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Resolving the chemical substructure of Orion-KL *(Corrigendum)*

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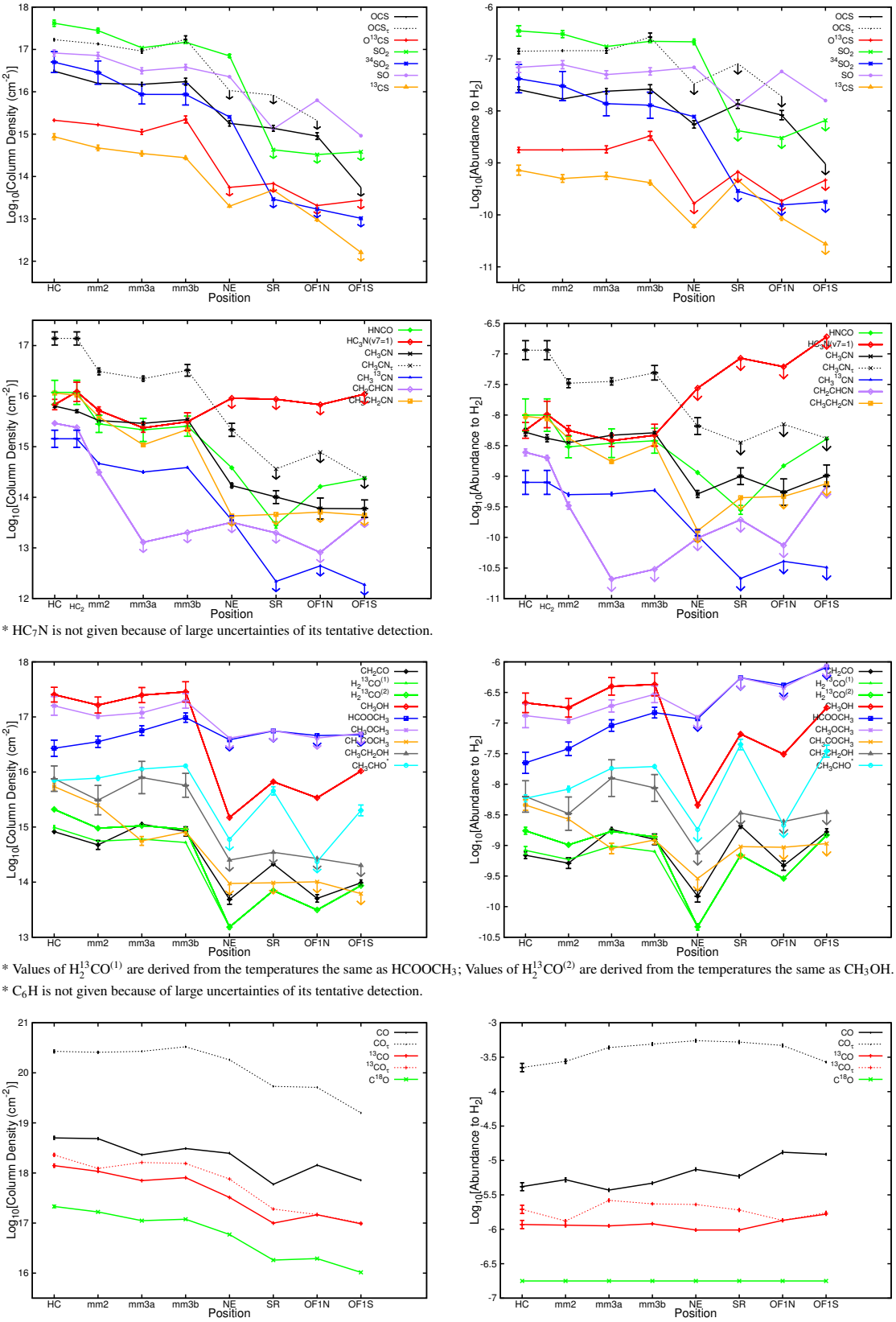
Key words. stars: formation – ISM: abundances – stars: massive – ISM: lines and bands – ISM: molecules – errata, addenda

In Fig. 10 of Feng et al. (2015), the dashed lines are incorrectly printed as the solid lines. These dashed lines show the opacity-corrected estimates of the molecular column densities and abundances towards eight positions in Orion-KL. See the correct figure next page.



* HC_7N is not given because of large uncertainties of its tentative detection.

* Values of $H_2{}^{13}CO^{(1)}$ are derived from the temperatures the same as $HCOOCH_3$; Values of $H_2{}^{13}CO^{(2)}$ are derived from the temperatures the same as CH_3OH.

* C_6H is not given because of large uncertainties of its tentative detection.

Fig. 10. Molecular column densities (*panels in the left column*) and abundances (*panels in the right column*) towards 8 positions in Orion-KL. The molecular abundances of each species with respect to H_2 are converted from $C^{18}O$, assuming a constant $N_{C^{18}O}/N_{H_2} \sim 1.79 \times 10^{-7}$ ratio. Column densities and abundances with the uncertainties are measured as mentioned in Sect. 4.4 and Tables A.5–A.7. Coloured solid lines show the differentiation for each isotopologue calculated from optically thin assumption (arrows highlight the upper limits), while dashed lines show the differentiation calculated with optical depth correction from Eq. (10). For abundant molecules (e.g., HNCO, SO, SO_2, CH_3OH) which we have not measured their line optical depths, their densities could be underestimated by a factor of 7–10.