

A sample of ultra steep spectrum sources selected from the Westerbork In the Southern Hemisphere (WISH) survey[★]

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Received 25 June 2002 / Accepted 1 August 2002

Abstract. The 352 MHz Westerbork In the Southern Hemisphere (WISH) survey is the southern extension of the WENSS, covering 1.60 sr between $-9^\circ < \delta < -26^\circ$ to a limiting flux density of ~ 18 mJy (5σ). Due to the very low elevation of the observations, the survey has a much lower resolution in declination than in right ascension ($54'' \times 54'' \cos \delta$). A correlation with the 1.4 GHz NVSS shows that the positional accuracy is less constrained in declination than in right ascension, but there is no significant systematic error. We present a source list containing 73570 sources. We correlate this WISH catalogue with the NVSS to construct a sample of faint Ultra Steep Spectrum (USS) sources, which is accessible for follow-up studies with large optical telescopes in the southern hemisphere. This sample is aimed at increasing the number of known high redshift radio galaxies to allow detailed follow-up studies of these massive galaxies and their environments in the early Universe.

Key words. surveys – radio continuum: general – radio continuum: galaxies – galaxies: active

1. Introduction

Powerful radio sources provide excellent targets to probe the formation and evolution of galaxies out to cosmological distances. The Hubble $K - z$ diagram of radio and near-IR selected galaxies shows that at $z \gtrsim 1$, the host galaxies of powerful radio sources are > 2 mag brighter than HDF field galaxies (De Breuck et al. 2002). Because there are strong arguments that this K -band emission is due to starlight, and not due to direct or scattered AGN contributions, high redshift radio galaxies (HzRGs) are among the most massive galaxies known at high redshift. This is consistent with the observations at low redshifts ($z \lesssim 1$), where radio galaxies are uniquely identified with massive ellipticals (e.g. Best et al. 1998; McLure & Dunlop 2000). Because HzRGs pinpoint over-dense regions in the early Universe, they have also been successfully used as tracers of proto-clusters at very high redshifts

(e.g. Le Fèvre et al. 1996; Pascarella et al. 1996; Pentericci et al. 2000; Venemans et al. 2002).

The first redshift surveys of radio sources targeted only the brightest objects in the sky (3CR Bennett 1962; Spinrad et al. 1985). Present-day surveys reach flux densities several orders of magnitude fainter than the 3CR (e.g. McCarthy et al. 1996; Lacy et al. 1999; Rawlings et al. 2001; Willott et al. 2002). However, because the optical spectroscopy of the host galaxies requires substantial integration times on 3–10 m class telescopes, these flux density limited surveys necessarily need to be limited in sky area. This excludes the objects with the lowest space density, such as the most distant luminous radio galaxies (e.g. Blundell et al. 1998; Jarvis et al. 2001b). To find such objects, additional “high redshift filters” need to be applied to the samples of radio surveys, at the expense of completeness. The most efficient filter is the selection of sources with ultra steep radio spectra (USS; $\alpha \lesssim -1$; $S \propto \nu^\alpha$). The success of this USS technique is mainly based on the k -correction of the generally concave radio spectrum of powerful radio galaxies. Several such USS samples have shown to be much more efficient in finding $z > 2$ radio galaxies than the complete samples (e.g. Röttgering et al. 1997; Stern et al. 1999; De Breuck et al. 2001; Jarvis et al. 2001a).

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* Table 2 is also available in electronic form at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsweb.u-strasbg.fr/cgi-bin/qcat?J/A+A/394/59>

★★ Marie Curie fellow.

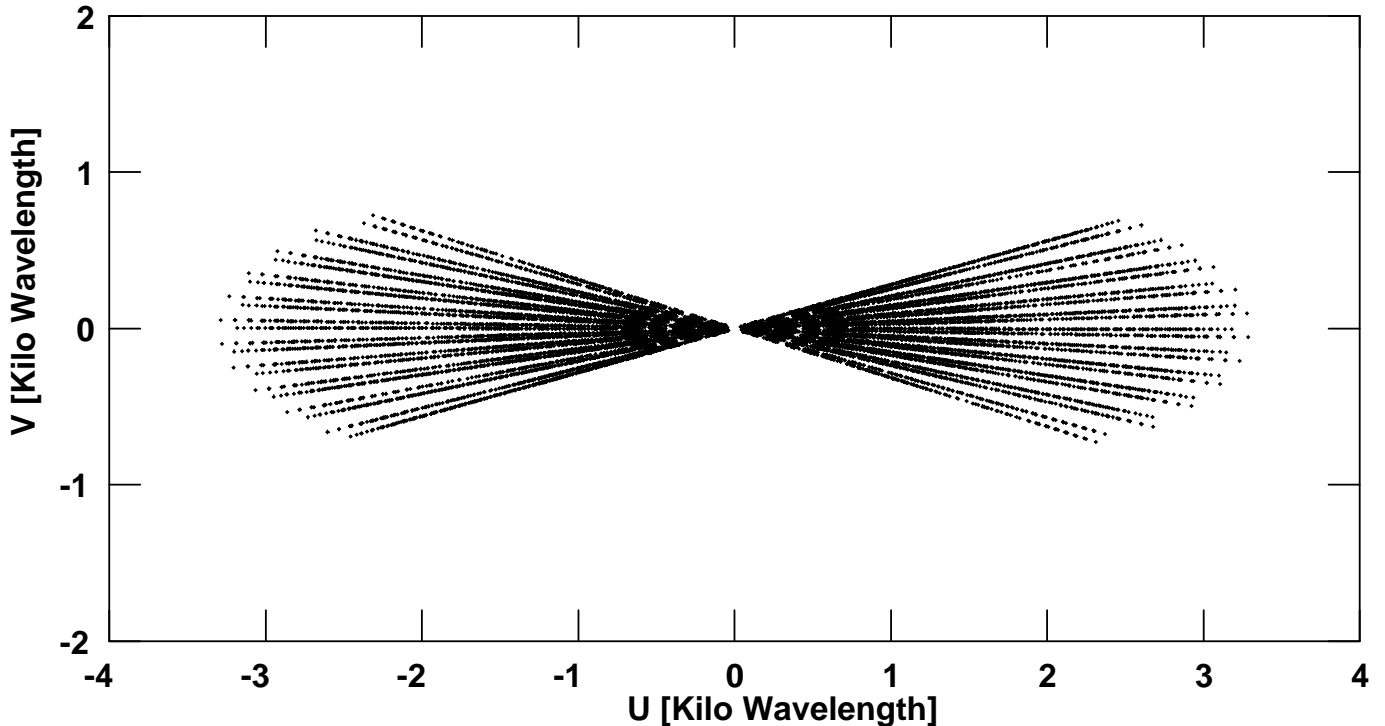


Fig. 1. Sample $u-v$ coverage plot of the field centered at $\alpha = 1^{\text{h}}28^{\text{m}}$, $\delta = -20^{\circ}$. The v -coverage is limited due to the low declination. Note the excellent radial coverage due to the bandwidth synthesis technique.

It is now possible to define large, well defined samples of USS sources using a new generation of large area radio surveys: the Westerbork Northern Sky Survey (WENSS; 325 MHz Rengelink et al. 1997), the Texas survey (Douglas et al. 1996, 365 MHz), The Sydney University Molonglo Sky Survey (SUMSS; 843 MHz Bock et al. 1999), the NRAO VLA Sky Survey (NVSS; 1.4 GHz Condon et al. 1998), and the Faint Images of the Radio Sky at Twenty centimeters (FIRST; 1.4 GHz Becker et al. 1995). De Breuck et al. (2000) have used these surveys to define a sample of 669 USS sources covering the entire sky outside the Galactic plane. However, their samples necessarily favour the northern hemisphere, because the WENSS survey, which is an order of magnitude deeper than the Texas survey, covers only the sky at $\delta > +29^{\circ}$. In this paper, we introduce the Westerbork In the Southern Hemisphere (WISH) survey, the southern extension of the WENSS. We use WISH in combination with NVSS to define a fainter sample of USS sample in the $-9^{\circ} < \delta < -26^{\circ}$ region, in analogy with the northern WENSS–NVSS sample of De Breuck et al. (2000). The construction of such a southern hemisphere sample is especially timely due to the advent of several 8 m class telescopes in the southern hemisphere, which can be used for the optical/near-IR identification and spectroscopy of the host galaxies.

The layout of this paper is as follows. In Sect. 2, we introduce the WISH survey, and compare the data products with the WENSS. In Sect. 3, we define the WISH–NVSS USS sample. Section 4 compares this new sample with previous samples and

Sect. 5 concludes with an overview of the planned observations of this sample.

2. The WISH survey

2.1. Motivation

The 325 MHz WENSS survey (Rengelink et al. 1997) has proved to be an extremely valuable tool for extra-galactic and galactic Astronomy. Some of the scientific applications include (i) the study of large-scale structure using radio sources (Rengelink 1999a), (ii) the search for high redshift radio galaxies (De Breuck et al. 2000), (iii) the selection of faint Gigahertz Peaked Spectrum sources (Snellen et al. 1998), (iv) the selection of Giant Radio Galaxies (Schoenmakers et al. 2001), (v) the construction of samples to search for gravitational lenses in the context of the Cosmic Lens All Sky Survey (Myers et al. 1995), (vi) the selection of optically bright galaxies with radio counterparts to construct the local radio luminosity functions of elliptical and spiral galaxies (de Ruiter et al. 1998), (vii) the study of linear polarization of the diffuse galactic radio background (Haverkorn et al. 2000), and (viii) the study of pulsars (Kouwenhoven 2000).

The WENSS survey is the most sensitive low-frequency survey covering a substantial fraction (25%) of the sky. The limitation of the sky coverage is due to the East-West construction of the WSRT. With the advent of several new large optical facilities in the southern hemisphere (VLT, Gemini South, Magellan), there is a clear need for equally sensitive samples in the South. Because the WSRT can also observe a part of the

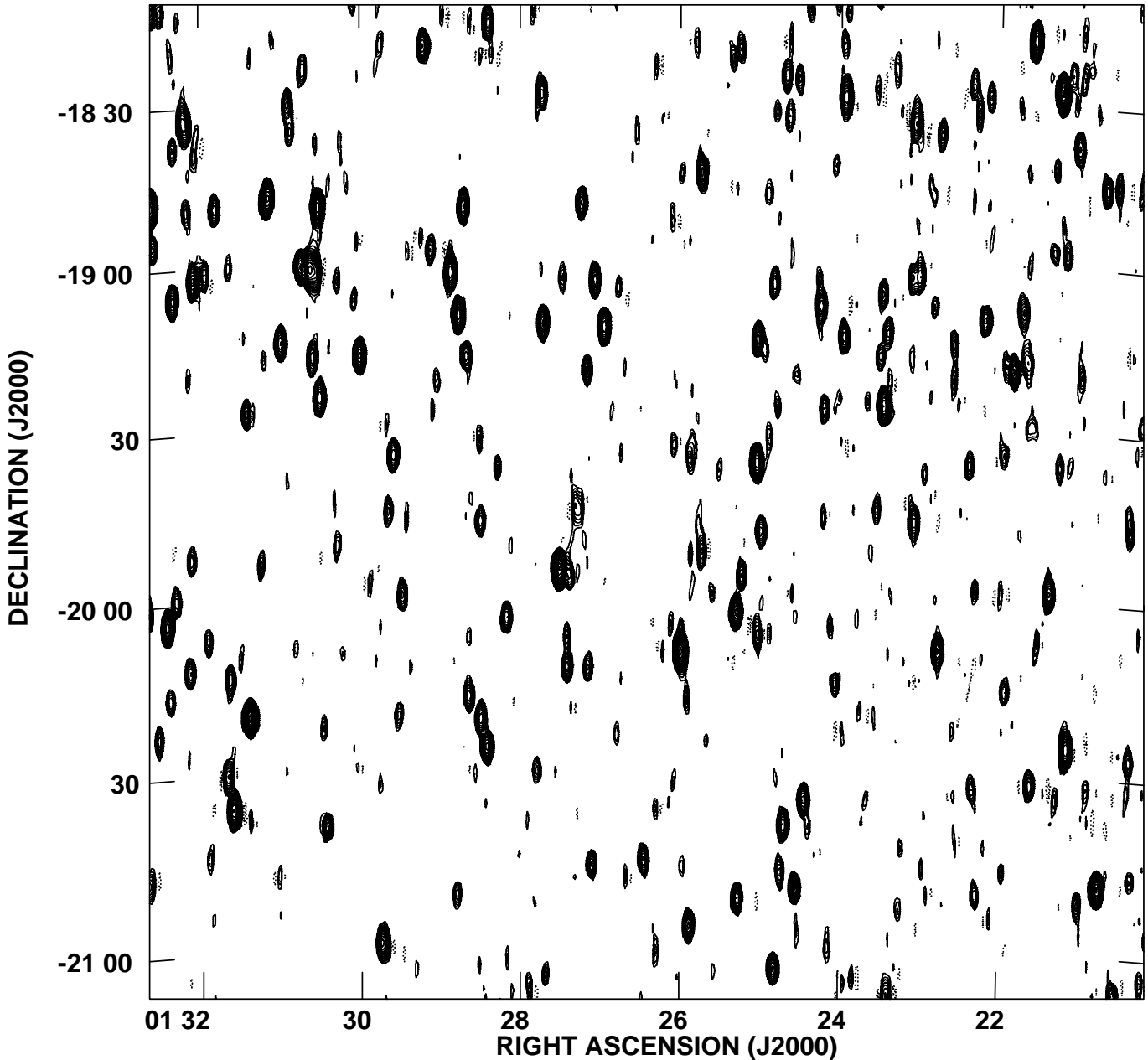


Fig. 2. Central part of the field centered at $\alpha = 1^{\text{h}}28^{\text{m}}$, $\delta = -20^{\circ}$. The contour scheme is a geometric progression in $\sqrt{2}$, which implies a factor 2 change in surface brightness every two contours (negative contours are dotted). The first contour is at 3σ with $\sigma = 2.0 \text{ mJy Beam}^{-1}$. Note that even with the elongated beam, we can still resolve several sources.

southern sky, we extended the WENSS by carrying out a similar Westerbork survey in the southern hemisphere, the WISH.

2.2. Survey design

The main goal of the WISH is to cover as large as possible an area that can be observed with the VLT. The limitations are the latitude of Westerbork, where the horizon is at $\delta = -37^{\circ}$, and an infinite elongation of the synthesized beam towards $\delta = 0^{\circ}$. We therefore imposed a northern limit of $\delta < -9^{\circ}$, to obtain a beam with a ratio of the major to minor axis of < 6 . In order to obtain at least 4 hours of hour angle coverage, we limited the survey to $\delta \gtrsim -26^{\circ}$. However, the uv -coverage is still

limited (see Fig. 1), which results in a synthesized beam with large near-in sidelobes. To limit the effects of this problem, we decided to avoid the Galactic Plane ($|b| > 10^{\circ}$) with its bright extended emission.

Apart from the above limitations, the design of WISH is based on that of WENSS. We used the broadband back-end with 8 bands of 5 MHz to improve the uv -coverage using bandwidth synthesis. The central frequencies of these 8 bands are 325.0, 333.0, 341.0, 347.0, 354.85, 366.6, 371.3, and 377.3 MHz. This bandwidth synthesis mosaicing technique was used before for the WENSS polar cap area ($\delta > +75^{\circ}$) and proved very successful (Rengelink 1999b). Because this requires only three configurations of the 9-A baseline, instead

of six for WENSS, WISH was also more efficient in observing time.

WISH consists of mosaics of 8×8 pointings each. With a grid-step of 1.25° , each mosaic covers $10^\circ \times 10^\circ = 100$ square degrees (10 degrees in Declination and 42 min in Right Ascension). Each pointing was observed for 20 s. With a move time of 10 s we have thus covered a mosaic once every 32 min. In an average observing time of 6 hours we have therefore obtained 11 cuts at each position. By observing in three different configurations (9A = 48 m, 72 m and 96 m) we have thus accumulated 33 cuts.

2.3. Observations and data reduction

The observations for WISH started in the Autumn 1997, and were concluded in the Spring of 1998. The observational techniques, data reduction, and source extraction process are identical to the polar cap region of the WENSS, and are discussed by Rengelink (1999b). Table 1 lists the 49 frames of the WISH with their respective field centers. Figure 1 shows the $u-v$ plane coverage of a sample field, and Fig. 2 the central $3^\circ \times 3^\circ$ of this field. Although the limited v -coverage leads to beam shapes elongated along the North-South direction, we can still resolve several sources into individual components.

The total sky area of WISH is 1.60 sr (half the WENSS area). Figure 3 shows the total sky coverage of WISH. Note that $\sim 38\%$ of the $\delta = -13^\circ$ fields could not be observed due to the limited observing time.

The final WISH catalogue contains 73 570 sources, and is available on the WENSS homepage (note that this catalogue has already been corrected for the flux density problem described in Sect. 2.4):

<http://www.strw.leidenuniv.nl/wenSS>

2.3.1. Noise

Figure 4 shows the distribution of the local noise level in WISH. The use of the broadband back-end for the WISH leads to slightly lower noise level compared with the main WENSS. However, we do not achieve the noise levels obtained in the WENSS polar cap region, which was observed with the same broadband system. This is probably due to the poorer $u-v$ coverage, resulting in higher sidelobe noise (see also Sect. 2.4).

2.3.2. Positional accuracy

To check the accuracy of the positions in WISH, we want to correlate the WISH with other catalogues with similar or better positional accuracy. Ideally, we would like to use accurate optical catalogues such as the USNO-A2.0, but due to the large WISH beam, there are on average more than one optical counterpart within one WISH beam. We therefore used other radio catalogues overlapping with the WISH area.

The only radio catalogues in the WISH area that have position accuracies $< 1''$ are the TEXAS (Douglas et al. 1996) and NVSS (Condon et al. 1998). Using the TEXAS, we find a mean position difference of $-0'.54$ (median $-0'.55$) in Right

Table 1. The 49 frames of the WISH. C is the flux density correction factor discussed in Sect. 2.4.3.

Frame	Mosaic center (B1950)		C
	RA	Dec	
SNH13_029	01 ^h 56 ^m 00 ^s	-13°00'	0.855
SNH13_089	05 ^h 56 ^m 00 ^s	-13°00'	0.809
SNH13_149	09 ^h 56 ^m 00 ^s	-13°00'	0.830
SNH13_179	11 ^h 56 ^m 00 ^s	-13°00'	0.835
SNH13_189	12 ^h 36 ^m 00 ^s	-13°00'	0.823
SNH13_199	13 ^h 16 ^m 00 ^s	-13°00'	0.807
SNH13_209	13 ^h 56 ^m 00 ^s	-13°00'	0.838
SNH13_219	14 ^h 36 ^m 00 ^s	-13°00'	0.860
SNH13_229	15 ^h 16 ^m 00 ^s	-13°00'	0.884
SNH13_239	15 ^h 56 ^m 00 ^s	-13°00'	0.846
SNH13_249	16 ^h 36 ^m 00 ^s	-13°00'	0.879
SNH13_259	17 ^h 16 ^m 00 ^s	-13°00'	0.831
SNH13_289	19 ^h 16 ^m 00 ^s	-13°00'	0.992
SNH13_299	19 ^h 56 ^m 00 ^s	-13°00'	0.903
SNH13_319	21 ^h 16 ^m 00 ^s	-13°00'	0.833
SNH13_329	21 ^h 56 ^m 00 ^s	-13°00'	0.822
SNH13_339	22 ^h 36 ^m 00 ^s	-13°00'	0.853
SNH13_349	23 ^h 16 ^m 00 ^s	-13°00'	0.858
SNH13_359	23 ^h 56 ^m 00 ^s	-13°00'	0.850
SNH20_000	00 ^h 00 ^m 00 ^s	-20°00'	0.818
SNH20_011	00 ^h 44 ^m 00 ^s	-20°00'	0.802
SNH20_021	01 ^h 24 ^m 00 ^s	-20°00'	0.816
SNH20_032	02 ^h 08 ^m 00 ^s	-20°00'	0.790
SNH20_042	02 ^h 48 ^m 00 ^s	-20°00'	0.755
SNH20_053	03 ^h 32 ^m 00 ^s	-20°00'	0.709
SNH20_063	04 ^h 12 ^m 00 ^s	-20°00'	0.691
SNH20_074	04 ^h 56 ^m 00 ^s	-20°00'	0.880
SNH20_084	05 ^h 36 ^m 00 ^s	-20°00'	0.855
SNH20_095	06 ^h 20 ^m 00 ^s	-20°00'	0.797
SNH20_127	08 ^h 28 ^m 00 ^s	-20°00'	0.811
SNH20_137	09 ^h 08 ^m 00 ^s	-20°00'	0.781
SNH20_148	09 ^h 52 ^m 00 ^s	-20°00'	0.803
SNH20_158	10 ^h 32 ^m 00 ^s	-20°00'	0.761
SNH20_169	11 ^h 16 ^m 00 ^s	-20°00'	0.716
SNH20_179	11 ^h 56 ^m 00 ^s	-20°00'	0.858
SNH20_190	12 ^h 40 ^m 00 ^s	-20°00'	0.870
SNH20_200	13 ^h 20 ^m 00 ^s	-20°00'	0.898
SNH20_211	14 ^h 04 ^m 00 ^s	-20°00'	1.039
SNH20_221	14 ^h 44 ^m 00 ^s	-20°00'	1.063
SNH20_232	15 ^h 28 ^m 00 ^s	-20°00'	1.128
SNH20_242	16 ^h 08 ^m 00 ^s	-20°00'	1.087
SNH20_253	16 ^h 52 ^m 00 ^s	-20°00'	1.122
SNH20_287	19 ^h 08 ^m 00 ^s	-20°00'	1.199
SNH20_297	19 ^h 48 ^m 00 ^s	-20°00'	0.974
SNH20_308	20 ^h 32 ^m 00 ^s	-20°00'	0.846
SNH20_318	21 ^h 12 ^m 00 ^s	-20°00'	0.937
SNH20_329	21 ^h 56 ^m 00 ^s	-20°00'	0.785
SNH20_339	22 ^h 36 ^m 00 ^s	-20°00'	0.904
SNH20_350	23 ^h 20 ^m 00 ^s	-20°00'	0.846

Ascension (RA) and $-0'.27$ (median $-0'.39$) in Declination (DEC). The NVSS has better position accuracy than the TEXAS, and we have therefore used NVSS sources as position calibrators for WISH. A comparison of the WISH and NVSS positions is therefore not independent, and should be treated

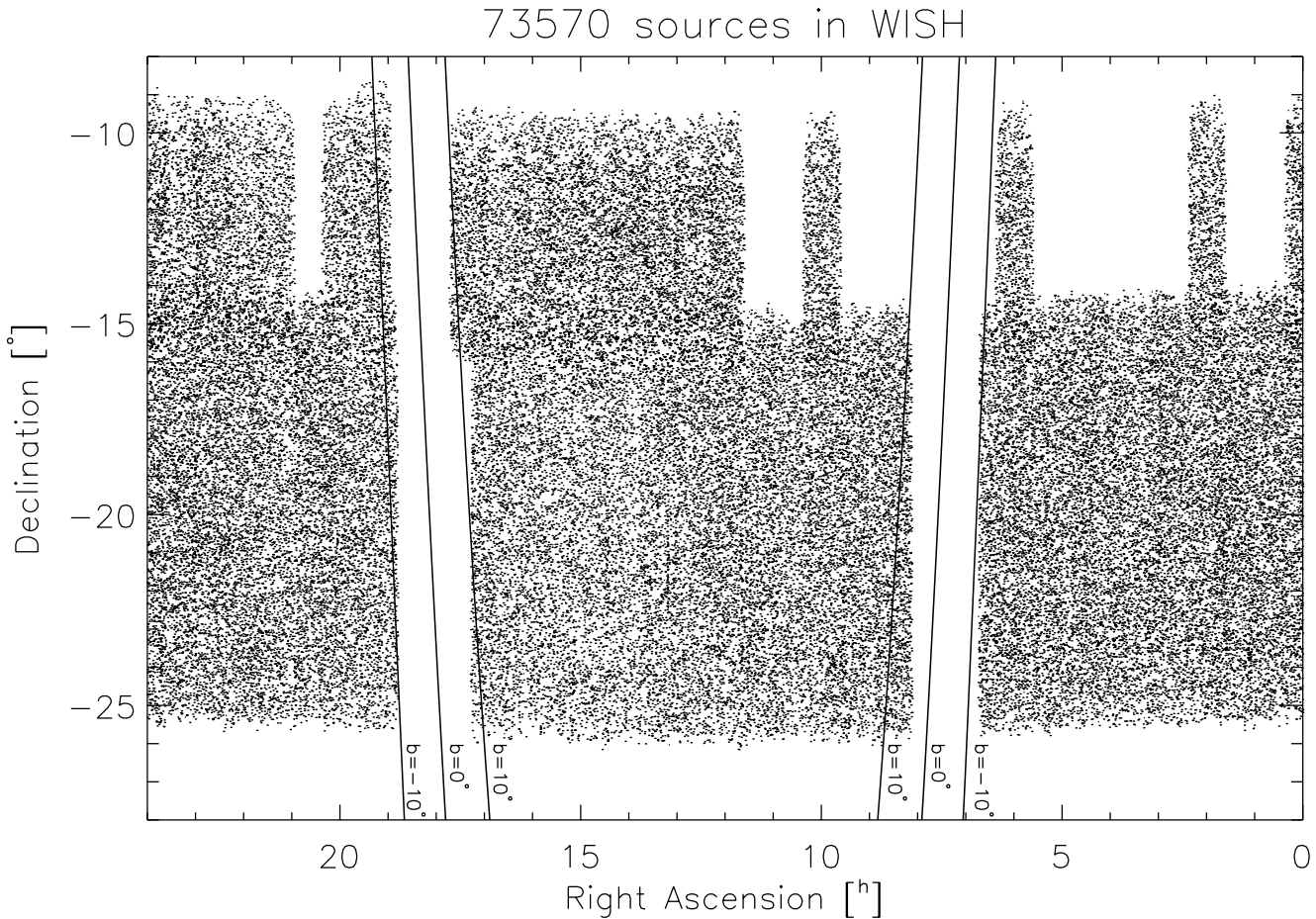


Fig. 3. Location of the 73 570 sources detected in WISH. The galactic plane ($|b| < 10^\circ$) was not observed.

with caution. We still perform such a comparison, as our sample of USS sources (Sect. 3) is based on this correlation.

To avoid problems due to the different resolution of WISH and NVSS, we only retained WISH sources fitted with a single Gaussian component (see Sect. 2.3.3), and excluded all objects with ≥ 2 NVSS sources within $200''$. Figure 5 shows the relative position differences. The most obvious feature is the strong asymmetry in the position differences due to the very elongated synthesized beam of the WISH. This effect is not seen in the comparison with the TEXAS survey, as this survey groups sources $< 2'$ as a single entry in the catalogue.

The mean offset between WISH and NVSS is $\langle \Delta(\text{RA}) \rangle = -0'.59$ (median $-0'.51$) and $\langle \Delta(\text{Dec}) \rangle = -0'.32$ (median $-0'.37$). These values are consistent with those found from the correlation with the TEXAS. The offset can also be compared with the systematic offset between the WENSS and NVSS positions: $\langle \Delta(\text{RA}) \rangle = -0.12$ and $\langle \Delta(\text{Dec}) \rangle = -0.09$. However, for consistency with the WENSS, and in order to retain independent positions for comparison with future surveys, we did not apply this correction to the WISH catalogue.

We find a significantly larger offset in RA, which is due to the lack of correction for polar motion and/or a timing problem (see also de Vries et al. 2002). Despite this large scatter due to the elongated beam, the offset in DEC is only half that in RA,

indicating that the large Gaussian fitting errors do not introduce a systematic bias.

2.3.3. Morphology

As can be seen from Fig. 2, WISH has poor resolution in declination. The mean ratio between the fitted major and minor axis $\langle b_{\text{maj}}/b_{\text{min}} \rangle = 5.00$, while for WENSS, $\langle b_{\text{maj}}/b_{\text{min}} \rangle = 1.89$. This illustrates the strong difference in ellipticity of the synthesised beam. Because we used the same source finding algorithm as WENSS, this leads to a large number of mis-classifications of WISH sources as resolved North-South sources. Because the NVSS already covers the same area at $45''$ resolution, the NVSS morphological parameters should be used in any selection based on such parameters. We include the WISH source dimensions only for consistency with the 2 WENSS catalogues.

2.4. Flux density

2.4.1. Comparison with other radio surveys

To check the flux calibration of the WISH, we correlated WISH with several other radio surveys. Because no previous 352 MHz

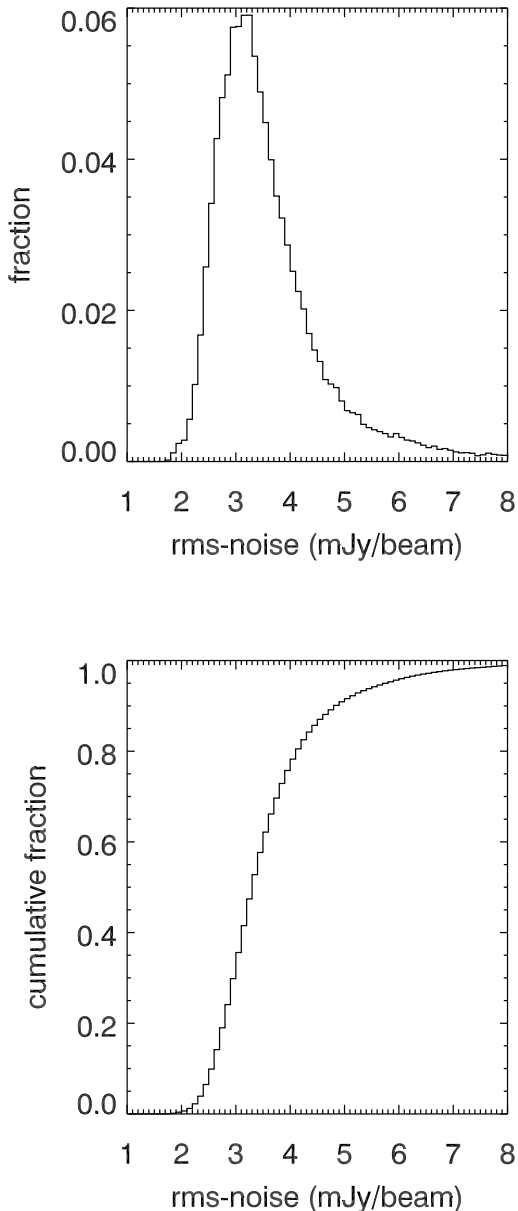


Fig. 4. The differential (top) and cumulative distribution (bottom) of rms-noise in the WISH.

survey is available covering this part of the sky, we have to use surveys at different frequencies. These are the 365 MHz Texas survey (Douglas et al. 1996), the 408 MHz MRC Large et al. (1981), and the 1.4 GHz NVSS (Condon et al. 1998). Both the Texas and MRC surveys use the TXS flux scale, which is related to the more commonly used Baars (BA) scale by the relation $S(\text{TXS}) = 0.9607 \times S(\text{BA})$ (Douglas et al. 1996). To compare the flux densities, we have put the Texas and MRC flux densities on the Baars scale, as used for WISH.

For the Texas survey, the difference in survey frequency should cause the WISH flux densities to be only $\sim 3\%$ brighter than the Texas flux densities (assuming $\alpha_{352}^{365} = -0.8$). The observed flux ratio is clearly higher: $\langle S_{\text{WISH}}/S_{\text{Texas}} = 1.071 \rangle$ (median 1.059). A similar result is seen for the MRC: $\langle S_{\text{WISH}}/S_{\text{MRC}} = 1.242 \rangle$ (median 1.262) with an expected value of 1.125. For the NVSS, the flux differences are larger,

but they also point towards a $\sim 16\%$ over-estimation of the WISH flux densities: $\langle S_{\text{WISH}}/S_{\text{NVSS}} = 3.500 \rangle$ (median 3.551) with an expected value of 3.018.

2.4.2. Cause of the flux density problem

We examined the dependence of this flux discrepancy on several source parameters such as flux density, morphology, and position difference with respect to the other surveys. None of these appear to influence this flux discrepancy.

We do find a strong variation of mean spectral index between individual survey frames (Table 1). There is no dependence on Galactic latitude or longitude, or on Galactic extinction, nor is there a dependence on distance from the field centre. There is a clear discontinuity across the frame boundaries, but no strong inter-frame correlations, suggesting the flux discrepancy depends rather on the epoch of observation than on the position on the sky.

The most obvious explication for these large flux density errors are the low elevations at which the WISH has been observed. This could lead to errors in the system temperature corrections, especially when using flux density calibrators observed at much higher elevations.

2.4.3. Correction using NVSS

Because none of the other radio surveys have the same frequency as WISH, we can only apply a statistical correction factor to the individual frames, assuming a constant spectral index between the two survey frequencies. Although they are close in survey frequency, we prefer not to use the Texas and MRC surveys, because they are much shallower than the WISH, and have $2 - 3\times$ poorer resolution. We shall therefore only use the NVSS in the following.

To determine the expected WISH–NVSS spectral indices, we use the spectral index distribution of the 13 600 sources in the WENSS–NVSS polar cap region. Here, we consider only the sources from the WENSS polar cap, because they were observed with the same broadband receiver system as WISH¹. As noted by Rengelink (1999b), this system leads to bandwidth smearing, which attenuates faint sources, leading to an underestimate of the flux densities of the order of 10% for the faintest ($S_{352} \lesssim 30$ mJy) sources. To avoid this problem, we therefore consider only sources with $S_{352} > 40$ mJy. As shown by De Breuck et al. (2000), this does not significantly affect the spectral index distribution. Because the steep part of the WENSS–NVSS spectral index distribution has a nearly Gaussian distribution, we use the fitted peak $GP_{\text{WENSS-NVSS}} = -0.786$ of the distribution to compare with the WISH–NVSS spectral indices. In each WISH frame, there are on average ~ 550 single component sources with $S_{352} > 40$ mJy and an NVSS counterpart. We also fit a Gaussian to their spectral index distribution to determine $GP_{\text{WISH-NVSS}}$ for each frame.

¹ We have checked that the WENSS polar cap region is not subject to this flux density problem; this implies that the problem is not due to the broadband receiver system itself.

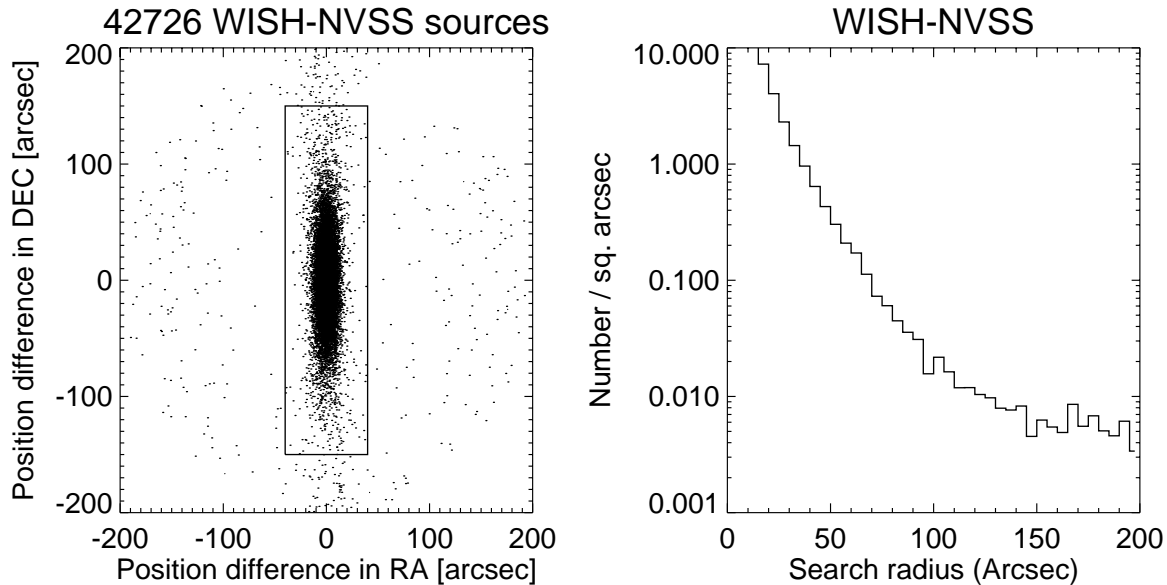


Fig. 5. *Left:* plot showing the relative position differences between WISH and NVSS. Only single component WISH sources (Sect. 2.3.3) with a single NVSS counterpart within 200'' are considered. Note the strong asymmetry of this distribution due to the very elongated WISH synthesized beam. The box indicates the region where we accept only 1 NVSS source as being associated with a single WISH source (Sect. 4). *Right:* the density of NVSS sources around a WISH source. Note that the distribution flattens from 150'' onwards.

Forcing these peaks to coincide with $GP_{\text{WENSS-NVSS}}$ yields correction factors

$$C = (1400/352)^{GP_{\text{WISH-NVSS}} - GP_{\text{WENSS-NVSS}}}$$

Table 1 lists these correction factors for each frame. They range from 0.691 to 1.199, with a mean of 0.868 (median 0.846), indicating the WISH flux densities are statistically overestimated by an average $\sim 13\%$.

We have also determined the correction factors comparing the mean and median spectral WISH–NVSS and WENSS–NVSS spectral indices. The resulting correction factors are consistent with the ones determined from the Gaussian peaks, with a mean difference of $\sim 1\%$, and a maximum difference of 7% and 2% for the correction factors based on the mean and median, respectively. We therefore estimate these statistical correction factors to be accurate to $< 2\%$.

We have applied these correction factors to flux densities listed in the WISH catalogue published on the WENSS homepage, and will consider only the corrected values in the remainder of this paper.

3. USS sample selection

We now use the WISH to define a sample of USS sources analogous to the WENSS–NVSS sample of De Breuck et al. (2000). We use the same selection criteria, i.e. only sources outside the Galactic Plane $|b| > 10^\circ$ with $\alpha_{352}^{1400} < -1.30$, $S_{1400} \geq 10$ mJy, and position differences between WISH and NVSS $< 10''$. We also need to exclude objects which are listed as a single source in WISH, but resolved in the NVSS. This can occur due to differences in resolution and source finding algorithms, and could easily introduce spurious USS sources. In the WENSS–NVSS, De Breuck et al. (2000) conservatively

excluded all WENSS sources with > 1 NVSS source within a 72'' from the WENSS position. Because the WISH synthesized beam is much more asymmetrical than the WENSS beam, we cannot use such a circular beam here. Based in Fig. 5, we adopt a box of 40'' in RA and 150'' in DEC to exclude the multiple-component sources. This results in a sample of 154 USS sources, listed in Table 2.

4. Discussion

Although we used almost the same selection criteria, the source density of the WISH–NVSS USS sample (78 sr^{-1}) is roughly half that of the WENSS–NVSS USS sample. This can be explained by the much larger position uncertainty in declination of the WISH survey (Sect. 2.3.2), which results in a much larger number of sources that fall outside our adopted circular 10'' search radius. Contrary to the asymmetrical exclusion box for multiple NVSS sources (Sect. 3), we keep a circular correlation search radius because the sources with larger position offsets could also have larger flux density uncertainties, and could also select sources with large angular sizes ($\geq 1'$), which are less likely to be at very high redshifts. This implicitly introduces a small bias against sources oriented North-South.

We have searched the literature for previous observations of the sources in this USS sample. There is a significant overlap (16 sources) with the shallower Texas–NVSS USS sample of De Breuck et al. (2000). This also includes the highest redshift radio galaxy known to date, TN J0924–2201 ($z = 5.19$ van Breugel et al. 1999). Six sources are also detected in the second data release of the 2MASS survey Jarrett et al. (2000), suggesting they are identified with low redshift galaxy clusters. This is consistent with the samples of De Breuck et al. (2000),

Table 2. The WISH–NVSS USS sample.

Name	S_{352} mJy	S_{1400} mJy	α_{352}^{1400}	α_{J2000} h m s	δ_{J2000} ° ' "	z	ID ^a	Reference
WN J0015–1528	104 ± 6	15.5 ± 0.7	−1.38 ± 0.06	00 15 59.65	−15 28 20.3			
WN J0017–2120	85 ± 5	13.7 ± 0.6	−1.32 ± 0.05	00 17 48.34	−21 20 54.9			
WN J0036–1835	101 ± 6	16.6 ± 1.0	−1.30 ± 0.06	00 36 54.52	−18 35 48.7			
WN J0037–1904	422 ± 18	64.9 ± 2.4	−1.36 ± 0.04	00 37 23.76	−19 04 32.2			
WN J0038–1540	406 ± 17	62.7 ± 1.9	−1.35 ± 0.04	00 38 47.15	−15 40 06.1			De Breuck et al. (2000)
WN J0138–1420	82 ± 6	10.7 ± 0.6	−1.47 ± 0.07	01 38 30.25	−14 20 07.6			
WN J0141–1406	296 ± 13	29.3 ± 1.3	−1.68 ± 0.05	01 41 38.02	−14 06 08.8			
WN J0151–1417	180 ± 9	18.4 ± 1.0	−1.65 ± 0.05	01 51 38.85	−14 17 44.7			
WN J0203–1315	215 ± 10	35.4 ± 1.2	−1.31 ± 0.04	02 03 28.09	−13 15 07.1			
WN J0220–2311	366 ± 16	58.1 ± 1.8	−1.33 ± 0.04	02 20 24.08	−23 11 39.6			
WN J0223–1539	104 ± 6	12.1 ± 0.6	−1.56 ± 0.05	02 23 31.53	−15 39 52.5			
WN J0224–1701	65 ± 6	10.0 ± 0.6	−1.36 ± 0.08	02 24 16.23	−17 01 01.3			
WN J0230–2001	413 ± 18	55.6 ± 2.1	−1.45 ± 0.04	02 30 45.73	−20 01 18.9			
WN J0230–1706	64 ± 6	10.5 ± 0.6	−1.31 ± 0.08	02 30 52.16	−17 06 26.8			
WN J0246–1649	368 ± 15	54.9 ± 1.7	−1.38 ± 0.04	02 46 52.89	−16 49 28.1			
WN J0306–1736	125 ± 6	19.8 ± 0.8	−1.34 ± 0.05	03 06 29.62	−17 36 38.6			
WN J0314–1849	143 ± 8	21.2 ± 0.8	−1.38 ± 0.05	03 14 13.66	−18 49 50.7			
WN J0316–2137	100 ± 6	15.6 ± 0.7	−1.34 ± 0.06	03 16 21.01	−21 37 41.3			
WN J0330–1810	96 ± 6	14.9 ± 0.6	−1.35 ± 0.05	03 30 50.47	−18 10 59.6			
WN J0335–2041	364 ± 16	53.5 ± 2.0	−1.39 ± 0.04	03 35 40.08	−20 41 10.4			
WN J0340–2159	313 ± 14	51.6 ± 2.0	−1.31 ± 0.04	03 40 14.93	−21 59 54.4			
WN J0341–1719	191 ± 9	27.6 ± 1.3	−1.40 ± 0.05	03 41 55.52	−17 19 27.0			
WN J0349–1801	94 ± 6	15.6 ± 0.7	−1.30 ± 0.06	03 49 25.11	−18 01 31.7			
WN J0414–2114	307 ± 13	47.0 ± 1.5	−1.36 ± 0.04	04 14 01.22	−21 14 53.9			
WN J0423–1537	217 ± 9	31.8 ± 1.7	−1.39 ± 0.05	04 23 32.85	−15 37 43.1			
WN J0456–2202	66 ± 5	10.9 ± 0.6	−1.31 ± 0.07	04 56 49.41	−22 02 55.5			
WN J0510–1838	6217 ± 254	634.0 ± 20.7	−1.65 ± 0.04	05 10 32.43	−18 38 42.5		2MASS	Jarrett et al. (2000)
WN J0526–1830	134 ± 6	19.6 ± 1.3	−1.39 ± 0.06	05 26 24.60	−18 30 40.1			
WN J0526–2145	112 ± 6	17.6 ± 0.7	−1.34 ± 0.05	05 26 48.80	−21 45 19.8			
WN J0528–1710	131 ± 7	21.8 ± 0.8	−1.30 ± 0.05	05 28 04.93	−17 10 04.1			
WN J0536–1357	279 ± 14	40.1 ± 1.6	−1.41 ± 0.05	05 36 51.96	−13 57 10.3		2MASS	Jarrett et al. (2000)
WN J0557–1124	351 ± 15	40.1 ± 1.9	−1.57 ± 0.05	05 57 00.72	−11 24 16.6			
WN J0602–2036	235 ± 11	37.7 ± 1.2	−1.33 ± 0.04	06 02 29.41	−20 36 46.3			
WN J0604–2015	222 ± 10	27.6 ± 1.0	−1.51 ± 0.04	06 04 57.58	−20 15 56.9			
WN J0621–1902	64 ± 5	10.3 ± 1.0	−1.32 ± 0.09	06 21 05.78	−19 02 03.6			
WN J0851–1728	101 ± 6	15.1 ± 0.7	−1.38 ± 0.05	08 51 25.09	−17 28 43.8			
WN J0903–1759	243 ± 11	38.2 ± 1.2	−1.34 ± 0.04	09 03 44.11	−17 59 52.5			
WN J0910–2228	469 ± 20	53.5 ± 1.7	−1.57 ± 0.04	09 10 34.15	−22 28 43.3		R	De Breuck et al. (2000)
WN J0912–1655	104 ± 7	11.8 ± 0.6	−1.58 ± 0.06	09 12 57.24	−16 55 54.8			
WN J0924–2201	454 ± 19	71.1 ± 2.2	−1.34 ± 0.04	09 24 19.94	−22 01 42.2	5.19	R, K	van Breugel et al. (1999)
WN J0938–1831	116 ± 7	18.4 ± 1.2	−1.33 ± 0.06	09 38 44.24	−18 31 36.2			
WN J0945–2118	160 ± 8	18.3 ± 0.7	−1.57 ± 0.04	09 45 29.43	−21 18 48.8		DSS	
WN J0949–2520	103 ± 8	15.9 ± 0.7	−1.35 ± 0.07	09 49 28.04	−25 20 13.4			
WN J0956–1500	440 ± 19	67.9 ± 2.1	−1.35 ± 0.04	09 56 01.46	−15 00 21.5			De Breuck et al. (2000)
WN J1001–1723	70 ± 5	10.8 ± 0.6	−1.36 ± 0.07	10 01 15.75	−17 23 27.2			
WN J1014–1830	86 ± 6	11.6 ± 0.6	−1.45 ± 0.06	10 14 10.48	−18 30 00.9		DSS	
WN J1026–2116	444 ± 19	60.2 ± 1.9	−1.45 ± 0.04	10 26 22.35	−21 16 07.8		R, K	De Breuck et al. (2000)
WN J1036–1837	446 ± 19	70.8 ± 2.2	−1.33 ± 0.04	10 36 16.19	−18 37 28.0			
WN J1047–1836	153 ± 8	20.4 ± 0.8	−1.46 ± 0.05	10 47 15.51	−18 36 31.1			
WN J1051–1517	103 ± 8	16.6 ± 0.7	−1.32 ± 0.06	10 51 45.62	−15 17 06.3			

Table 2. continued.

Name	S_{352} mJy	S_{1400} mJy	α_{352}^{1400}	α_{J2000} h m s	δ_{J2000} ° ' "	z	ID	Reference
WN J1052–1812	118 ± 7	14.4 ± 0.6	-1.52 ± 0.05	10 52 00.82	-18 12 32.3			
WN J1053–1656	226 ± 10	35.8 ± 1.5	-1.34 ± 0.05	10 53 19.77	-16 56 40.2			
WN J1103–2004	189 ± 9	30.7 ± 1.0	-1.32 ± 0.04	11 03 17.51	-20 04 58.9			
WN J1104–1913	203 ± 10	29.3 ± 1.4	-1.40 ± 0.05	11 04 07.25	-19 13 47.6		DSS	
WN J1109–1917	1324 ± 54	197.4 ± 5.9	-1.38 ± 0.04	11 09 49.93	-19 17 53.7			
WN J1123–2154	299 ± 13	49.3 ± 1.6	-1.30 ± 0.04	11 23 10.11	-21 54 05.6			
WN J1127–2126	81 ± 6	11.8 ± 0.6	-1.40 ± 0.06	11 27 54.80	-21 26 21.8		2MASS	Jarrett et al. (2000)
WN J1132–2102	219 ± 10	30.8 ± 1.0	-1.42 ± 0.04	11 32 52.66	-21 02 45.0			
WN J1138–1324	83 ± 8	10.0 ± 0.6	-1.53 ± 0.08	11 38 05.42	-13 24 23.5			
WN J1143–2143	423 ± 18	67.6 ± 2.5	-1.33 ± 0.04	11 43 17.43	-21 43 31.2			
WN J1148–2114	101 ± 6	15.5 ± 0.7	-1.36 ± 0.05	11 48 13.50	-21 14 03.7			
WN J1150–1317	205 ± 9	31.1 ± 1.0	-1.37 ± 0.04	11 50 09.59	-13 17 53.9			
WN J1151–2547	86 ± 9	11.8 ± 1.0	-1.44 ± 0.10	11 51 46.42	-25 47 52.2		DSS	
WN J1152–1558	116 ± 6	14.1 ± 1.5	-1.52 ± 0.09	11 52 50.03	-15 58 00.9			
WN J1200–1125	74 ± 5	10.2 ± 0.6	-1.44 ± 0.06	12 00 54.27	-11 25 48.6			
WN J1211–1126	91 ± 5	14.0 ± 1.3	-1.36 ± 0.08	12 11 32.05	-11 26 01.5			
WN J1222–2129	101 ± 7	14.3 ± 0.6	-1.42 ± 0.06	12 22 48.22	-21 29 10.0			
WN J1222–2531	99 ± 13	11.6 ± 0.6	-1.55 ± 0.10	12 22 50.24	-25 31 18.0		2MASS	Abell et al. (1989)
WN J1224–2101	103 ± 6	16.9 ± 0.7	-1.31 ± 0.05	12 24 30.26	-21 01 41.9			
WN J1225–1429	245 ± 11	31.2 ± 1.0	-1.49 ± 0.04	12 25 31.71	-14 29 15.4			
WN J1230–1038	65 ± 6	10.3 ± 0.6	-1.33 ± 0.08	12 30 00.61	-10 38 58.5			
WN J1242–1616	196 ± 9	30.9 ± 1.0	-1.34 ± 0.04	12 42 13.34	-16 16 57.4			
WN J1242–1646	119 ± 7	19.7 ± 0.8	-1.30 ± 0.05	12 42 27.56	-16 46 36.4			
WN J1246–1245	166 ± 8	27.4 ± 1.0	-1.30 ± 0.04	12 46 17.32	-12 45 24.6			
WN J1246–2426	112 ± 7	13.8 ± 0.6	-1.52 ± 0.05	12 46 44.15	-24 26 17.8			
WN J1252–2057	70 ± 6	11.3 ± 0.6	-1.32 ± 0.07	12 52 02.13	-20 57 12.0			
WN J1252–2015	99 ± 6	15.9 ± 1.0	-1.33 ± 0.06	12 52 50.30	-20 15 02.8			
WN J1254–1604	84 ± 6	12.8 ± 0.6	-1.36 ± 0.06	12 54 11.90	-16 04 44.2			
WN J1255–1256	184 ± 9	25.1 ± 0.9	-1.44 ± 0.04	12 55 33.48	-12 56 37.0			
WN J1255–1913	108 ± 6	10.8 ± 0.6	-1.67 ± 0.06	12 55 52.66	-19 13 00.6			
WN J1255–1222	2292 ± 94	321.6 ± 9.7	-1.42 ± 0.04	12 55 55.45	-12 22 54.7			
WN J1300–1146	317 ± 15	48.7 ± 1.5	-1.36 ± 0.04	13 00 46.89	-11 46 22.8			
WN J1305–1441	121 ± 7	19.8 ± 0.8	-1.31 ± 0.05	13 05 56.83	-14 41 57.9			
WN J1315–1511	67 ± 7	10.2 ± 0.6	-1.36 ± 0.09	13 15 43.95	-15 11 53.9			
WN J1326–2330	523 ± 22	80.2 ± 2.5	-1.36 ± 0.04	13 26 25.24	-23 30 23.4		DSS	De Breuck et al. (2000)
WN J1331–1947	488 ± 20	70.8 ± 2.2	-1.40 ± 0.04	13 31 47.18	-19 47 26.5			
WN J1345–2423	130 ± 9	19.1 ± 0.8	-1.39 ± 0.06	13 45 47.68	-24 23 45.6			
WN J1349–2049	203 ± 10	25.4 ± 0.9	-1.51 ± 0.04	13 49 02.91	-20 49 11.1		DSS	
WN J1351–1205	183 ± 9	19.6 ± 1.2	-1.62 ± 0.06	13 51 14.67	-12 05 53.6			
WN J1353–2152	347 ± 15	46.0 ± 1.8	-1.46 ± 0.04	13 53 55.92	-21 52 12.6			
WN J1416–1130	912 ± 37	142.0 ± 5.5	-1.35 ± 0.04	14 16 04.09	-11 30 30.2			
WN J1417–1522	89 ± 10	14.3 ± 0.7	-1.32 ± 0.09	14 17 29.38	-15 22 30.0			
WN J1429–1909	99 ± 7	14.8 ± 0.6	-1.37 ± 0.06	14 29 07.73	-19 09 50.5			
WN J1437–2040	118 ± 7	12.5 ± 1.0	-1.63 ± 0.07	14 37 06.06	-20 40 21.0			
WN J1446–2003	1164 ± 48	182.1 ± 5.5	-1.34 ± 0.04	14 46 48.51	-20 03 37.7	0.753	R	Giraud et al. (1996)
WN J1451–2301	128 ± 7	14.6 ± 1.2	-1.57 ± 0.07	14 51 15.54	-23 01 46.5			
WN J1454–1130	76 ± 6	10.2 ± 1.1	-1.46 ± 0.09	14 54 49.59	-11 30 03.0			
WN J1457–0954	120 ± 9	15.5 ± 1.4	-1.48 ± 0.09	14 57 53.71	-09 54 22.9			
WN J1513–1801	544 ± 23	83.0 ± 2.5	-1.36 ± 0.04	15 13 55.31	-18 01 07.9			De Breuck et al. (2000)
WN J1516–2110	230 ± 10	34.1 ± 1.1	-1.38 ± 0.04	15 16 42.32	-21 10 27.4			

Table 2. continued.

Name	S_{352} mJy	S_{1400} mJy	α_{352}^{1400}	α_{J2000} h m s	δ_{J2000} ° ' "	z	ID	Reference
WN J1518–1225	105 ± 6	10.5 ± 0.6	-1.67 ± 0.06	15 18 43.43	-12 25 35.6			
WN J1544–2104	243 ± 11	24.3 ± 1.2	-1.67 ± 0.05	15 44 04.77	-21 04 11.2			
WN J1551–1422	83 ± 5	10.7 ± 0.6	-1.48 ± 0.06	15 51 58.78	-14 22 48.3			
WN J1557–1349	90 ± 6	13.3 ± 0.6	-1.39 ± 0.06	15 57 41.72	-13 49 54.8			
WN J1558–1142	1316 ± 54	201.8 ± 6.1	-1.36 ± 0.04	15 58 00.63	-11 42 24.9			
WN J1603–1500	127 ± 7	17.4 ± 0.7	-1.44 ± 0.05	16 03 04.78	-15 00 53.8			
WN J1605–1952	241 ± 11	39.2 ± 1.3	-1.32 ± 0.04	16 05 53.98	-19 52 00.6			
WN J1631–1332	81 ± 5	12.6 ± 0.6	-1.35 ± 0.06	16 31 49.83	-13 32 28.7			
WN J1633–1517	93 ± 7	12.0 ± 0.6	-1.48 ± 0.06	16 33 15.09	-15 17 25.9			
WN J1634–1107	172 ± 9	17.4 ± 0.7	-1.66 ± 0.05	16 34 41.37	-11 07 12.9			
WN J1637–1931	327 ± 14	35.8 ± 1.2	-1.60 ± 0.04	16 37 44.85	-19 31 22.9			De Breuck et al. (2000)
WN J1653–1756	73 ± 5	10.1 ± 1.2	-1.44 ± 0.10	16 53 50.11	-17 56 04.6			
WN J1653–1155	1179 ± 48	191.5 ± 5.8	-1.32 ± 0.04	16 53 52.81	-11 55 59.1			De Breuck et al. (2000)
WN J1702–1414	156 ± 11	25.0 ± 0.9	-1.33 ± 0.06	17 02 51.45	-14 14 55.8			
WN J1932–1931	4841 ± 198	789.4 ± 23.7	-1.31 ± 0.04	19 32 07.22	-19 31 49.6			De Breuck et al. (2000)
WN J1939–1457	193 ± 11	28.8 ± 1.0	-1.38 ± 0.05	19 39 13.18	-14 57 27.3			
WN J1942–1514	385 ± 22	33.5 ± 1.1	-1.77 ± 0.05	19 42 07.50	-15 14 37.9			
WN J1954–1207	636 ± 27	98.9 ± 3.0	-1.35 ± 0.04	19 54 24.24	-12 07 49.3		K	De Breuck et al. (2000)
WN J1956–1308	96 ± 6	15.9 ± 0.7	-1.31 ± 0.06	19 56 06.62	-13 08 08.4			
WN J2002–1842	81 ± 5	11.4 ± 0.6	-1.42 ± 0.06	20 02 56.00	-18 42 47.8			
WN J2007–1840	230 ± 10	34.6 ± 1.1	-1.37 ± 0.04	20 07 00.16	-18 40 57.3			
WN J2007–1843	102 ± 5	16.7 ± 1.0	-1.31 ± 0.06	20 07 31.49	-18 43 43.9			
WN J2007–1316	922 ± 38	112.9 ± 3.4	-1.52 ± 0.04	20 07 53.23	-13 16 45.0		K	De Breuck et al. (2000)
WN J2014–2115	311 ± 14	47.3 ± 1.5	-1.36 ± 0.04	20 14 31.69	-21 15 02.6			De Breuck et al. (2000)
WN J2027–1909	86 ± 5	11.7 ± 0.6	-1.44 ± 0.06	20 27 01.50	-19 09 53.8			
WN J2045–1948	103 ± 6	15.5 ± 0.7	-1.37 ± 0.05	20 45 56.24	-19 48 19.6			
WN J2052–2306	356 ± 15	57.0 ± 2.1	-1.33 ± 0.04	20 52 50.08	-23 06 22.8			
WN J2054–2006	71 ± 6	10.4 ± 0.6	-1.39 ± 0.08	20 54 35.75	-20 06 09.4			
WN J2054–1939	74 ± 6	12.2 ± 0.6	-1.31 ± 0.07	20 54 52.30	-19 39 52.6			
WN J2103–1917	1358 ± 56	223.8 ± 6.7	-1.31 ± 0.04	21 03 42.91	-19 17 47.5			De Breuck et al. (2000)
WN J2104–2037	183 ± 8	27.6 ± 0.9	-1.37 ± 0.04	21 04 13.33	-20 37 27.7			
WN J2105–1057	136 ± 8	16.9 ± 0.7	-1.51 ± 0.05	21 05 45.16	-10 57 33.5			
WN J2106–1040	322 ± 14	50.0 ± 1.9	-1.35 ± 0.04	21 06 39.96	-10 40 43.5		DSS	
WN J2114–2127	203 ± 9	33.1 ± 1.1	-1.31 ± 0.04	21 14 26.43	-21 27 50.9			
WN J2116–1519	73 ± 11	11.0 ± 0.9	-1.37 ± 0.12	21 16 13.07	-15 19 38.6			
WN J2133–1656	234 ± 10	28.3 ± 1.3	-1.53 ± 0.05	21 33 24.84	-16 56 22.2			
WN J2137–1246	71 ± 6	11.2 ± 0.6	-1.34 ± 0.07	21 37 30.62	-12 46 43.7			
WN J2137–1350	258 ± 12	35.6 ± 1.2	-1.43 ± 0.04	21 37 47.42	-13 50 28.7			
WN J2139–1205	500 ± 21	82.0 ± 2.5	-1.31 ± 0.04	21 39 58.18	-12 05 28.6			
WN J2141–1425	114 ± 9	18.9 ± 0.7	-1.30 ± 0.06	21 41 18.88	-14 25 48.1			
WN J2144–1818	201 ± 9	30.9 ± 1.3	-1.36 ± 0.05	21 44 18.91	-18 18 36.6			
WN J2145–2240	192 ± 9	24.3 ± 1.2	-1.50 ± 0.05	21 45 15.94	-22 40 26.4			
WN J2204–1004	220 ± 12	28.1 ± 0.9	-1.49 ± 0.05	22 04 00.00	-10 04 21.3			
WN J2214–2353	188 ± 9	29.6 ± 1.3	-1.34 ± 0.05	22 14 14.03	-23 53 25.6			
WN J2216–1725	458 ± 20	13.4 ± 1.1	-2.56 ± 0.07	22 16 57.57	-17 25 21.4	0.1301	2MASS	Jarrett et al. (2000)
WN J2217–1913	284 ± 12	39.2 ± 1.3	-1.43 ± 0.04	22 17 28.22	-19 13 20.9			De Breuck et al. (2000)
WN J2226–1151	77 ± 5	12.5 ± 0.6	-1.31 ± 0.06	22 26 04.17	-11 51 02.3			
WN J2244–2520	84 ± 8	11.5 ± 0.6	-1.44 ± 0.08	22 44 28.75	-25 20 45.3			
WN J2302–2229	229 ± 10	37.9 ± 1.2	-1.30 ± 0.04	23 02 49.64	-22 29 44.4			
WN J2308–1546	104 ± 6	16.2 ± 0.7	-1.35 ± 0.05	23 08 16.46	-15 46 30.3			

Table 2. continued.

Name	S_{352} mJy	S_{1400} mJy	α_{352}^{1400}	α_{J2000} h m s	δ_{J2000} ° ' "	z	ID	Reference
WN J2320–1436	78 ± 6	10.1 ± 0.6	-1.48 ± 0.07	23 20 22.54	-14 36 19.6			
WN J2349–1542	129 ± 7	17.2 ± 1.2	-1.46 ± 0.06	23 49 40.09	-15 42 51.3			
WN J2350–1321	388 ± 16	55.1 ± 1.7	-1.41 ± 0.04	23 50 15.92	-13 21 15.9			De Breuck et al. (2000)
WN J2350–1238	104 ± 6	10.5 ± 0.6	-1.66 ± 0.06	23 50 49.70	-12 38 33.1		2MASS	Jarrett et al. (2000)

^a We provide either the name of the large sky survey, or the filter of literature observations.

who found that >3% of their USS sources are associated with nearby galaxy clusters.

5. Conclusions

The WISH survey is the deepest low-frequency survey covering roughly a quarter of the area between the WENSS and SUMSS surveys ($-30^\circ < \text{Dec} < +28^\circ$). Because it has been observed at very low elevations, it has relatively poor declination resolution compared with the NVSS. It is sufficiently deep to provide spectral index information for $\sim 42\,000$ sources in common with the NVSS. This provides a unique sample of faint USS sources which can be observed with southern hemisphere telescopes. We have obtained VLA and/or ATCA snapshot observations of a first batch of 69 sources from this sample to provide more accurate positions and morphological information needed for the identification of the host galaxies. A campaign of K -band identifications with CTIO, and optical spectroscopy with the VLT has already identified several new $z > 3$ radio galaxies from this sample (de Vries et al., in preparation).

Acknowledgements. We thank Wim de Vries of useful discussions. The Westerbork Synthesis Radio Telescope (WSRT) is operated by the Netherlands Foundation for Research in Astronomy (NFRA) with financial support of the Netherlands Organization for Scientific Research (NWO). This work was supported by a Marie Curie Fellowship of the European Community programme “Improving Human Research Potential and the Socio-Economic Knowledge Base” under contract number HPMF-CT-2000-00721. The work by W.v.B. at IGPP/LLNL was performed under the auspices of the U.S. Department of Energy, National Nuclear Security Administration by the University of California, Lawrence Livermore National Laboratory under contract No. W-7405-Eng-48. This work was carried out in the context of EARA, the European Association for Research in Astronomy.

References

Abell, G., Corwin, H., & Olowin, R. 1989, *ApJS*, 70, 1
 Becker, R., White, R., & Helfand, D. 1995, *ApJ*, 450, 559
 Bennett, A. 1962, *MmRAS*, 68, 163
 Best, P., Longair, M., & Röttgering, H. 1998, *MNRAS*, 295, 549
 Blundell, K., Rawlings, S., Eales, S., Taylor, G., & Bradley, A. 1998, *MNRAS*, 295, 265
 Bock, D., Large, M., & Sadler, E. 1999, *AJ*, 117, 1578
 Condon, J., Cotton, W. D., Greisen, E. W., et al. 1998, *AJ*, 115, 1693
 De Breuck, C., van Breugel, W., Röttgering, H., & Miley, G. 2000, *A&AS*, 143, 303

De Breuck, C., van Breugel, W., Röttgering, H., et al. 2001, *AJ*, 121, 1241
 De Breuck, C., van Breugel, W., Stanford, S. A., et al. 2002, *AJ*, 123, 637
 de Ruiter, H., Parma, P., Stirpe, G. M., et al. 1998, *A&A*, 339, 34
 de Vries, W., Morganti, R., Röttgering, H. J. A., et al. 2002, *AJ*, 123, 1784
 Douglas, J. N., Bash, F. N., Bozyan, F. A., Torrence, G. W., & Wolfe, C. 1996, *AJ*, 111, 1945
 Giraud, E., Melnick, J., Gopal-Krishna, & van Drom, E. 1996, *A&A*, 311, 446
 Haverkorn, M., Katgert, P., & de Bruyn, A. 2000, *A&A*, 356, L13
 Jarrett, T., Chester, T., Cutri, R., et al. 2000, *AJ*, 119, 2498
 Jarvis, M., Rawlings, S., Lacy, M., et al. 2001a, *MNRAS*, 326, 1563
 Jarvis, M., Rawlings, S., Willott, C., et al. 2001b, *MNRAS*, 326, 1585
 Kouwenhoven, M. 2000, *A&AS*, 145, 243
 Lacy, M., Rawlings, S., Hill, G., et al. 1999, *MNRAS*, 308, 1096
 Large, M. I., Mills, B. Y., Little, A. G., Crawford, D. F., & Sutton, J. M. 1981, *MNRAS*, 194, 693
 Le Fèvre, O., Deltorn, J., Crampton, D., & Dickinson, M. 1996, *ApJ*, 471, L11
 McCarthy, P., Kapahi, V., van Breugel, W., et al. 1996, *ApJS*, 107, 19
 McLure, R., & Dunlop, J. 2000, *MNRAS*, 317, 249
 Myers, S., Fassnacht, C. D., Djorgovski, S. G., et al. 1995, *ApJ*, 447, L5
 Pascarelle, S., Windhorst, R., Driver, S., Ostrander, E., & Keel, W. 1996, *ApJ*, 456, L21
 Pentericci, L., Kurk, J. D., Röttgering, H. J. A., et al. 2000, *A&A*, 361, L25
 Rawlings, S., Eales, S., & Lacy, M. 2001, *MNRAS*, 322, 523
 Rengelink, R., Tang, Y., de Bruyn, A. G., et al. 1997, *A&A*, 124, 259
 Rengelink, R. 1999a, *The Most Distant Radio Galaxies*, 399
 Rengelink, R. 1999b, Ph.D. Thesis, Universiteit Leiden, available at <ftp://ftp.strw.leidenuniv.nl/pub/wenSS/thesis>
 Röttgering, H., van Ojik, R., Miley, G., et al. 1997, *A&A*, 326, 505
 Schoenmakers, A., de Bruyn, A., Röttgering, & van der Laan, H. 2001, *A&A*, 374, 861
 Snellen, I., Schilizzi, R. T., de Bruyn, A. G., et al. 1998, *A&AS*, 131, 435
 Spinrad, H., Djorgovski, S., Marr, J., & Aguilar, L. 1985, *PASP*, 97, 932
 Stern, D., Dey, A., Spinrad, H., et al. 1999, *AJ*, 117, 1122
 van Breugel, W., De Breuck, C., Stanford, S. A., et al. 1999, *ApJ*, 518, L61
 Venemans, B., Kurk, J. D., Miley, G. K., et al. 2002, *ApJ*, 569, L11
 Willott, C., Rawlings, S., Blundell, K., et al. 2002, *MNRAS*, in press [astro-ph/0205507]