

Neutral hydrogen in dwarf galaxies

II. The kinematics of HI

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Abstract. This paper is the second in a series presenting a sample of 29 late-type dwarf galaxies observed with the Westerbork Synthesis Radio Telescope in the 21-cm line of neutral hydrogen (HI). Here we present rotation curves, maps of the velocity field and maps of the velocity dispersion across the sample galaxies.

Key words. galaxies: irregular – galaxies: dwarf – galaxies: kinematics and dynamics

1. Introduction

In this paper we present the kinematical information extracted from HI maps of the dwarf galaxy sample described by Stil & Israel (2002; hereafter Paper I). As neutral atomic hydrogen is one of the most extended observable components of a galaxy, its line emission provides an excellent tool to probe galaxy mass distributions. The dopplershift of a line profile is a direct measure of the projected rotation velocity at the position sampled. Its linewidth is a measure of the macroscopic chaotic motion of the gas and reflects its physical condition. Both rotation and velocity dispersion can be measured out to the edge of the HI distribution, which usually extends far beyond the stellar distribution. In the outer regions, rotational velocities are more or less constant, implying the presence of large amounts of matter even at the largest distances to the center. These rotation velocities are about three times higher than expected if only stars and gas would contribute to the mass; thus the amount of “unseen” or “dark” mass is about an order of magnitude higher than that associated with luminous matter.

It is of interest to determine the detailed kinematical conditions governing late-type dwarf galaxies. These galaxies are gas-rich, dynamically simple and relatively easy to observe. Their rotation curves trace the distribution of dark matter more directly than those of more massive galaxies, where the mass-to-light ratio of the stellar disk and bulge is a critical parameter. Interestingly, it has been suggested

that in some dwarf galaxies, the stellar contribution to the total mass is quite small (Broeils 1992), and that this contribution decreases with decreasing maximum rotation curve velocity (Persic et al. 1996; Swaters 1999).

2. Velocity field and rotation curve

We determined intensity-weighted mean velocities in order to make maximum use of the information contained in spectra with limited signal-to-noise ratios. Noise was further suppressed by only using data from the areas delineated by the cleaning masks (cf. Paper I). The intensity-weighted velocity fields of our galaxy sample are shown in Fig. 4.

We have determined rotation curves by iteratively fitting to the observed velocity field the parameters found with the tilted-ring method (Warner et al. 1973), incorporated in the GIPSY packages as the task ROTCUR (Begeman 1987). Note that the inclination and the location of the kinematic center can be determined only if the rotation curve flattens at the outer radii (i.e. shows the onset of differential rotation). In order to obtain robust solutions, we performed a large number of fits where each parameter was free in 25 to 30 fits, allowing calculation of the rms scatter of each parameter at all radii. Solution of $v(r) \sin i$ can also directly be compared to observed major axis position-velocity maps (Sect. 3). Although the product $v(r) \sin i$ is well-constrained, the separate solutions for $v(r)$ and i in general are not unique. In particular, at inclination angles $\leq 50^\circ$, $v(r)$ and i cannot be fitted independently (Begeman 1987).

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Table 1. HI disk parameters from tilted-ring fits; inclination fitted.

Name	resolution	kinematic center		incl.	fixed/free	PA	v_{sys}	scale
	arcsec	α_{1950}	δ_{1950}	degr.		degr.	km s ⁻¹	kpc/'
[1]	[2]	[3]	[4]	[5]	[6]	[7]	[8]	[9]
DDO 46	13.5	7 ^h 38 ^m 00 ^s .9 ± 0 ^s .4	40°13'37" ± 5"	45	fixed	270 ± 5	362 ± 2	1.4
DDO 47	13.5	7 ^h 39 ^m 03 ^s .1	16°55'13"	30	fixed	318 ± 8	272 ± 2	0.58
	30	7 ^h 39 ^m 03 ^s .1	16°55'13"	30	fixed	316 ± 9	272 ± 2	
DDO 48	13.5	7 ^h 54 ^m 46 ^s .4 ± 0 ^s .2	58°10'43" ± 3"	75 ± 1	free	356 ± 1	1087 ± 1	4.6
	30	7 ^h 54 ^m 46 ^s .4 ± 0 ^s .3	58°10'43" ± 10"	80 ± 5	free	356 ± 2	1088 ± 2	
NGC 2537	30	8 ^h 09 ^m 42 ^s .6 ± 0 ^s .4	46°08'40" ± 4"	42 ± 3	free	174 ± 6	444 ± 1	1.9
UGC 4278	13.5	8 ^h 10 ^m 27 ^s .4 ± 0 ^s .5	45°53'52" ± 4"	80 ± 3	free	351 ± 1	564 ± 1	1.9
NGC 2976	13.5	9 ^h 43 ^m 08 ^s .0 ± 1 ^s .0	68°08'57" ± 5"	65 ± 3	free	326 ± 2	4 ± 2	1.0
	30	9 ^h 43 ^m 07 ^s .7 ± 0 ^s .4	68°08'52" ± 4"	62 ± 2	free	325 ± 4	4 ± 1	
DDO 83	13.5	10 ^h 33 ^m 54 ^s .4 ± 0 ^s .1	31°48'24" ± 2"	66 ± 2	free	59 ± 3	582 ± 2	2.6
DDO 87	30	10 ^h 46 ^m 17 ^s .0 ± 0 ^s .5	65°47'35" ± 5"	63 ± 4	free	239 ± 3	338 ± 2	1.0
DDO 123	13.5	12 ^h 23 ^m 47 ^s .0 ± 0 ^s .3	58°35'51" ± 7"	25	fixed	201 ± 4	722 ± 1	3.3
DDO 133	30	12 ^h 30 ^m 26 ^s .8 ± 1 ^s .2	31°48'46" ± 15"	20	fixed	353 ± 4	330 ± 3	1.5
DDO 168	30	13 ^h 12 ^m 15 ^s .0 ± 1 ^s .5	46°11'24" ± 6"	63 ± 3	free	276 ± 1	190 ± 3	1.0
DDO 185	30	13 ^h 52 ^m 53 ^s .3 ± 0 ^s .5	54°08'30" ± 4"	66 ± 5	free	18 ± 2	140 ± 2	2.0
DDO 217	30	23 ^h 27 ^m 32 ^s .1 ± 1 ^s .1	40°42'56" ± 15"	46 ± 5	free	35 ± 3	428 ± 2	2.7

Column designations: [1] Object name; [2] resolution of dataset used; [3] and [4] right ascension and declination (epoch 1950) of kinematic center and their rms scatter between radii; [5] inclination in degrees and its error if it was a free parameter in the fit; [6] flag indicating whether the inclination was a free parameter; [7] position angle in degrees and its rms scatter between radii; [8] heliocentric systemic velocity in km s⁻¹ and its rms scatter between radii; [9] radial scale of the rotation curve in kpc/arcmin.

Notes: DDO 168 position angle is mean for radii 150" and 180"; DDO 217 position angle is mean for radii larger than 150 arcsec.

Depending on galaxy HI extent and surface brightness, we performed fits on full-resolution (13.5") or low-resolution (30") data. As a consistency check, we used both low and full resolution data for DDO 47, DDO 48 and NGC 2976. The results are presented in Tables 1 and 2. In total, 13 out of 29 galaxies in the sample could so be fitted. For the remainder, a complete analysis was not feasible. There, we required the rotation center to coincide with the HI center of mass which was found to be correct at least for all the objects listed in Table 1. In addition, we assigned to these objects a low (30°), average (60°) or high (80°) inclination based on the appearance of the HI isophotes. The results of fits to the high-resolution data, but restricted by these assumptions, are given in Tables 3 and 4. The center coordinates in Table 3 are those of the HI-intensity-weighted mean position in α_{1950} and δ_{1950} . They typically change only by a few arcseconds if different intensity thresholds are applied. The errors quoted for position angle and systemic velocity represent their scatter between different radii. Note that the velocities listed in Table 4 have *not* been corrected for inclination. These $v(r) \sin i$ values, as indeed also the position angle, depend only weakly on the assumed inclination.

3. Major axis position-velocity diagrams

The projected rotation curve, $v(r) \sin i$, can be visualized as a position-velocity slice. We include such maps in Fig. 4, which show intensity as a function of position along the

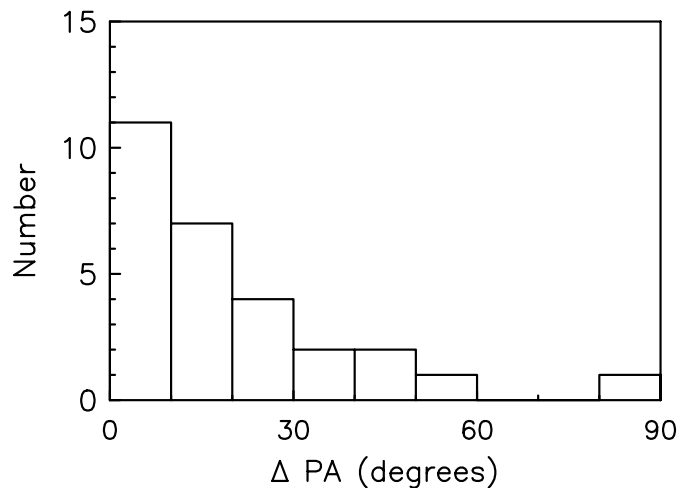


Fig. 1. Histogram of the difference in position angle between the major axis of the $N_{\text{HI}} = 3 \times 10^{20} \text{ cm}^{-2}$ isophote and the velocity gradient for 27 objects (Mkn 178 and DDO 101 were excluded because of missing data). The angle ΔPA is the smallest angle between the velocity gradient and the major axis. The median difference is $\Delta\text{PA} = 15^\circ$. The object with $\Delta\text{PA} = 90^\circ$ is NGC 2537.

kinematic major axis and heliocentric velocity. Position angles and central positions were taken from Tables 1 or 3, in order of preference. We have also included in Fig. 4 the rotation curve points listed in Tables 2 and 4.

Table 2. Rotation velocities (km s⁻¹) from tilted ring fits.

radius "	DDO 46 HR	DDO 47 HR	DDO 47 LR	DDO 48 HR	DDO 48 LR	NGC 2537 LR	UGC 4278 HR	NGC 2976 HR	NGC 2976 LR
15	15.8 ± 5.2	20.6 ± 0.1		30.0 ± 7.6			8 ± 5	20.9 ± 5	
30	36.0 ± 4.4	24.5 ± 1.5	24.4 ± 0.4	47.5 ± 0.3	40.9 ± 3.9	46.2 ± 2.0	25.6 ± 1.5	28.8 ± 3	29.6 ± 1.3
45	41.2 ± 3.2	27.1 ± 3.0		59.4 ± 0.3			40.1 ± 0.5	40.4 ± 1.8	
60	44.3 ± 3.7	25.7 ± 1.6	27.1 ± 0.5	66.8 ± 0.4	63.3 ± 9.8	53.4 ± 4.3	50.1 ± 1.4	49.6 ± 2.8	51.6 ± 0.3
75	46.0 ± 3.7	30.4 ± 1.0		72.0 ± 0.2			59.3 ± 0.5	61.5 ± 0.4	
90	44.8 ± 6.6	36.1 ± 0.5	36.4 ± 0.5	74.9 ± 0.1	74.3 ± 6.7	63.2 ± 3.3	68.5 ± 1.1	69.1 ± 1.2	69.2 ± 0.1
105		41.7 ± 0.5		76.6 ± 0.5			76.4 ± 0.3	71.1 ± 0.4	
120		48.2 ± 1.1	47.9 ± 0.5			71.9 ± 5.2	82.5 ± 2.5	71.5 ± 0.2	71.7 ± 0.3
135		55.0 ± 0.5					87.4 ± 2.1		
150		62.0 ± 0.8	60.6 ± 0.4				87.6 ± 1.8		
165		66.4 ± 0.7							
180			68.0 ± 1.3						
radius "	DDO 83 HR	DDO 87 LR	DDO 123 HR	DDO 133 LR	DDO 168 LR	DDO 185 LR	DDO 217 LR		
15	23.7 ± 7.5		26.1 ± 0.9						
30	35.6 ± 0.4	17.6 ± 1.3	35.5 ± 4.8	34.6 ± 5.7	11.2 ± 2.0	17.8 ± 4.7	26.1 ± 7.5		
45	46.1 ± 0.2		50.8 ± 0.4						
60	51.7 ± 0.3	30.0 ± 0.6	41.8 ± 0.2	50.3 ± 14	26.6 ± 2.0	28.1 ± 0.4	58.7 ± 6.0		
75	51.0 ± 3.0		49.3 ± 0.8						
90		32.3 ± 0.4	55.5 ± 0.2	65.8 ± 11	29.7 ± 1.1	40.1 ± 0.5	66.2 ± 3.8		
105			59.9 ± 0.7						
120		34.4 ± 1.8	63.1 ± 3.0	75.7 ± 2.1	39.4 ± 1.4	47.3 ± 0.3	68.1 ± 3.3		
150				75.1 ± 4.2	42.4 ± 1.0	51.9 ± 2.3	72.5 ± 2.8		
180					44.3 ± 0.5		74.9 ± 3.3		
210							77.0 ± 0.8		
240							78.1 ± 2.1		

Note: HR is based on high-resolution (13.5") fits; LR is based on low-resolution (30") fits.

Virtually all of the observed galaxies show a velocity gradient. However, in poorly-ordered velocity fields, it is not clear whether or not this rotation represents. In Fig. 1 we show, however, that observed velocity gradients tend to align with the major axis of the HI isophotes, although a few significant misalignments exist suggesting a rotating disk (e.g. DDO 47, DDO 63, DDO 165 and also Sextans A as shown by Skillman et al. 1988). We suspect that much of the width of the distribution in Fig. 1 is caused by errors in the determination of major axis position angle from the HI distribution, although the magnitude of ΔPA is not correlated with the total brightness of the galaxy in HI (or in blue light).

The position-velocity maps of three very-low luminosity dwarf galaxies ($M_B > -14$) betray high ratio of rotational to random velocities: those of DDO 47, DDO 52 and DDO 87. The existence of rotationally supported dwarfs of such low luminosity is remarkable if simple arguments based on the luminosity-linewidth relation for large spirals are applied to dwarf galaxies (Lo et al. 1993; Stil & Israel, in preparation). In addition, the rotation curves of DDO 52 and DDO 87 show clear signs of flattening. They illustrate the increase in the number of extremely low luminosity galaxies known to be largely supported by

rotation (cf. Carignan & Beaulieu 1989; Côté 1997). Indeed, with the exception of DDO 47 Fig. 4 as well as Tables 2 and 4 include few convincing examples of pure solid-body rotation.

4. Velocity dispersion

The linewidth in a single synthesized beam carries information about the velocity dispersion of the ensemble of HI clouds contained within the beam. This information is, unfortunately, confused with the signatures of other effects such as warps or flares of the HI layer, the extent to which HI structures are resolved, and the rotation velocity gradient over the beam. Warps and flares bring HI from different galactocentric radii into the line of sight, but they only cause confusion in highly inclined galaxies.

Excluding lines of sight with low signal-to-noise ratios, we have used Gaussian fitting to obtain the mean velocity dispersions tabulated in Table 5. To first order, we corrected the maps for finite velocity resolution and linear velocity gradients over the beam according to the relation

$$\sigma^2 = \sigma_{\text{obs}}^2 - \sigma_{\text{inst}}^2 - \frac{1}{2}b^2(\nabla v)^2$$

Table 3. HI disk parameters from restricted tilted-ring/HI isophote fits; inclination assumed.

Name	HI center of mass		i	PA _{kin}	v_{sys}	D_{HI}	q_{HI}	PA _{HI}	Scale
	α_{1950}	δ_{1950}							
[1]	[2]	[3]	[4]	[5]	[6]	[7]	[8]	[9]	[10]
D 22	2 ^h 29 ^m 47 ^s .2	38°27'35"	80	178 ± 1	564 ± 2	157 ± 9	0.47 ± 0.04	181 ± 2	2.9
D 43	7 ^h 24 ^m 50 ^s .2	40°52'19"	30	296 ± 4	355 ± 1	145 ± 3	0.93 ± 0.03	335 ± 11	1.4
D 46*	7 ^h 38 ^m 00 ^s .9	40°13'30"	30	273 ± 4	363 ± 2	189 ± 5	0.97 ± 0.05	226 ± 45	1.4
D 47*	7 ^h 39 ^m 03 ^s .1	16°55'13"	30	319 ± 8	272 ± 1	403 ± 4	0.85 ± 0.01	266 ± 3	0.58
D 48*	7 ^h 54 ^m 46 ^s .8	58°10'39"	80	357 ± 3	1086 ± 1	245 ± 12	0.33 ± 0.02	354 ± 2	4.6
N 2537*	8 ^h 09 ^m 42 ^s .3	46°08'32"	30	171 ± 2	446 ± 1	219 ± 5	0.80 ± 0.04	261 ± 6	1.9
D 52	8 ^h 25 ^m 06 ^s .4	42°01'17"	60	5 ± 3	394 ± 1	204 ± 9	0.58 ± 0.04	359 ± 3	1.5
D 63	9 ^h 36 ^m 01 ^s .9	71°25'07"	30	38 ± 15	140 ± 2	307 ± 6	0.76 ± 0.02	5 ± 2	1.0
N 2976*	9 ^h 43 ^m 11 ^s .7	68°09'06"	60	321 ± 2	3 ± 2	316 ± 5	0.62 ± 0.01	326 ± 1	1.0
D 64	9 ^h 47 ^m 26 ^s .3	31°43'19"	60	97 ± 7	517 ± 1	210 ± 11	0.56 ± 0.05	84 ± 4	1.8
D 68	9 ^h 53 ^m 52 ^s .8	29°03'42"	60	29 ± 10	504 ± 3	300 ± 9	0.56 ± 0.03	19 ± 2	1.8
D 73	10 ^h 06 ^m 39 ^s .5	30°23'50"	60	66 ± 3	1378 ± 2	135 ± 3	0.94 ± 0.03	38 ± 15	5.2
D 83*	10 ^h 33 ^m 54 ^s .7	31°48'24"	60	55 ± 3	584 ± 1	190 ± 4	0.55 ± 0.02	56 ± 2	2.6
D 87*	10 ^h 46 ^m 16 ^s .3	65°47'34"	60	239 ± 5	339 ± 1	246 ± 5	0.86 ± 0.03	224 ± 6	1.0
M 178	11 ^h 30 ^m 45 ^s .1	49°31'06"	–	–	–	91 ± 4	0.70 ± 0.06	146 ± 7	1.5
N 3738	11 ^h 33 ^m 04 ^s .7	54°48'10"	60	270	225 ± 4	182 ± 6	0.80 ± 0.05	274 ± 10	1.5
D 101	11 ^h 53 ^m 07 ^s .1	31°47'41"	–	–	–	–	–	–	2.1
D 123*	12 ^h 23 ^m 45 ^s .9	58°35'52"	18	198 ± 5	723 ± 1	240 ± 2	0.97 ± 0.01	197 ± 13	3.3
M 209	12 ^h 23 ^m 50 ^s .2	48°46'25"	30	225	285 ± 8	125 ± 5	0.86 ± 0.08	206 ± 20	1.4
D 125	12 ^h 25 ^m 14 ^s .0	43°46'17"	60	135 ± 2	195 ± 1	281 ± 3	0.66 ± 0.02	118 ± 2	1.3
D 133*	12 ^h 30 ^m 27 ^s .6	31°48'54"	30	350 ± 5	331 ± 1	296 ± 5	0.86 ± 0.02	18 ± 5	1.5
D 165	13 ^h 04 ^m 40 ^s .2	67°58'20"	–	135	22	253 ± 5	0.79 ± 0.03	93 ± 4	1.3
D 166	13 ^h 11 ^m 00 ^s .0	36°28'36"	30	41 ± 9	942 ± 3	200 ± 3	0.75 ± 0.02	46 ± 3	4.7
D 168*	13 ^h 12 ^m 14 ^s .7	46°11'07"	60	276 ± 1	191 ± 2	366 ± 5	0.82 ± 0.02	304 ± 3	1.0
D 185*	13 ^h 52 ^m 53 ^s .5	54°08'24"	80	39 ± 5	137 ± 3	302 ± 13	0.37 ± 0.02	17 ± 1	2.0
D 190	14 ^h 22 ^m 48 ^s .2	44°45'09"	60	149 ± 6	149 ± 1	201 ± 4	0.86 ± 0.03	150 ± 7	1.7
D 216	23 ^h 26 ^m 02 ^s .9	14°28'02"	60	135	–	286 ± 17	0.32 ± 0.02	130 ± 1	0.29
D 217*	23 ^h 27 ^m 32 ^s .8	40°43'03"	30	43 ± 8	432 ± 2	472 ± 5	0.75 ± 0.01	27 ± 2	2.7

Column designations [1] Object name; [2] and [3] right ascension and declination (epoch 1950) of kinematic center and its rms scatter between radii; [4] inclination in degrees (see text for explanation); [5] position angle of kinematic major axis in degrees and its rms scatter between radii; [6] systemic velocity in km s⁻¹ and its rms scatter between radii; [7] major axis of ellipse fitted to the $N_{\text{HI}} = 3 \times 10^{20}$ cm⁻² contour in the HI column density map; [8] axial ratio of the ellipse in [7]; [9] position angle of the ellipse major axis in [7]; [10] radial scale of the rotation curve in kpc/arcmin.

Notes: If no error is given, the value is a best estimate. For galaxies marked with an asterisk, more accurate values are given in Table 1; they are included here only for comparison purposes.

where σ_{obs} is the dispersion of a Gaussian $e^{-\frac{1}{2}v^2/\sigma_{\text{obs}}^2}$ fitted to the line profile at each position, $\sigma_{\text{inst}} = 2 \cdot 0.8493 \cdot \Delta v$ is the dispersion of a Gaussian corresponding to the velocity resolution of the Hanning-smoothed data, and ∇v is the local velocity gradient over the beam, assumed to be of the form e^{-x^2/b^2} . We calculated the velocity gradient at each position from model velocity fields constructed with the rotation curves presented in the previous section. No correction for galaxy inclination was applied. The procedure is described in more detail in Appendix A.

The high velocity dispersion in NGC 3738 is probably an artifact resulting from too small a beam/velocity gradient correction caused by the marginally resolved steep velocity gradient of the galaxy. Excluding NGC 3738, we find a mean velocity dispersion of 8.6 ± 0.34 km s⁻¹ at 13'' resolution and 9.5 ± 0.38 km s⁻¹ at 27'' resolution. Although the difference between the two results is

significant, its magnitude is only 10%. Moreover, these values are consistent with those in the literature (Shostak & van der Kruit 1984; Skillman et al. 1988). As the physical area contributing to the measured velocity dispersion increases with galaxy distance squared, we show in Fig. 2 mean velocity dispersions as a function of beam surface area for both the 13'' and 27'' resolution maps. The systematic increase in mean velocity dispersion with physical beam area increasing by two orders of magnitude is no more than about 2 km s⁻¹. As our results are thus effectively insensitive to linear resolution on scales of ~ 0.1 kpc and larger, we may compare the velocity dispersions of galaxies at various distances without fear of introducing large systematic effects.

Finally, we show mean velocity dispersions as a function of inclination in Fig. 3. We used tilted-ring inclinations and, lacking these, inclinations estimated from

Table 4. Velocity $v(r) \sin(i)$ from restricted rotation curve fits.

radius	DDO 22	DDO 43	DDO 46*	DDO 47*	DDO 48*	NGC 2537*	DDO 52	DDO 63	NGC 2976*
"									
30	8.1 ± 2.4	11.7 ± 2.0	23.7 ± 3.3	12.2 ± 2.0	42.4 ± 6.5	33.6 ± 3.2	23.7 ± 3.6	6.3 ± 2.3	25.6 ± 5.6
60	18.7 ± 2.8	15.7 ± 2.8	29.9 ± 2.1	13.4 ± 2.3	62.7 ± 5.0	38.9 ± 2.5	32.7 ± 2.6	8.0 ± 1.9	44.7 ± 3.8
90		17.5 ± 3.9	30.2 ± 3.6	18.3 ± 1.7	73.0 ± 5.0	43.2 ± 3.5	38.2 ± 2.5	7.1 ± 2.6	57.2 ± 5.3
120				24.1 ± 2.2			43.3 ± 5.0	8.6 ± 3.2	61.0 ± 7.6
150				30.6 ± 2.1				8.0 ± 5.3	54.0 ± 11
radius	DDO 64	DDO 68	DDO 73	DDO 83*	DDO 87*	NGC 3738	DDO 123*	Mk 209	DDO 125
"									
30	22.0 ± 3.1	20.5 ± 4.3	20.0 ± 1.6	32.6 ± 4.8	16.9 ± 3.3	39.0 ± 6.2	14.0 ± 2.5	14.3 ± 5.0	5.4 ± 1.8
60	33.4 ± 2.9	28.4 ± 5.7	26.6 ± 2.4	43.7 ± 2.5	26.7 ± 1.8		18.7 ± 1.7	17.1 ± 5.0	8.7 ± 1.1
90	41.8 ± 7.4	42.9 ± 5.6		46.7 ± 3.3	28.8 ± 2.9		22.6 ± 2.6		9.8 ± 1.6
120		45.9 ± 5.6			30.2 ± 4.0		27.5 ± 2.7		11.2 ± 2.7
150					31.0 ± 5.8				
radius	DDO 133*	DDO 165	DDO 166	DDO 168*	DDO 185*	DDO 190	DDO 216	DDO 217*	
"									
30	11.5 ± 3.0	9	22.6 ± 4.3	10.7 ± 2.5	11.7 ± 2.2	12.3 ± 1.8	1 ± 1	24.1 ± 4.7	
60	17.7 ± 1.9	17	30.2 ± 3.1	20.3 ± 2.4	22.9 ± 3.0	19.4 ± 2.1	4 ± 2	38.1 ± 3.1	
90	22.9 ± 2.8	26	35.4 ± 5.0	24.8 ± 2.2	36.2 ± 3.2	24.7 ± 2.5	7 ± 5	41.3 ± 2.8	
120	26.6 ± 2.2			35.6 ± 5.5			7 ± 5	43.4 ± 2.9	
150				37.3 ± 6.7				47.0 ± 2.0	

Notes: see text for fit procedure details. Velocities have *not* been corrected for inclination, as opposed to results presented in Table 2. For galaxies marked with an asterisk, more accurate values are given in Table 2; values given here are only for comparison purposes. The rotation velocity of DDO 165 was determined from the position-velocity map in position angle 135° .

Table 5. Mean velocity dispersions from single Gaussian fits.

Name	$\langle \sigma \rangle_{13''}$	$\langle \sigma \rangle_{27''}$	Name	$\langle \sigma \rangle_{13''}$	$\langle \sigma \rangle_{27''}$	Name	$\langle \sigma \rangle_{13''}$	$\langle \sigma \rangle_{27''}$
[1]	[2]	[3]	[1]	[2]	[3]	[1]	[2]	[3]
	km s^{-1}	km s^{-1}		km s^{-1}	km s^{-1}		km s^{-1}	km s^{-1}
DDO 22	12.0	11.2	DDO 64	10.8 ± 3.3	10.9 ± 2.5	DDO 125	6.2	7.3 ± 1.5
DDO 43	8.4 ± 2.2	9.0 ± 1.7	DDO 68	10.9 ± 3.8	12.6 ± 4.4	DDO 133	–	8.2 ± 2.0
DDO 46	9.3 ± 2.3	10.0 ± 2.2	DDO 73	5.8	8.5 ± 2.8	DDO 165	9.3 ± 2.9	12.0 ± 3.8
DDO 47	8.2 ± 1.9	8.7 ± 3.2	DDO 83	9.9 ± 2.4	10.0 ± 2.5	DDO 166	9.1	11.6 ± 3.4
DDO 48	9.3 ± 3.1	10.1	DDO 87	–	6.0 ± 2.0	DDO 168	9.9 ± 2.9	10.6 ± 3.3
NGC 2537	7.2	11.5 ± 5.2	Mkn 178	–	7.6	DDO 185	8.1 ± 1.6	8.8 ± 1.7
DDO 52	6.8 ± 2.8	7.4 ± 2.3	NGC 3738	(12.2)	(18.2)	DDO 190	9.2 ± 2.4	10.0 ± 2.4
DDO 63	7.1 ± 2.1	8.9 ± 2.3	DDO 123	7.4 ± 2.5	9.0 ± 1.9	DDO 216	6.3 ± 1.7	5.4 ± 2.6
NGC 2976	11.1 ± 3.5	11.8 ± 3.2	Mkn 209	–	11.5	DDO 217	7.5 ± 2.7	8.6 ± 2.7

Notes: results are corrected for finite velocity resolution and velocity gradients over the beam as discussed in the text. Velocity dispersion scatter values are given only if the area considered was at least ten times the synthesized beam area. Column [2] gives mean velocity dispersion and rms scatter of the velocity dispersion over the galaxy from the full-resolution data, and Col. [3] the corresponding mean from the low-resolution data.

optical axial ratios assuming an intrinsic axial ratio of 0.15. The result is not sensitive to the exact value of this intrinsic axial ratio. For instance, use of the higher values suggested by the work of Staveley-Smith et al. (1992), increases $\sin i$ values by at most 0.08. The upper and lower panels in Fig. 3 show velocity dispersions before and after the correction for inclination. The upper panel shows velocity dispersions increasing at the highest inclinations ($\sin i > 0.9$; $i > 65^\circ$). This increase has disappeared completely in the corrected set in the lower panel. The widths

of local line profiles depend on inclination only through the observed velocity gradient over the beam, which is steeper on average for high inclination angles. Thus, *all* observed galaxies, irrespective their absolute luminosity ($-12.8 \text{ mag} > M_B > -17.6 \text{ mag}$) are have mean velocity dispersions of about 10 km s^{-1} , very similar to that of spiral galaxy disks. We will return to this result in a forthcoming paper.

The velocity dispersion maps shown in Fig. 4 are corrected for the local velocity gradient over the beam.

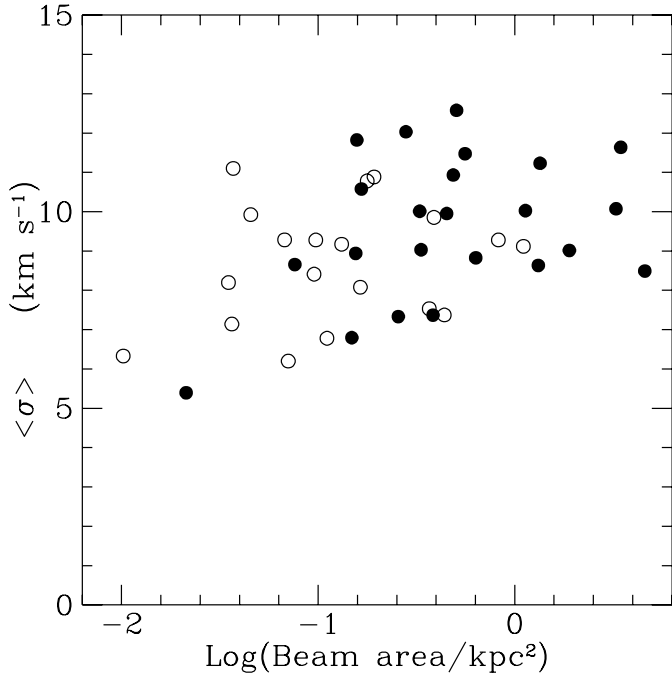


Fig. 2. The mean velocity dispersion in the low resolution maps is plotted as a function of the area of the synthesized beam. The velocity dispersions have been corrected for instrumental resolution and the velocity gradient over the beam according to Eq. (4). The beam area is defined as $\frac{1}{4}\pi b_\alpha b_\delta$, where b_α and b_δ are the FWHM beamsizes in right ascension and declination. Open symbols represent the 13'' resolution data, filled symbols are the 27'' resolution data. The pair of points with $\log(\text{Beam area}) < -1.6$ represents the Local Group dwarf DDO 216.

Note that the steep inner rotation curves of NGC 2537 and NGC 3738 are not completely resolved, resulting in artificially large linewidths.

5. Individual objects

DDO 22: The HI axial ratio suggests that this galaxy is seen at a high inclination. A region of high HI column density is present in the southern side of the galaxy. The kinematic and HI major axes appear to be misaligned.

DDO 43: The HI axial ratio indicates a face-on orientation. Parallel velocity contours in a regular velocity field mark a rapidly rising rotation curve. The kinematic and optical major axes are roughly perpendicular.

DDO 46: The HI distribution shows a U-shaped high-column density ridge. The velocity field is reasonably symmetric with respect to the kinematic minor axis, but asymmetric with respect to the major axis. The rotation curve is almost flat at the edge of the HI distribution.

DDO 47: Spiral structure in the outer HI disk was reported by Puche & Westpfahl (1994). No optical emission appears associated with the spiral arms. Regularly shaped, elliptical outer HI isophotes suggest a disk seen at low inclination. A number of high column density regions are

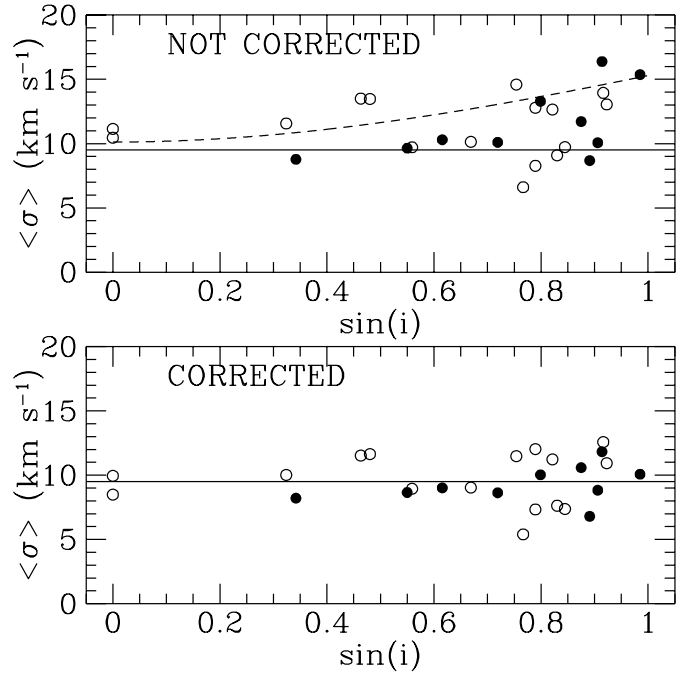


Fig. 3. The mean velocity dispersion as a function of $\sin(i)$ for the low resolution maps. Closed symbols are objects for which the inclinations were determined with a tilted ring fit. The inclination of the remaining objects (open symbols) was calculated from the optical axial ratios listed in Melisse & Israel (1994) and an intrinsic axial ratio of 0.15. The upper panel shows the uncorrected mean velocity dispersions. The lower panel shows the corrected values. The dashed line in the upper panel is the relation expected for a galaxy with a solid body rotation curve with a slope $1 \text{ km s}^{-1} \text{ arcsec}^{-1}$ (if seen edge-on) and a velocity dispersion of 9.5 km s^{-1} .

distributed evenly over the disk. There is a deep hole in the HI distribution at $\alpha = 7^{\text{h}}39^{\text{m}}8^{\text{s}}.1$, $\delta = 16^{\circ}54'31''$. The velocity field presents one of the few clear examples of solid-body rotation, as does the major axis XV diagram.

DDO 48: This is probably a nearly edge-on disk with a regular velocity field. The rotation curve rises rapidly near the center and flattens gradually outwards.

NGC 2537: The HI distribution consists of a U-shaped high column-density ridge. The velocity field, irregular in the north, exhibits the characteristic spider shape of a flattening rotation curve, also evident in the XV map. NGC 2537 is sometimes classified as a Blue Compact Dwarf (BCD) galaxy and is also known as the *Bear Paw Galaxy* (Schorn 1988). UGC 4278 is a nearby companion. NGC 2537A occurring just east of NGC 2537 on PSS plates is not visible in the HI data. The large linewidths in the center are an artifact of the unresolved velocity gradient.

UGC 4278: The rotation curve may show a turnover. NGC 2537 is $16'.6$ away at $\text{PA} = 328^{\circ}$.

DDO 52: Most of the HI is in a low-column-density disk, with a peak $N_{\text{HI}} = 1.5 \times 10^{21} \text{ cm}^{-2}$. The velocity field is regular with a hint of a flattening rotation curve. The XV

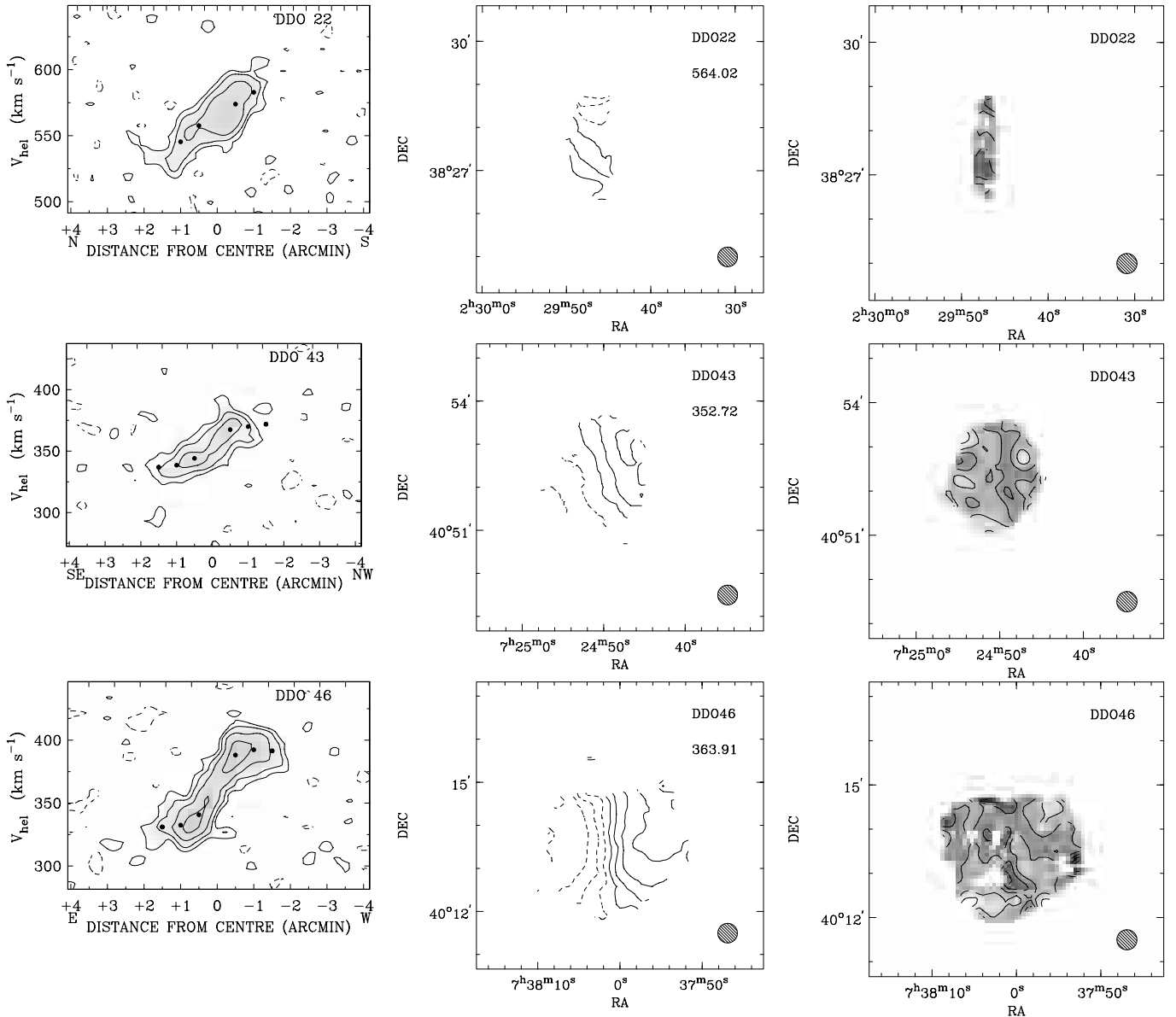


Fig. 4. For each galaxy, all maps at $27''$ resolution: (left) major axis position-velocity maps with contours at -2σ $2^N\sigma$, with $N = 0, 1, 2, \dots$ and σ the rms noise in empty channel maps; grayscales are logarithmic; tickmarks on the upper horizontal axis mark intervals of 1 kpc. (center) velocity fields with contour intervals of 8 km s^{-1} ; dashed contours mark approaching velocities; the first solid contour is at the system velocity which is printed under each object's name; WSRT beamsize is indicated by the hatched ellipse. (right): velocity dispersion maps with contour levels at 5, 7.5, 10, 12.5 $\dots \text{ km s}^{-1}$; grayscales are linear from 5 km s^{-1} to 20 km s^{-1} .

diagram shows a rather high ratio of rotational to random velocity, despite its small amplitude.

DDO 63: The HI is concentrated in a ring with a high column density, with a five times lower central minimum ($N_{\text{HI}} = 2.4 \times 10^{20} \text{ cm}^{-2}$) at $\alpha = 9^{\text{h}}36^{\text{m}}3^{\text{s}}.1$ $\delta = 71^{\circ}24'44''$. The kinematic and HI major axes are misaligned by about 30° . The rotation velocity is comparable to the HI velocity dispersion. The ring dominates the XV map. DDO 63 has also been studied by Puche & Westpfahl (1994) and Tully et al. (1978). The latter find the same kinematic/HI axis misalignment, and a comparable well-ordered, small-amplitude rotation velocity.

NGC 2976: The outer HI isophotes are well represented by ellipses. The two continuum sources (Paper I) are near to but not precisely coincident with the high HI column density regions ($N_{\text{HI}} \approx 3.5 \times 10^{21} \text{ cm}^{-2}$) on either side of the galaxy at $\alpha = 9^{\text{h}}43^{\text{m}}1^{\text{s}}.2$, $\delta = 68^{\circ}9'44''$ (NW) and $\alpha = 9^{\text{h}}43^{\text{m}}19^{\text{s}}.1$, $\delta = 68^{\circ}7'49''$ (SE). The emission at upper left in the velocity field map is unrelated Galactic foreground HI. The rotation curve flattens near the edge of the HI disk. A detailed optical study of NGC 2976 was presented by Bronkalla et al. (1992).

DDO 64: A high-column-density ridge dominates the HI structure of this probably nearly edge-on galaxy. The velocity field is poorly resolved along the minor

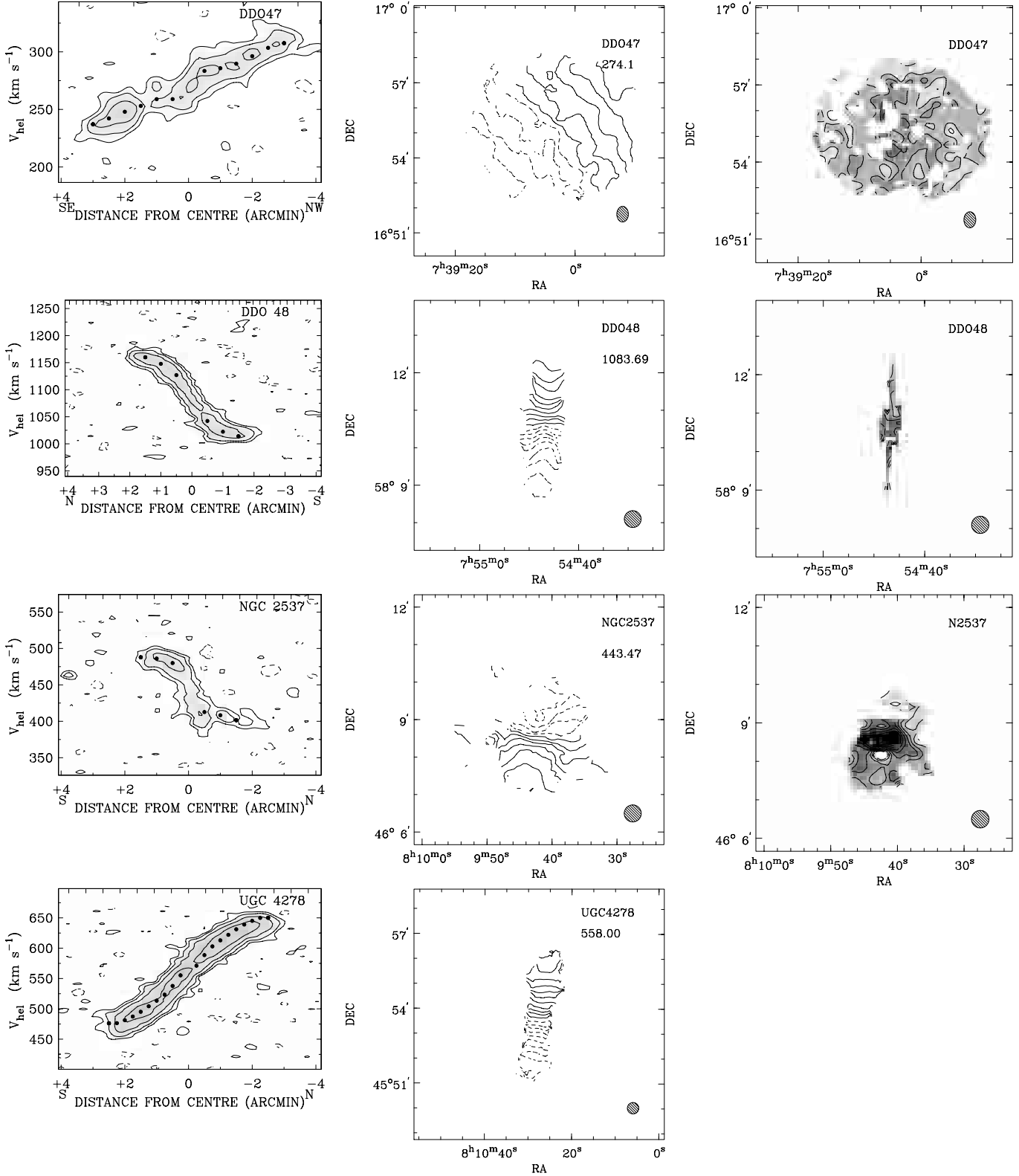


Fig. 4. continued.

axis. The XV map shows nearly solid-body rotation with only a hint of flattening. The feature south of DDO 64 is UGC 5272B, the feature north of it is probably noise.

DDO 68: At low column density levels, DDO 68 is reasonably symmetrical. High-column-density regions are found

in the north and east of the galaxy. A deep hole devoid of HI emission is located at $\alpha = 9^{\text{h}}53^{\text{m}}52^{\text{s}}.2$, $\delta = 29^{\circ}4'15''$. Low signal-to-noise regions contribute to the irregular appearance of the velocity field. The ratio of rotational to random velocity is low.

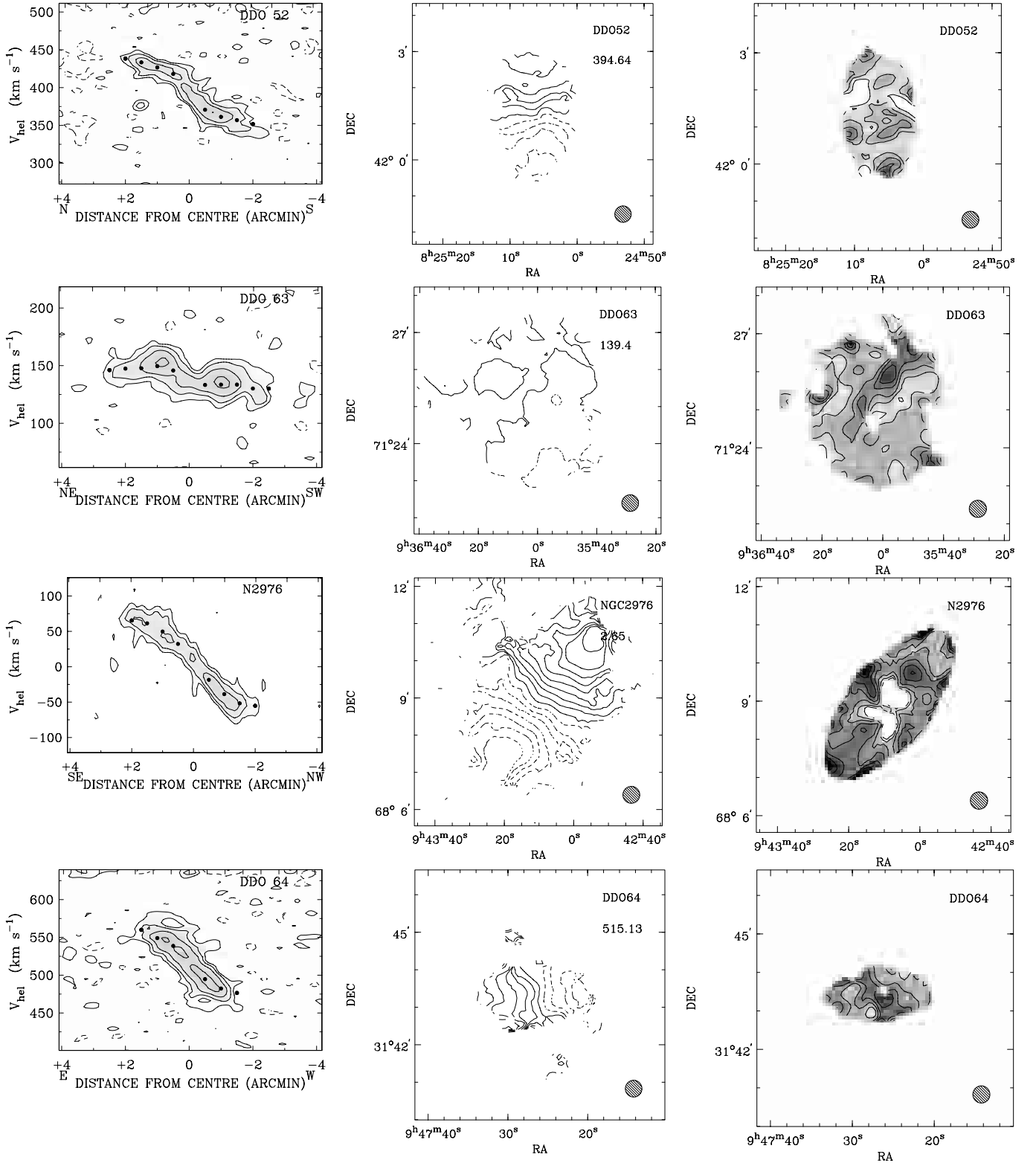


Fig. 4. continued.

DDO 73: The HI isophotes suggest a nearly face-on orientation. However, both velocity field and XV map indicate projected rotational velocities considerably in excess of the HI velocity dispersion.

DDO 83: The HI column density is high throughout the galaxy. The velocity field is regular and spider-like. The rotation curve rises strongly near the center and flattens at the edge of the HI disk.

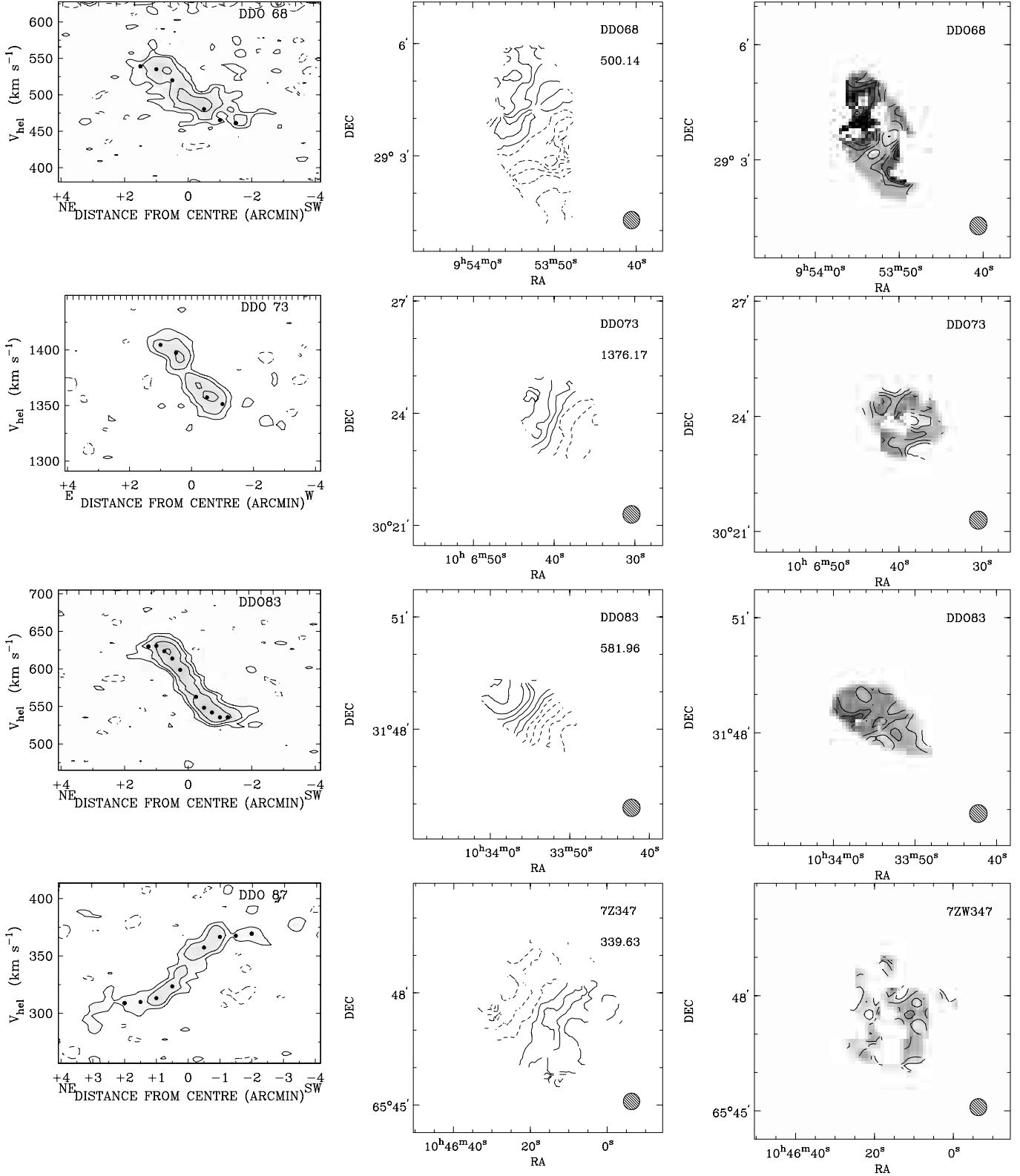


Fig. 4. continued.

DDO 87: The HI is distributed over a low-column-density disk with a number of small high column density regions, unresolved at 13'' resolution. The fragmented appearance of the velocity field is the result of low signal-to-noise

ratios. The XV map shows a nearly flat rotation curve in the outer regions. The low luminosity of DDO 87 is inferred from its association with the M 81 group (Huchtmeier & Skillman 1998).

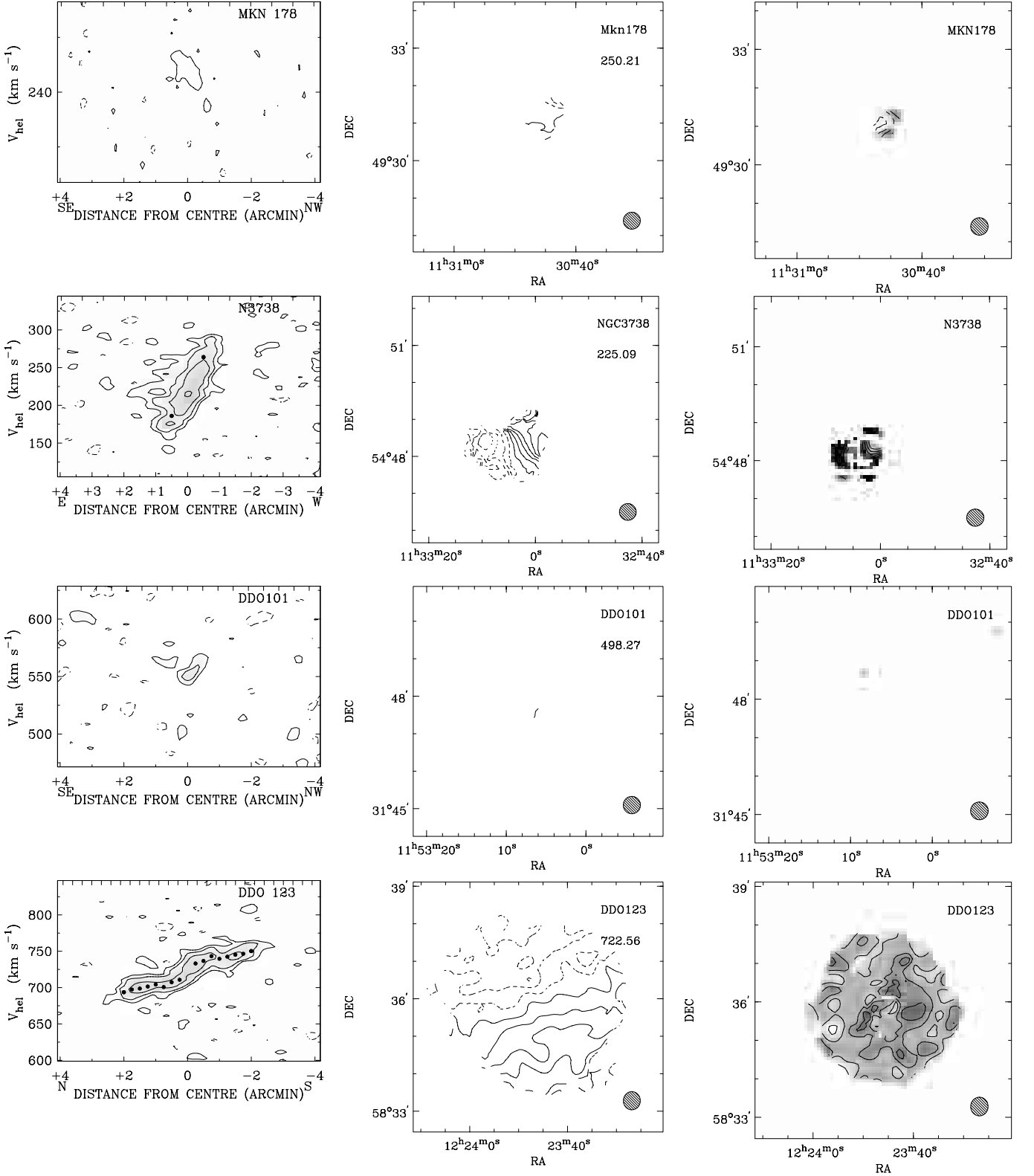


Fig. 4. continued.

Mk178: This galaxy is poorly resolved spatially. Its HI structure, kinematics and luminosity are similar to DDO 63, DDO 125 and DDO 165.

NGC 3738: The HI column density is high everywhere, with a central peak $N_{\text{HI}} = 4.5 \times 10^{21} \text{ cm}^{-2}$. The velocity

profile and the XV map indicate a steep velocity gradient, which is difficult to fit with the tilted ring method. The indicated rotation velocity was fitted by eye to the XV map. The high velocity dispersion is an artifact of the large velocity gradient in this marginally resolved galaxy.

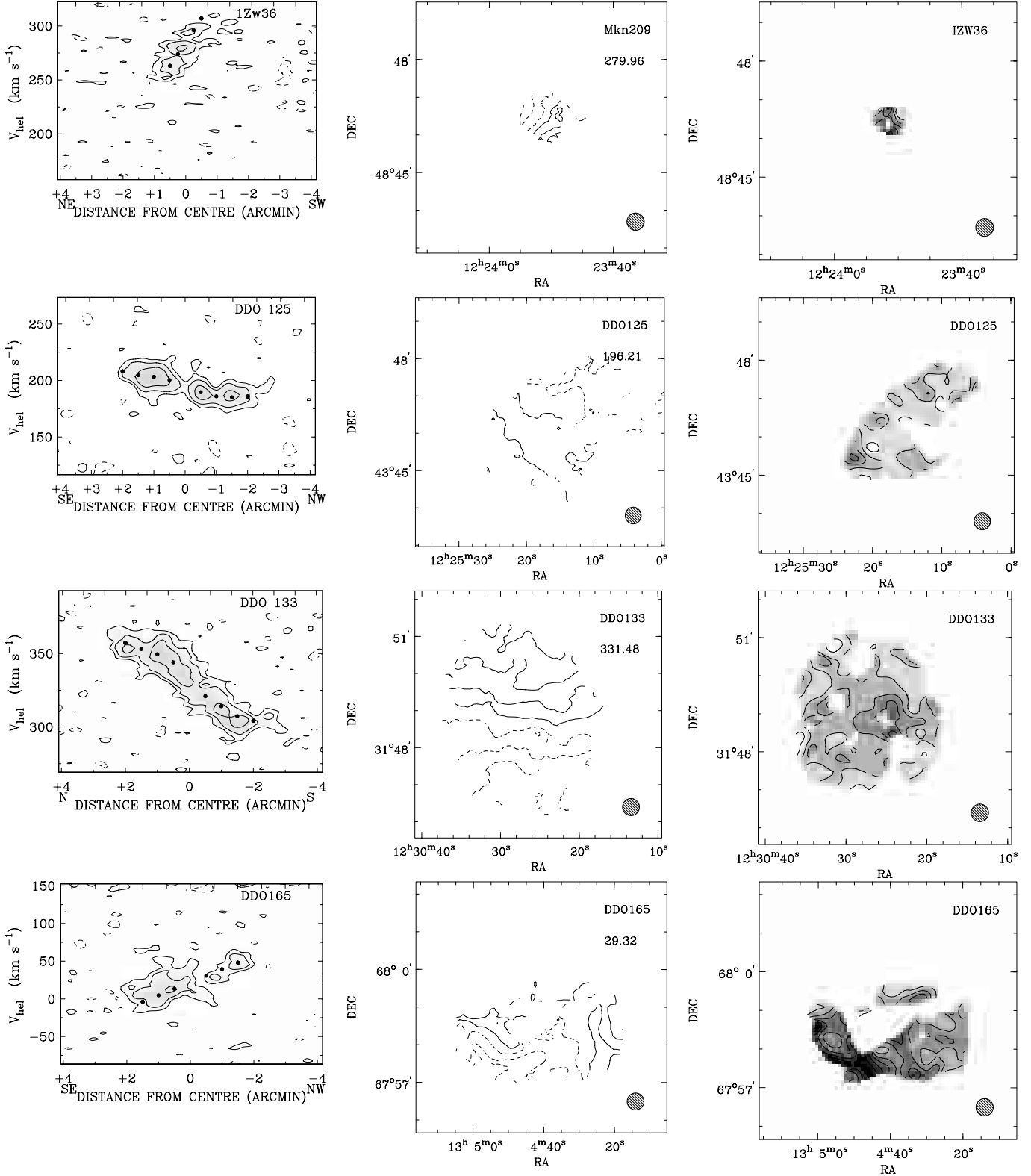


Fig. 4. continued.

DDO 101: HI extent is too limited to show structure.

DDO 123: The HI is distributed evenly throughout the disk of this face-on galaxy. Irregularities in the velocity field coincide with low signal-to-noise regions. The rotation curve rises out to the edge of the HI disk.

Mk 209: The peak of the high-column-density region is $N_{\text{HI}} = 2.9 \times 10^{21} \text{ cm}^{-2}$. The velocity field appears regular but could not be fitted properly because of insufficient resolution. The velocity gradient was fitted manually to the XV map; it is consistent with a solid body rotation

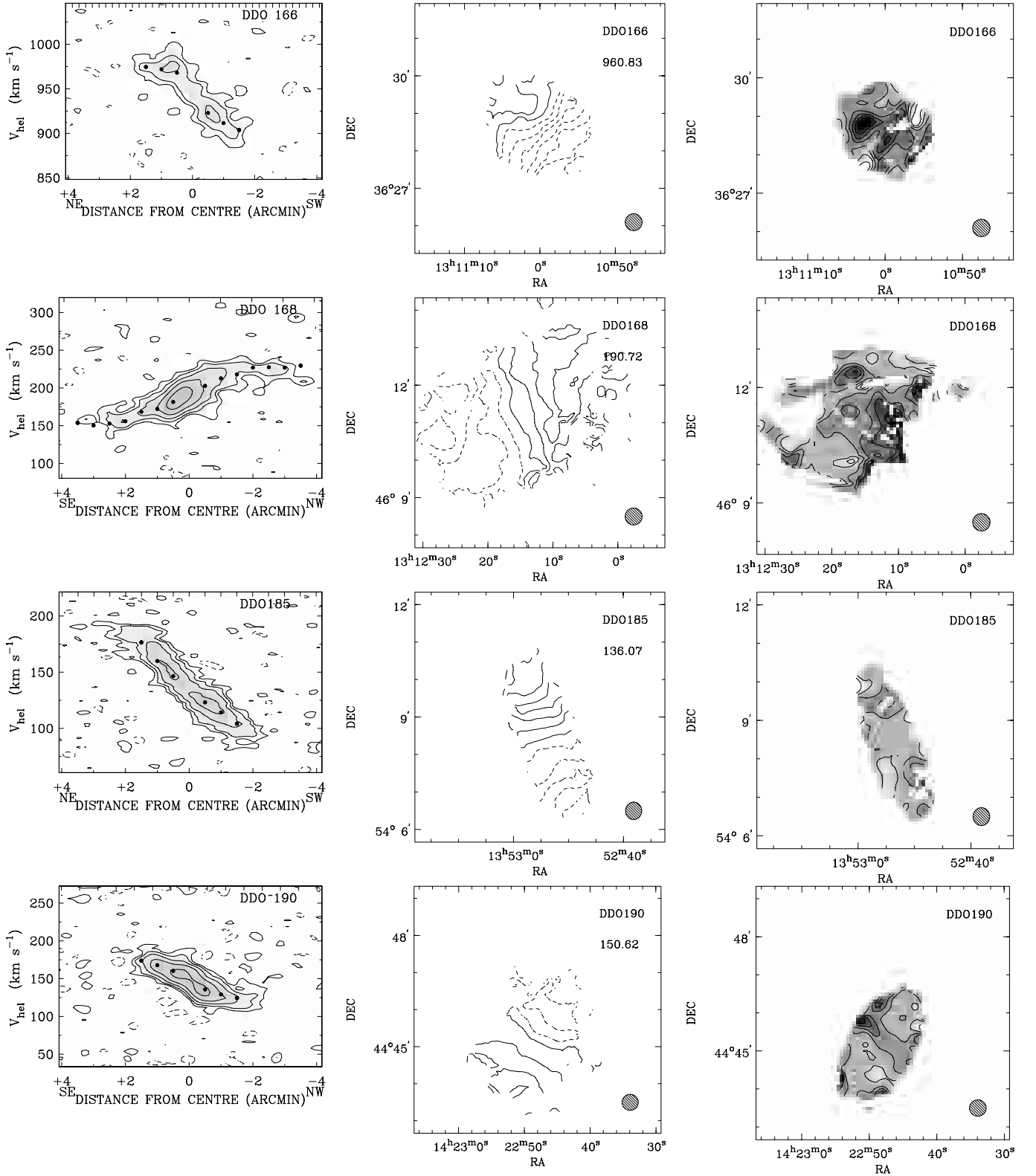


Fig. 4. continued.

curve. HI in Mk 209 (IZw 36) has also been observed with the VLA by Viallefond et al. (1987).

DDO 125: The HI is mainly concentrated in two high-column-density regions, separated by a low-column

density center, suggesting a fragmented ring. The velocity field shows a velocity gradient along the major axis. Rotation is clearly demonstrated by the XV map. HI in DDO 125 was studied in detail by Tully et al. (1978).

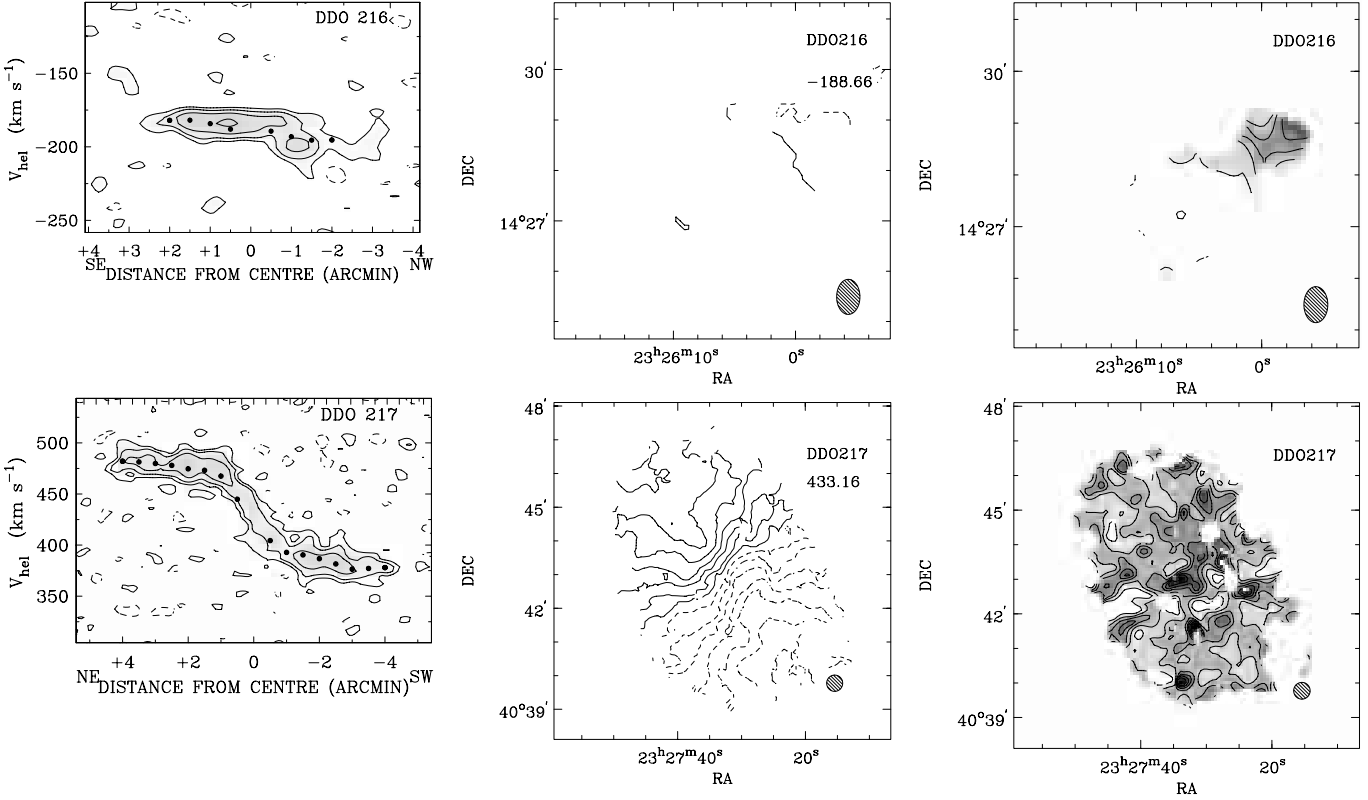


Fig. 4. continued.

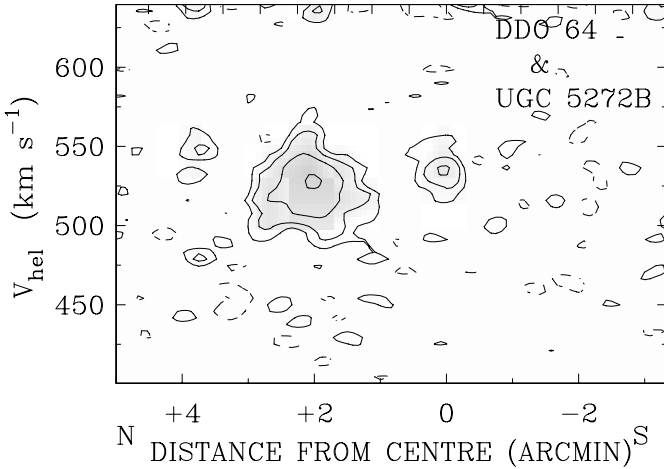


Fig. 5. Position-velocity map through UGC 5272B in position angle 21 degrees, with contour levels at -5.07 , 5.07 (2σ), 10.14 , 20.28 and 40.56 mJy per beam. UGC 5272B is visible at the center of the slice ($0'$); DDO 64 is the larger patch to the left (north) of UGC 5272B. The slice also intersects the feature which can be seen north of DDO 64 in the HI column density map. Not much is evident in the X-V map.

DDO 133: The HI isophotes are well-represented by ellipses at the level of $N_{\text{HI}} = 3 \times 10^{20}$ cm⁻². The rotation curve flattens slightly outwards.

DDO 165: Most of the HI is located in a ring. The velocity field is highly irregular. The maximum velocity gradient is along a line from SE to NW in position angle 120° – 140° .

The rotation velocity was fitted to the XV map assuming solid-body rotation. However, the emission in the XV map is mainly due to the ring.

DDO 166: A ridge of high-column-density HI extends over the eastern side of this face-on galaxy. The velocity field has a strong gradient, in spite of the small inclination suggested by the HI isophotes. The isovelocity contours are twisted into an S-shape at the eastern side of the galaxy. Thean et al. (1997) have published VLA HI maps of NGC 5033, DDO 166 and UGC 8314.

DDO 168: Two very high HI column density regions ($N_{\text{HI}} \geq 6 \times 10^{21}$ cm⁻²) occur near the center of the galaxy at $\alpha = 13^{\text{h}}12^{\text{m}}16^{\text{s}}.8$, $\delta = 46^\circ 11' 0''$ and $\alpha = 13^{\text{h}}12^{\text{m}}15^{\text{s}}.8$, $\delta = 46^\circ 11' 30''$. The position angle of the velocity gradient changes over the disk by approximately 20° . The peculiar structure of DDO 168 is not unique. Similar very-high column density regions combined with twisted velocity fields have been observed in other “amorphous galaxies” such as NGC 1140 (Hunter et al. 1994) and IZw 18 (Viallefond et al. 1987; Van Zee et al. 1998). A detailed study of the mass distribution in DDO 168 was performed by Broeils (1992).

DDO 185: The HI column density map suggests a disk seen at a high inclination. The isovelocity contours are regularly spaced, consistent with the rising rotation curve shown by the position-velocity map. DDO 185 (=Holmberg IV) was used as a calibrator galaxy for the Tully-Fisher relation by Kraan-Korteweg et al. (1988).

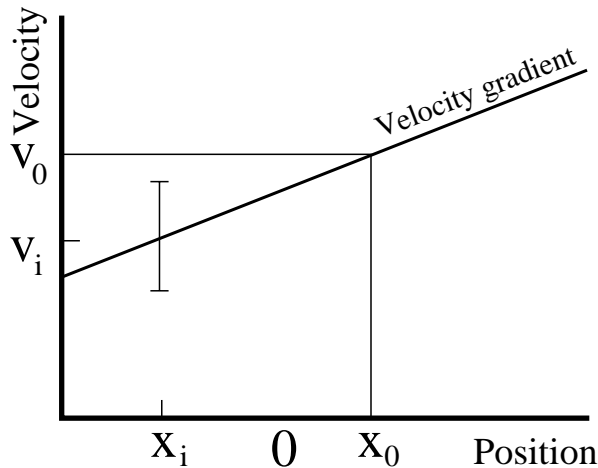


Fig. 6. Definition of symbols used in the calculation of the broadening of the line profile by a velocity gradient over the synthesized beam. The bar at the position x_i, v_i indicates the velocity dispersion σ_i at that position. The center of the beam is $x = 0$ by definition.

DDO 190: The highest column densities are found on the west. The velocity field is somewhat irregular, but rotation is clearly visible. The XV map shows a hint of flattening of the rotation curve on the NW side.

DDO 216: The HI is located in the southern half of the (optical) galaxy (Sandage 1986; Lo et al. 1993). There is little sign of rotation in the velocity field and XV maps. In fact, the velocity gradient suggested by the XV map may represent a single HI cloud at a discrepant velocity. For a VLA study of DDO 216, see Lo et al. (1993).

DDO 217: The HI is distributed relatively evenly over the disk. The velocity field shows differential rotation over most of the galaxy. The difference in position angle between the inner region and the outer disk is clearly visible in the velocity field map. The XV map shows a rapid rise of the rotation velocity near the center and a gradual increase in the outer disk.

Appendix A: Correction for the velocity gradient

We assume a well-resolved velocity field so that second and higher order derivatives of the velocity field can be neglected. We also assume an HI disk of negligible thickness so that all lines of sight cross the disk at a single radius, implying the one-dimensional linear situation sketched in Fig. 6. We choose the x -axis along the velocity gradient, with the zeropoint at the center of the Gaussian beam. The intensity-weighted mean velocity over the beam corresponds to that of position x_0 , which is not necessarily the center of the beam.

We now divide the beam into many (N) lines of sight, each with a large number of identical elements (M) with velocities v_{ik} , $k = 1 \dots M$ at position x_i , $i = 1 \dots N$. This definition includes an implicit integration over the

coordinate perpendicular to the velocity gradient. The elements are identified with individual HI clouds of very small intrinsic velocity dispersion. At every position x_i we define the mean velocity v_i as $v_i = \frac{1}{M} \sum_k v_{ik}$, which is related to the intensity-weighted mean velocity v_0 and the velocity gradient ∇v through

$$v_i - v_0 = (x_i - x_0) \nabla v.$$

The velocity dispersion of the elements at position x_i is $\sigma_i = \frac{1}{M-1} \sum_k (v_{ik} - v_i)^2$. Substitution of $v_i = v_0 + (x_i - x_0) \nabla v$ and evaluation of the cross-product yields

$$\sigma_i^2 = \frac{1}{M-1} \sum_{k=1}^M \{ (v_{ik} - v_0)^2 - 2(v_{ik} - v_0)(x_i - x_0) \nabla v + (x_i - x_0)^2 (\nabla v)^2 \}.$$

The third term is independent of k . For the second term we may write

$$\begin{aligned} -2(x_i - x_0) \nabla v \sum_{k=1}^M (v_{ik} - v_0) &= \\ -\frac{2M}{M-1} (v_i - v_0)(x_i - x_0) \nabla v &= \\ = -\frac{2M}{M-1} (x_i - x_0)^2 (\nabla v)^2. & \end{aligned}$$

Therefore, we have

$$\sigma_i^2 = \frac{1}{M-1} \sum_{k=1}^M (v_{ik} - v_0)^2 - \frac{M}{M-1} (x_i - x_0)^2 (\nabla v)^2.$$

The intensity-weighted mean velocity dispersion over the beam is

$$\langle \sigma^2 \rangle = \frac{\sum_{i=1}^N w_i \sigma_i^2}{\sum_{i=1}^N w_i}$$

with weight $w_i = e^{-x^2/b^2} I_i$. Therefore, with $M \gg 1$

$$\begin{aligned} \langle \sigma^2 \rangle &= \frac{\sum_{i=1}^N w_i \sum_{k=1}^M (v_{ik} - v_0)^2}{(M-1) \sum_{i=1}^N w_i} \\ &\quad - (\nabla v)^2 \frac{\sum_{i=1}^N w_i (x_i - x_0)^2}{\sum_{i=1}^N w_i}. \end{aligned}$$

This equation is of the general form

$$\langle \sigma^2 \rangle = \sigma_{\text{obs}}^2 - \Omega_I b^2 (\nabla v)^2,$$

where σ_{obs} is the observed dispersion of a local line profile, corrected for the instrumental spectral resolution. The second term is the line broadening due to the velocity gradient over the beam. Note that the velocity gradient ∇v is a function of position if the galaxy is not in solid-body rotation, necessitating use of a model velocity field constructed from the rotation curve in order to calculate ∇v at every position. The coefficient $\Omega_I > 0$ is a weighted

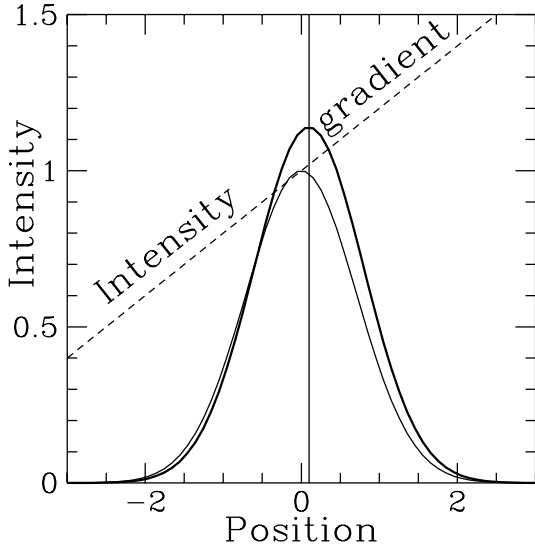


Fig. 7. The effect of an intensity gradient $I(x) = 1 + 0.2 \cdot x$ (dashed line; $a = 0.2$ in Table 6) over the beam (thin curve). The thick curve is the product of the intensity and the beamshape function with an arbitrary scaling. The vertical solid line marks the position $x_0 = 0.1$. The small difference between the thin and the thick curve (scaling in intensity is free) is the reason that the value of Ω_I is not sensitive to an intensity gradient over the synthesized beam. In this case, $\Omega_I = 0.490$ although $I(x=1)/I(x=-1) = 1.5$.

mean of the intensity distribution over the beam. If we assume a constant intensity (i.e. $I_i = I_0$) over the Gaussian beam $e^{-(x^2/b^2)}$, we have $\Omega_I = \frac{1}{2}$:

$$\langle \sigma^2 \rangle = \sigma_{\text{obs}}^2 - \frac{1}{2} b^2 (\nabla v)^2.$$

The error introduced by the assumption of a constant intensity can be estimated by calculating the correction for simple analytical intensity distributions. For any assumed intensity distribution

$$x_0 = \frac{\sum_{i=1}^N w_i x_i}{\sum_{i=1}^N w_i}$$

and $\langle \sigma^2 \rangle$ can be calculated. The coefficients Ω_I are given for three types of intensity distribution in Table 6.

Shallow intensity gradients do not make much of a difference. To first order, the effect of an intensity gradient is to shift the distribution of w_i in the direction of the intensity gradient. Since the beam function falls off rapidly for large x , only large gradients produce a significant difference with constant intensity. The greatest effect on the correction for the velocity gradient is brought about by the symmetric distribution. If the emission is highly concentrated towards the center of the beam, the velocity gradient has no effect. On the other hand, if the emission is concentrated in the wings of the beam, the effect of the velocity gradient is maximal. If the intensity does not change more than 50% over the beam, the error in Ω_I introduced by the assumption of constant intensity is of the order of 10%.

Table 6. Values of Ω_I for a gradient, a minimum and an jump in the intensity distribution at the center of the synthesized beam.

	Int. gradient	Int. minimum	Int. jump
a	$I(x) = 1 + ax$	$I(x) = 1 + ax^2$	$I(x) = 1 - aH(x)$
0.0	0.500	0.500	0.500
0.1	0.498	0.548	0.497
0.2	0.490	0.591	0.487
0.5	0.439	0.700	0.421
1.0	0.349	0.833	0.182

Note: scale factor a defines the magnitude of the intensity change over the beam. The position x is in units of beamsize b . At positions where the indicated functional forms are negative, the intensity was set to zero. The symbol $H(x)$ is used for the heaviside function $H(x) = \frac{x}{|x|}$.

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