the stars most appropriate to be used as distance indicators, the Cepheids turned out to be too few to warrant a particular study. As far as RR Lyræ are concerned, we have determined their mean distance and found it to be ~ 5 kpc. Yet the statistics being quite limited, we are considering pursuing this effort by launching a longer search for variable stars based on the results of this first campaign.

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